

# Faraday rotation variations along radio jets: the magnetic field in galaxy and group halos

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**Abstract.** Our modelling of FRI radio jets as decelerating, relativistic flows allows us to derive their orientations accurately. We present images of Faraday rotation for two of these objects (3C 31 and NGC 315) and show that the fluctuations of rotation measure (RM) are larger in the fainter (receding) jets, as expected if the rotation occurs in the hot galaxy/group halos. The gas density is much lower in NGC 315 and the RM fluctuations are only just detectable.

**Key words:** galaxies: jets – radio continuum:galaxies – magnetic fields – polarization – MHD

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## 1. Introduction

In our models of FRI radio jets as relativistic flows (Laing, Canvin & Bridle 2006), the observed differences in brightness and linear polarization between the jets close to the nucleus are produced by relativistic aberration and Doppler beaming and we can determine the inclination,  $\theta$ . Given that we know the geometry and the external density profile (from X-ray observations), imaging of Faraday rotation measure (RM) can determine the distribution of magnetic-field irregularities in the surrounding hot plasma. In this paper, we summarize our RM imaging for two sources: 3C 31 (Laing & Bridle 2002) and NGC 315 (Canvin et al. 2005).

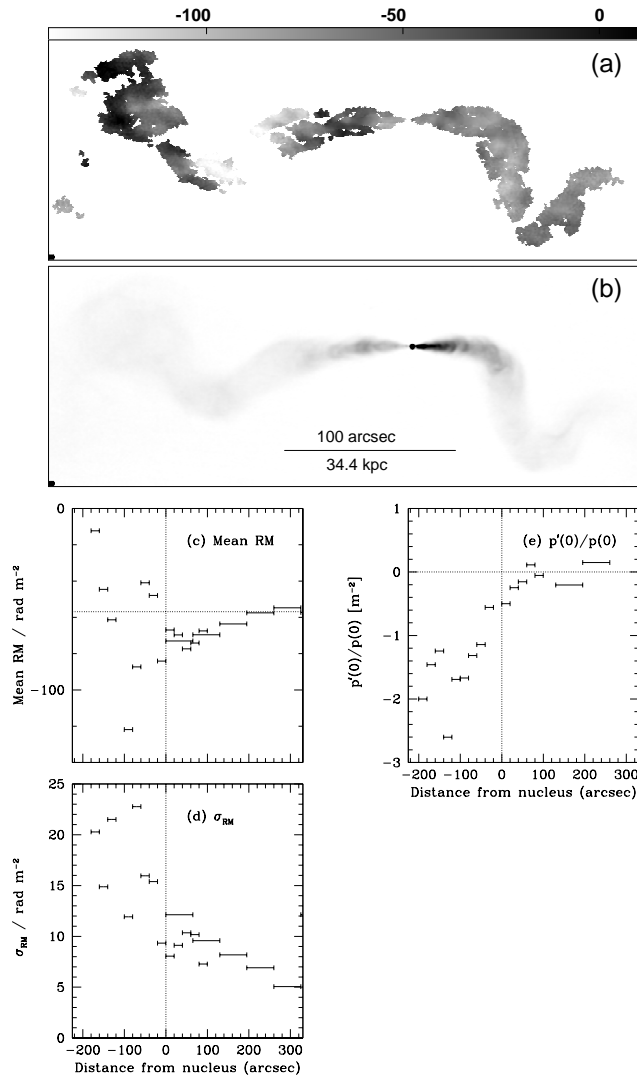
## 2. 3C 31

3C 31 is an FRI radio galaxy at a redshift of 0.0169 (0.344 kpc/arcsec for  $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ ). Our models of the inner jets give an inclination of  $\theta = 52^\circ$ . Our RM image, derived from 6-frequency observations between 1.365 and 8.4 GHz, is shown in Fig. 1(a), with an  $I$  image at the same resolution for comparison in Fig. 1(b). The E-vector position angle is accurately proportional to  $\lambda^2$  everywhere, indicating foreground rotation. There is structure in the RM image on a

range of spatial scales from 5 to  $>50$  arcsec and the RM fluctuations are higher by a factor of 2–3 on the counter-jet side (Fig. 1c and d). There is a small amount of depolarization on the counter-jet side. We estimate this by making a first-order linear approximation to the variation of degree of polarization,  $p$ , with  $\lambda^2$ :  $p(\lambda^2)/p(0) \approx 1 + [p'(0)/p(0)]\lambda^2$  (Fig. 1e). We model the thermal X-ray emission as the sum of two components: one associated with the galaxy (core radius  $r_c = 3.6$  arcsec), the second with the surrounding group, with  $r_c = 154$  arcsec (Hardcastle et al. 2002). The Faraday rotation variations are therefore plausibly associated with the group component. Applying the model described by Laing et al. (2006), we can estimate the central magnetic field strength:  $B_0/nT \approx 0.9(l/\text{kpc})^{-1/2}$ , where  $l$  is the correlation length of the field.

## 3. NGC 315

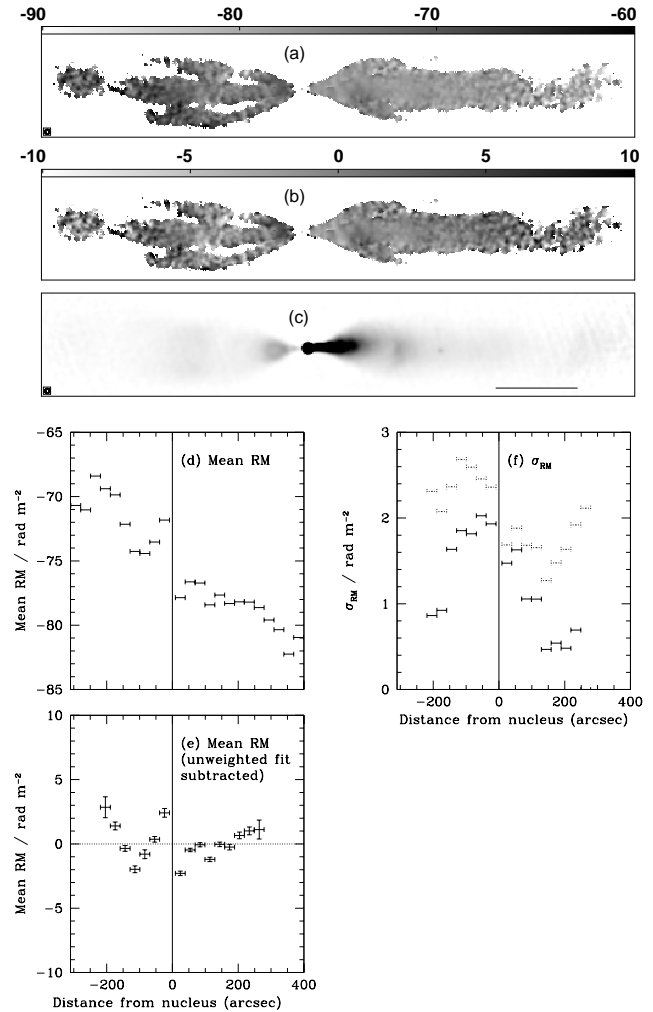
The giant FRI radio galaxy NGC 315 is at a redshift of 0.01648 (0.335 kpc/arcsec). We infer an angle to the line of sight of  $38^\circ$  (Canvin et al. 2005). Our RM images, derived from 5-frequency observations in the range 1.365–5 GHz, are shown in Fig. 2, together with profiles along the jet axis. The mean RM and a linear gradient along the jet are almost certainly Galactic. After removing these components, we detect RM fluctuations on scales of 10–100 arcsec, but the typical amplitudes are only  $\approx 2 \text{ rad m}^{-2}$  (Fig. 2e and f). The thermal X-ray emission from NGC 315 has a very



**Fig. 1.** RM and depolarization for 3C 31. All diagrams have the approaching (brighter) jet on the right. (a) RM image at a resolution of 1.5 arcsec, made using the Pacerman algorithm of Dolag et al. (2005). (b)  $I$  image of the same area. (c) Profile of mean RM, determined in boxes along the jet axis. (d) The rms RM, calculated over the same areas as in panel (c). (e) Profile of  $p'(0)/p(0)$ , as described in the text.

small core radius,  $r_c = 1.55$  arcsec (Worrall, Birkinshaw & Hardcastle 2003). No X-ray emission has yet been detected from the galaxy group associated with NGC 315 (Miller et al. 2002), but the small amplitude of the RM fluctuations and the fact that they are larger for the receding jet are both consistent with an origin in a tenuous group gas component with a scale size  $\approx 200$  arcsec. As we do not know the density of this component, we can only constrain the product  $(n_0/m^{-3})^2 (B_0/nT)^2 (l/kpc) \approx 700$ , where  $n_0$  is the central density of the group component and  $B_0$  and  $l$  are again the central field strength and correlation length, respectively (Laing et al. 2006).

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**Fig. 2.** RM data for NGC 315. All diagrams have the approaching jet on the right. (a) RM image at a resolution of 5.5 arcsec. (b) As (a), but with a linear variation subtracted. (c)  $I$  image covering the same area (a scale of 100 arcsec or 34.4 kpc is indicated by the horizontal line). (d) Profile of RM along the jet axis. (e) As (d), but with a linear variation subtracted and with low s/n data omitted. (f) The rms RM for the same data points as in panel (e). The dotted lines show the raw values; the full lines are the values after a first-order correction for fitting error. Full details are given by Laing et al. (2006).

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