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STRUCTURES, SPECTRAL INDEXES, AND OPTICAL IDENTIFICATIONS OF RADIO SOURCES SELECTED FROM THE B3 CATALOGUE

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ABSTRACT

This paper presents the first results of a large program of optical identifications of radio sources from the *Third Bologna (B3) Catalogue.* A homogeneous sample of 1103 radio sources, selected from the B3 Catalogue at 408 MHz and covering the flux-density range from 2.0 to 0.1 Jy, has been mapped at 1465 MHz using the VLA at 15 arcsec resolution. Both positonal and structural information have been derived and are presented here. Optical identifications have been attempted for all of the sources using the Palomar Sky Survey prints, resulting in 354 proposed identifications. The run of identification percentage, largest angular size (LAS), and spectral index with flux density is analyzed. Evidence is given that sizable intrinsic radio–optical displacements are possible for steep-spectrum unresolved sources. The fraction of "empty field" identifications (fainter than $m_r = 20$) increases from 15% for $S_{408} \geqslant 10$ Jy to 86% in the 20–60 mJy interval. The 408–1465 MHz spectra become steeper with decreasing flux density to about 50 mJy, and become flatter at lower flux densities. Correspondingly, a small discontinuity is found in the behavior of the apparent diameter versus flux-density relation, at a similar flux density. A negative correlation is found between the metric size and redshift for radio galaxies, confirming an earlier result of Grueff *et al.* (1977).

I. INTRODUCTION

It is generally recognized that radio-source counts are difficult to reconcile with the counts expected in a Euclidean space uniformly filled with radiating sources. When the predictions are based on relativistic metrics, in the generally accepted framework of an expanding universe, the difficulties are much aggravated.

Since no reasonable universe model (i.e., metric) has been devised explaining the empirically measured radio-source counts, a widely accepted explanation of the discrepancy has been that the space density of radio sources, or their absolute luminosity, or both, must vary with cosmological epoch.

Much effort has been devoted since the pioneering work of Oort (1961), Longair (1966), and Rowan-Robinson (1967) to define the properties of the cosmological evolution of the radio sources.

The source counts are now known with high accuracy at many frequencies (see, for example, Kellermann and Wall (1986) for a review on the subject), and recent advances have extended them over a range of as many as five orders of magnitude in flux density, reaching a surface density of 10^6/sr (Windhorst *et al.* 1985; Weistrop *et al.*, 1987, and references therein).

However, modeling the counts with a cosmological evolution requires a detailed knowledge of the source absolute-luminosity distribution at each flux density, which can only be attained through radio-source optical identifications. This has proved to be a formidable task; after two decades of effort, it can be considered complete only for a sample limited to a few dozen of the strongest radio sources on the sky (Gunn *et al.* 1981; Peacock *et al.* 1981; Spinrad *et al.* 1985).

A physical understanding of a model cosmological evolution needs a detailed study of physical properties of individual sources at large redshifts, a process now only in the beginning stages (see, for example, Djorgovski 1988, and references therein).

At present, these studies are mostly restricted to the 3CR sample, which includes sources spanning a very wide range in absolute power and redshift, albeit with limited statistics. However, due to the very steep source counts, most sources in the 3CR Catalog fall within a rather restricted flux-density range (close to the Catalog limit) so that, in practice, sources with a given absolute power are only observed in a narrow redshift range. To overcome this limitation, and to extend our knowledge of the power–redshift plane (P–z), it is necessary to extend the optical identification process down to sources with much fainter flux densities.

The basic aspects of the problem are the same at all frequencies; much progress, however, has been made at low frequencies, not only for obvious, historical reasons, but also because the evolution as shown by the source counts is especially evident at low frequency.

At 408 MHz, systematic optical identifications have been made in the past on the B2 Catalogue (Grueff and Vigotti 1975, and references therein; Allington-Smith *et al.* 1982) and on the various 5C catalogs (see Benn *et al.* (1982) for comprehensive references on 5C source counts). These works were relevant to flux densities in the range 5–1 Jy and 1–0.01 Jy, respectively. This latter interval (reaching the faintest flux density at which the 408 MHz counts are known) has been the subject of a recent, in-depth study (Benn *et al.* 1988, and references therein).

The number of sources involved in these studies, however, is not very large, especially in the interval from 1 to 0.1 Jy, where only a few hundred sources (giving a few dozen identifications) have been studied. The first section of the recently published B3 Catalogue (Ficarra *et al.* 1985), covering a sky area of 0.78 sr and listing over 13 000 sources down to a 408 MHz flux density of 0.1 Jy, is ideally suited to provide a complete source sample in this flux-density range. A large observational program has been started at the NRAO's Very Large Array (VLA) radio telescope, aimed at obtaining accurate medium- and high-resolution maps of a suitably large sample of B3 sources. This paper describes a set of medium-resolution VLA observations of a complete subsample of 1103 B3 sources, and the results of an optical identification program performed on the POSS prints.

The sample-selection criterion and the possible biases affecting it are analyzed in Sec. II. Section III deals with the VLA observations, radio-data reduction, and discussion of some sources of error. The identification procedure is described in Sec. IV; the data themselves, and further sources of error that affect the identification procedure, are described in Sec. V. In Sec. VI, the data are analyzed and compared to results from previous studies at the same frequency. A full discussion of these results in the frame of cosmological evolution of the radio-source properties, and a comparison with similar data at other observing frequencies, are deferred to future papers.

II. THE DEFINITION OF THE SAMPLE

Five complete subsamples of the B3 Catalogue, separated by equal increments in logarithmic flux density, were defined down to the Catalogue limit of 0.1 Jy. The selection criterion adopted consists in fixing the sample's declination limits in such a way as to yield a roughly equal number of sources in each flux-density interval.

In addition, we restricted the right ascension range to exclude sky areas at low galactic latitude, because obscuration makes these unsuitable for optical identifications, and confusion makes them unsuitable for VLA snapshot-mode observations. Thus the samples were selected only from the R.A. intervals 23^h00^m–03^h00^m and 07^h00^m–15^h00^m. Table I shows the sample definitions, and the number of sources in each sample. All the sky-area limits are at 1950.0 epoch, except the declination limits for sample 4, which correspond to the whole B3 Catalogue, and are referred to 1978.0 epoch.

Also, small regions of incompleteness are present in each sample; see Ficarra *et al.* (1985) for details on these. The percentage of sky area affected is negligible, since the main incompleteness areas in B3 are entirely within the limits of the excluded R.A. intervals.

TABLE I. Definition of the B3 samples.

Sample #	Declination range	Flux-density range	Number of sources
0	+ 39°38′ to + 40°00′	$0.1 \leqslant S < 0.2$	196
1	+ 38°50′ to + 40°00′	$0.2 \leqslant S < 0.4$	268
2	+ 38°00′ to + 40°48′	$0.4 \leqslant S < 0.8$	257
3	+ 38°00′ to + 44°30′	$0.8 \leqslant S < 1.6$	217
4	+ 37°15′ to + 47°37′	$1.6 \leqslant S$	165

We note that the number of sources in sample 0 and sample 4 is somewhat smaller; for sample 4, the number is limited to the extension of the whole B3 Catalogue. For sample 0, the low number is due to the fact that the samples were originally defined by the provisional B3 list, and to ensure completeness, even very uncertain sources were included; most of these were later deleted from the final Catalogue. This final trimming significantly affected only the faintest flux-density interval. In fact, observing these uncertain sources helped to define a completeness limit for the B3 Catalogue itself, as discussed in the original paper.

The samples now include all the B3 Catalogue sources meeting the above selection criteria, with no further selection. Specifically, the Catalogue includes a number of sources marked with a four-digit number (see the original paper) indicating confusion, higher noise, and, in general, possible problems; these sources were all included. Consequently, the five samples should be a fair statistical sample of the extragalactic source population at large as present in the B3 Catalogue.

It is perhaps worthwhile to point out that, owing to the Bologna telescope interferometric response (a HPFW of $3' \times 5'$ with a rather peculiar *UV* plane coverage), there are substantial selection effects against diffuse sources with typical size exceeding about 2 arcmin, as discussed in the B3 paper. Such diffuse sources appear, in retrospect, to be very rare, and statistics such as the median angular size are most likely unbiased by this selection effect. However, they must be carefully evaluated to obtain any meaningful estimate of statistics related to the "tail" of the size distribution toward large angular sizes. This statement, of course, applies practically to any deep low-frequency catalog of radio sources, and most notably to the 5C, Molonglo (Robertson 1977), and Texas (Douglas *et al.* 1980) catalogs.

III. THE VLA OBSERVATIONS

Ideally, one would need to observe each source at two frequencies at least, to obtain full spectral and structural information. Also, mapping only at the highest VLA angular resolution would not be adequate, because these maps would be insensitive to low-brightness features. Furthermore, the VLA is capable of producing maps with a very high dynamic range and with full polarization information, and of course it would be extremely interesting to exploit both these capabilities, to gain information about the physics driving the radio sources.

To extend such careful mapping to a sample of over one thousand sources would require a very large amount of observing time; the program was thus limited to short integration time, "snapshot" maps at a single frequency, and two array configurations, A and C. The frequency selected was 1465 MHz, with a bandwidth of 50 MHz; this was the lowest frequency available at the VLA when the observations were planned (1980) but still different enough from the B3 frequency that useful spectral-index information could be obtained.

Integration time was generally 3 min per map, regardless of the 408 MHz flux density. The expected map noise was 0.3 mJy/beam, while the faintest B3 sources, for a spectral index of $\alpha = -1.2$ and an angular area size of four beams (at C configuration), would have a brightness of about 6 mJy/beam. These would be unusual parameters for a radio source, as we show later. In some cases, however, sources proved to be fainter than this limit, and others were confused by sources outside the initial map area. To deal with these cases, about one hundred sources were reobserved at lower resolution with the VLA in its D configuration.

Most sources were observed close to meridian transit, and thus also close to zenith. The resulting beam shapes were therefore very similar for all maps, with minimal ellipticity and a HPFW of 1.4, 14, and 42 arcsec, respectively, for A, C, and D configurations. The flux-density scales of all observing sessions were normalized to that of Baars *et al.* (1977) by observing 3C 286, whose flux density on this scale is 14.51 Jy at 1465 MHz.

All observing was done between November 1981 and March 1983; the present paper deals with the lower-resolution C and D configuration total-intensity maps only, which are fully adequate for the purpose of optical identifications on the POSS prints. A detailed radio-structure classification must await the reduction of the A configuration and polarization data.

Data reduction was first performed in a standardized way, using a standard CLEAN algorithm (Clark 1980) and a map size of 256×256 pixels, giving a field of about 20×20 arcmin. For most sources, this procedure gave a good-quality map, with a rms noise usually ~ 1 mJy per beam. Very faint and/or diffuse sources, or heavily confused fields, were remapped and deconvolved using the ungridded-subtraction CLEAN algorithm (Schwab 1984) that is implemented in the AIPS task MX. This algorithm significantly reduces the rms fluctuations due to sidelobes of confusing sources in VLA snapshot maps.

In addition, about one hundred fields were reobserved in D configuration; this gave useful results not only in the cases of very diffuse sources, but also for very crowded and confused fields. Source parameters (position, peak, and integrated flux density) were measured by hand on the maps, which generally had a scale of about 2 arcsec/mm. Integrated flux densities for extended sources and components were obtained by multiplying the peak flux density by the half-power area, normalized to that of a point source.

Errors in positions, although larger than those achievable with a digitized measurement, were smaller than the errors in optical positions, as we will show later. Errors in flux density introduced by the manual measurement were also smaller than those intrinsic in the observations. A detailed error analysis will be given in Sec. V.

IV. THE IDENTIFICATION PROCEDURE

The nominal radio positions were at first located on the PSS prints with an accuracy of about 10 arcsec, and a 10× enlargement of both E and O prints was made. The optical position of the three optical objects closest to the radio position, and of three reference stars (14–17 mag) a few minutes of arc away, were measured with the Bologna *x–y* plate measuring machine, with reference to FK4 stars and with an rms accuracy better than 1 arcsec. Transparent overlays were produced at the scale of the print enlargements (about 6 arcsec/mm) on which the measured optical objects were marked, together with the original B3 position and the position(s) of the radio component(s) measured at VLA. By means of these overlays, all optical candidates within a given searching area were selected, and their optical position remeasured on the PSS prints.

The final optical positions were measured using a linear interpolation from ten FK4 reference stars conveniently located around each object. Figures 1(a) and 1(b) show the histogram of the differences in positions for objects close to PSS print boundaries, which can be measured on two or more plates.

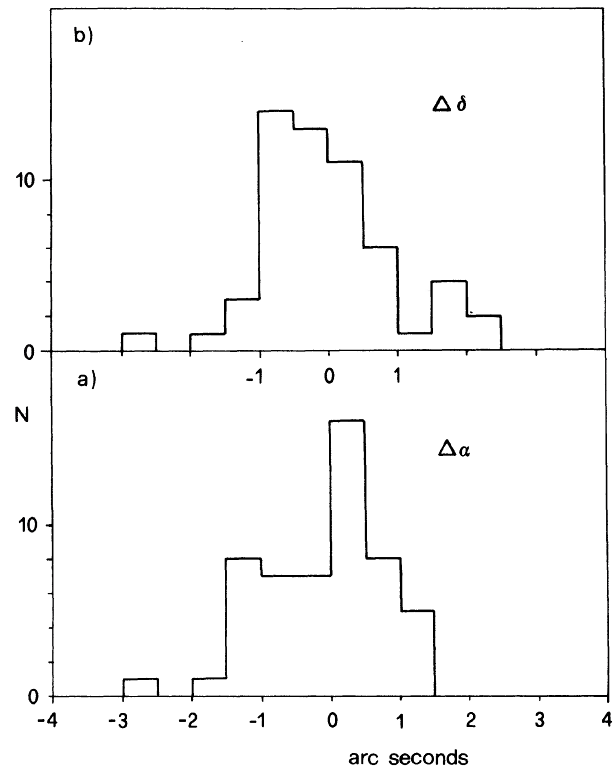

FIG. 1. Histograms of the differences in R.A. ($\Delta\alpha$) and in Dec. ($\Delta\delta$) for the optical objects (identifications and secondary reference stars) that have been measured on more than one plate.

The rms errors on a single measurement, inferred from Fig. 1, are 0.61 arcsec in R.A. and 0.68 arcsec in Dec. These are probably overestimated, since objects close to plate borders normally are not measurable in ideal conditions (i.e., choosing the ten FK4 stars to be evenly distributed around the object). We note that a *local, linear* fit for each object was found preferable to a higher-order, full-print solution, which gave inferior results even when many reference stars were measured. This is certainly due to relatively large print distortion; the problem is also present on plates, and in this case it must be caused by the optics of the 48″ Schmidt telescope (see discussion in Benn *et al.* 1988). The search for optical identification has been done differently according to the radio-source morphological classification and largest angular size (LAS).

a) LAS < 60 arcsec (916 Sources)

For single-peak sources a single circular area of radius 10 arcsec was inspected; for multiple-peak sources (double and triple sources) two 10 arcsec radius areas were inspected, centered at the unweighted average position (midpoint) and at the brightness-weighted average position (radio centroid); in addition, an area of 4 arcsec radius was inspected, centered at each radio component.

b) LAS≥60 arcsec (126 Sources)

Only 15 such sources are single peaked, the remainder being double, triple, or more complex. The basic searching method was the same as before, but the radius of the searching area for the centroid and midpoint positions was taken to be 20% of the LAS. Sometimes, an even larger radio–optical displacement was accepted, when a distorted radio image and a very bright optical candidate suggested a very probable identification (see below).

If more than one optical counterpart was found for a radio source, preference was given to midpoint or centroid positions, and to bright objects; however, all objects found are reported in the Notes to Table II (see Sec. V). At this stage, emphasis was placed on completeness of possible candidates rather than on reliability of identifications.

Magnitudes of starlike objects were obtained by reference to star images in Selected Area 57 (see Grueff *et al.* 1984); they were measured by eyeball estimate on the enlargements, and are therefore not very accurate, although previous experience indicated that the probable error is about 0.5 mag.

For galaxies, the image size has been measured, and from it the redshift has been inferred, following an established procedure for radio galaxies (Sandage 1972; Grueff and Vigotti 1977). The method has been recalibrated using spectroscopic redshifts available for 33 radio galaxies in our sample. Some of these are available in the literature, and many were measured by Kron (1986) as an optical followup of this program. Figure 2 shows the correlation between such spectroscopic redshifts and our estimate.

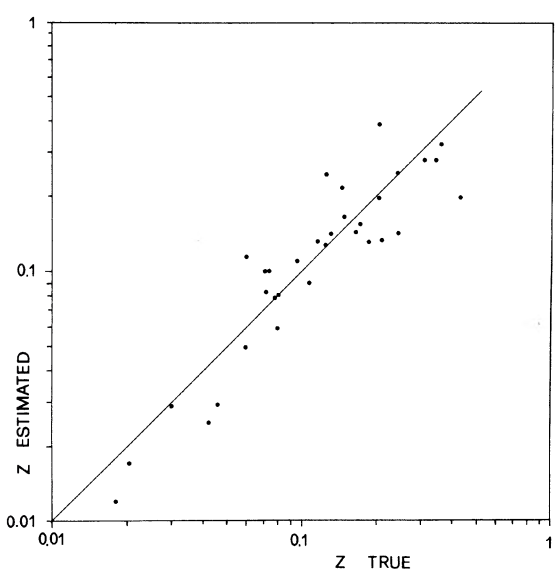

FIG. 2. Correlation between spectroscopic redshifts ($z_{\rm true}$) and redshifts estimated by measuring image size on the Palomar Sky Survey prints, for all galaxies in the B3 sample with spectroscopic redshift. The solid line is not a fit, but it is $z_{\rm true} = z_{\rm estimated}$.

V. THE RADIO AND OPTICAL DATA

Table II lists all the radio and optical data, as follows:

Column 1. The IAU name. This was not given in the B3 Catalogue, and was obtained from the B3 position (hours, minutes, degrees, and tenths of degrees) adding a letter (A, B, etc.) in order of R.A., in case of ambiguity. Note that this designation does not always match the new VLA position, although this happens rarely (see, for example, 0158 + 394). In the case of physically connected multiple sources, a single name is given if they correspond to a single source in the B3 Catalogue; however, some sources listed separately in B3 happen to be parts of the same source (see, for example, the sources 0050 + 401, 0050 + 403, and 0050 + 402B, which are physically connected) and in this case they retain their individual B3 names. If the B3 name is underlined, this is a reference to a note.

Column 2. The 4C name; if the source is also in the 3CR Catalogue, the 3CR name is given instead. 3CR names are recognizable by the format (integer number, or at most one decimal digit), while 4C names are in the DD.NN format. Identification with a 4C source is sometimes ambiguous, as discussed in the B3 Catalogue.

Column 3. The digit indicates the source parent subsample, as per Table I. This number indicates inclusion in the sample as originally selected; for example, the physical double 0218 + 399(A + B) is listed as

0218 + 399A 1,

0218 + 399B 2,

since the respective B3 flux densities are 0.36 and 0.49 Jy, and the components were originally selected as parts of two different samples. The VLA map showed them to be part of the same source; consequently the total 408 MHz flux density is now listed as 0.85 Jy (column 13) and in any subsequent analysis the source will be considered part of sample 3, $0.8 \leqslant S < 1.6$ Jy. Another example is given by the following double source 0218 + 402(A + B); only one component (A) was originally included in sample 2 (although both components are separately listed in B3) since component B did not meet the selection criteria as given in Sec. II.

Columns 4–9. The source (or component) position, at epoch 1950.0.

Columns 10 and 11. The source (or component) flux density at 1.46 GHz; units are milliJanskys. A " + " preceding the flux density indicates that the source (component) was resolved, and an integrated flux density is given.

Column 12. A simple morphological classification of the radio structure, as follows: U = unresolved, the synthesized beam is not appreciably broadened; R = resolved, only one emission peak is present, but the antenna beam is enlarged; D = double source, two emission peaks are present and they appear to be physically related; T = triple source, three peaks present and possibly related.

This coarse classification was adequate for the purpose of optical identifications on the PSS; a much more detailed structural classification will be possible when the A configuration observations are analyzed.

The source largest angular size (LAS), measured in arcseconds, follows the classification; U implies an (average) upper limit of 5 arcsec; R sources were given a size by measuring the largest HPFW, and deconvolving it quadratically with the antenna beam; for D and T sources the separation between the (outermost) components is given.

Column 13. The total 408 MHz flux density; generally it is taken from B3, but for extended, strong sources it has been remeasured integrating the B3 map on the structure revealed by the VLA data. Units are Janskys.

Column 14. The spectral index, computed between 1465 and 408 MHz; for multiple sources, the 1465 MHz flux densities have been added.

Column 15. The proposed optical identification: G = galaxy, B = starlike, blue object, N = starlike object of neutral color. A question mark indicates an uncertain optical classification, it does *not* necessarily indicate an uncertain identification. An asterisk preceding the optical classification indicates that the identification was accepted even if it does not meet a uniform criterion of positional agreement (see below).

Column 16. For N and B objects, the red magnitude is given; for galaxies, the redshift is given, estimated as explained in Sec. IV. However, if a spectroscopic redshift is available in the literature, it is given in this column, preceded by an asterisk. In this case, the reference given (see below) is to the redshift measurement.

Columns 17 and 18. The seconds of R.A. and arcseconds of Dec. of the *optical* object; hours and minutes in R.A., and degrees and arcminutes in Dec. are generally given by the radio position. This allows the recovery of full optical positional information.

Columns 19 and 20. Radio–optical coordinates differences, in arcseconds.

Column 21. A single digit indicates, for multiple sources, the radio position adopted (see Sec. IV) as follows: 1 = first component; 2 = second component; 3 = third component; 4 = unweighted average position (midpoint); 5 = weighted average position (radio centroid).

Column 22. A reference code number, underlined for better readability (see Table II(b)) if the source has been identified previously. For sources with a measured redshift, only the reference for the redshift is given; for 3CR sources, the reader is referred also to Spinrad *et al.* (1985).

Table II lists all the B3 sources meeting the selection criteria discussed in Sec. II; however, some were not observed; for various random reasons, some were not detected. For these, the columns from 4 to 12 are left blank, and are discussed in a note to the table. For double and triple sources, the columns 12–21 are filled only for the first component; the relevant information obviously refers to the source as a whole.

In Fig. 3, all the source maps are shown that were considered not to be adequately described by the data in Table II; the position of the proposed identification (if any) is marked with a cross.

Errors in optical positions have been already discussed in Sec. IV; radio-position uncertainties can now be conveniently estimated by reference to optical positions of safe, unambiguous identifications. For this purpose, we select the radio sources in Table II meeting all the following three criteria: (i) unresolved (U); (ii) flat radio spectra, namely $\alpha \geqslant -0.5$; (iii) identified with an object brighter than 19 mag. The last requirement minimizes the number of spurious identifications, while the first two help to eliminate possible cases of intrinsic radio–optical displacements.

There are 24 such sources in Table II, while the expected random-coincidence number is about 0.2. The results of the radio–optical comparison for these objects are given in Table III.

One can see that the R–O rms displacement is 0″.68 in R.A. and 0″.82 in Dec., with negligible systematic shift in R.A. and a small, probably not statistically significant shift in Dec.

Considering the errors given in Sec. IV for the optical, these R–O differences indicate a radio-position rms error of about 0.30–0.45 arcsec, close to what is expected from VLA data at this frequency and configuration, and still negligible compared to the optical errors.

Table III also shows some interesting results on radio–optical positional differences for *steep*-spectrum unresolved sources (same selection as above, but now $\alpha < -0.50$); their rms dispersion is very much increased (second row).

However, the expected number of spurious identifications in the sample is now of the order of 3. The last line in Table III shows the results obtained upon deletion of the four objects with the largest radio–optical displacements; one can see that the R–O rms displacement is still significantly larger for steep spectra. Thus, it appears that sizable *intrinsic* radio–optical displacements are possible for steep-spectrum sources, even if they look *unresolved* in a 14 arcsec beam map.

A discussion of errors in 408 MHz flux densities is given in Ficarra *et al.* (1985); the errors in the 1.46 GHz flux densities can be estimated by comparison with published data for 3CR sources.

There are 22 3CR sources in Table II; Table IV shows the results of a comparison with flux densities taken from Kellermann *et al.* (1969, hereafter referred to as KPW). The sample has been split according to 1.5 GHz flux density (median 1.8 Jy) and according to angular size (median 17 arcsec). The ratio R = KPW/VLA is given, using uncorrected KPW data.

After a 4% correction due to a small difference in observing frequency, we are left with a 13% difference, possibly due to resolution losses in the VLA data, confusion in the KPW data, or both, although the ratio appears to be independent of flux density, and only vaguely correlated with source angular size.

For example, the largest angular size is that of 3C 46, which gives $R = 0.97$; the average for the four unresolved 3C sources (3C 186, 216, 299, and 470) is $R = 1.13$. Moreover, the rms noise and confusion error quoted by KPW is 0.1 Jy, of the order of less than 7% in our 3CR sample. Thus, most of the observed *dispersion* should be due to the present measurements. An attempt to check the reliability of the manual measuring procedure has been made, by remeasuring with a proper fitting algorithm a number of representative maps. The measuring error has been found to be negligible for unresolved or barely resolved (<15 arcsec) source components. Consequently, the manual measuring procedure cannot be the source of the large observed dispersion. None of the 3CR sources in our sample is known to be variable.

A rather large sky area has been searched for optical identification, as described in Sec. IV; however, only identifications thought to be very reliable are retained in Table II. Specifically, for LAS < 60 arcsec the radius of the centroid and midpoint areas has been restricted to 6 arcsec; for $60 \leqslant \mathrm{LAS} < 120$, objects fainter than $m = 18$ ($z \approx 0.25$ for galaxies) have been searched in areas with radius = LAS/10 (centroid and midpoint positions), while objects brighter than these limits have been accepted up to a radius of LAS/5. Identifications with a component were always restricted to a 4 arcsec radius; in all cases, additional objects out to larger radio– optical displacements are listed in the notes to Table II, although very few of them are likely to be associated with the radio sources.

A few identifications, for which the above conditions on positional agreement are not met, but which are considered to be very probable on other grounds, are also listed in Table

TABLE II. The radio and optical data.

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0000+394		0	00	00	7.8	39	31	26		37	D190	.19	-.74	N	17.5	8.25	6.3	.0	-13.8	4	
			00	00	8.7	39	28	19		37	D										
0000+399		0	00	00	57.7	39	56	44	+	46	R 13	.15	-.93								
0001+395		0	00	01	41.4	39	33	10	+	30	R 12	.13	-1.15								
0001+398		1	00	01	43.7	39	49	17	+	30	D 36	.22	-.83	G?	0.198	45.28	30.5	-.1	-.8	5	
			00	01	46.3	39	49	38	+	46	D										
0003+380		2	00	03	22.3	38	03	33		347	U	.49	-.27	N	16.9	22.35	33.0	-.6	.0		17
0003+387	38.01	3	00	03	45.8	38	43	45		406	U	1.36	-.95								
0003+397		0										.11									
0004+380A	38.02	2	00	04	1.4	38	02	25		264	U	.79	-.86								
0005+383B		2	00	05	47.4	38	20	30	+	270	R 89	.53	-.53	G	0.080	47.47	28.6	-.8	1.4		
0006+397		3	00	06	28.5	39	45	05		495	U	1.15	-.66	B	18.7	28.54	5.2	-.5	-.2		50
0008+392		1	00	08	2.4	39	17	45		52	D 31	.20	-.77								
			00	08	0.4	39	17	24		23	D										
0010+392		2	00	10	7.1	39	16	15		137	U	.44	-.91								
0010+395		0	00	10	34.8	39	31	27		8	D 25	.14	-1.52								
			00	10	37.0	39	31	25	+	12	D										
0010+402		3	00	10	14.3	40	15	52	+	150	D 57	1.00	-.94	G	0.328	16.48	7.6	-.3	-.6	4	
			00	10	18.6	40	16	22	+	150	D										
0010+405	40.01	4	00	10	54.2	40	34	56	+	1087	R 15	3.26	-.86	G	0.165	54.29	56.2	-1.0	-.2		
0013+387		2	00	13	22.8	38	43	47		213	U	.74	-.98	B	18.6	22.79	47.0	.1	.0		
0013+393		1	00	13	11.2	39	20	27		101	U	.27	-.77								
0014+395		1	00	14	16.4	39	31	23		53	U	.21	-1.08								
0015+399		0	00	15	3.9	39	59	13	+	48	D 73	.14	-.65								
			00	15	3.5	40	00	26	+	13	D										
0017+395		0	00	17	10.0	39	35	15	+	46	R 10	.15	-.93								
0017+432		3	00	17	26.4	43	13	29	+	162	D 65	.80	-.92								
			00	17	32.3	43	13	33	+	84	D										
0018+393		1	00	18	41.4	39	21	54	+	32	T 61	.20	-.96								
			00	18	44.1	39	22	00		5	T										
			00	18	46.7	39	21	51	+	22	T										
0019+391		1	00	19	21.0	39	09	08		81	U	.29	-1.00	B?	21.0	20.64	9.3	4.2	-1.3		

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0019+431	43.01	4	00	19	7.3	43	11	51	+	202	D 28	2.22	-1.34	B	18.5	8.17	47.1	4.1	-.1	4	8
			00	19	9.8	43	11	43	+	202	D										
0020+437	43.02	3	00	20	49.7	43	46	13	+	233	D 44	1.22	-1.02								
			00	20	53.0	43	46	39	+	99	D										
0021+383		2	00	21	34.1	38	18	33		197	U	.59	-.86								
0021+395		0	00	21	38.7	39	32	29	+	54	R 6	.16	-.85								
0022+390		3	00	22	46.7	39	02	59		966	U	1.10	-.10	B	18.4	46.64	58.9	.7	.1		17
0022+394		0	00	22	24.4	39	29	18		45	U	.15	-.94								
0022+399		1	00	22	3.4	39	59	35		95	U	.20	-.58	B	18.5	3.51	36.4	-1.3	-1.4		
0022+424		3	00	22	54.9	42	27	15		370	U	.80	-.60	N	19.5	55.10	16.0	-2.2	-1.0		
0023+382		2	00	23	30.0	38	14	55		247	U	.65	-.76	N	19.0	29.83	58.0	2.0	-3.0		
0025+394		0	00	25	52.8	39	28	15	+	24	D 23	.12	-.84	N	20.0	53.55	5.0	-.6	1.5	4	
			00	25	54.2	39	27	58		17	D										
0026+397		0	00	26	50.6	39	46	09		14	D 64	.16	-1.42	G	0.198	50.89	14.0	-2.2	-3.0	2	
			00	26	50.7	39	47	11		12	D										
0027+380		2	00	27	41.9	38	00	55		87	D 14	.58	-1.08	B	19.3	42.10	47.2	.0	.8	4	
			00	27	42.3	38	00	41		60	D										
0027+395		2	00	27	3.4	39	32	04	+	170	R 10	.41	-.69	B	18.4	3.23	3.4	2.0	.6		6
0028+390		2	00	28	56.9	39	02	31		138	U	.47	-.96								
0028+394		2	00	28	29.3	39	25	25	+	77	D 25	.45	-.82								
			00	28	31.5	39	25	21	+	82	D										
0028+409	40.02	3	00	28	7.6	40	54	20		218	U	1.26	-1.38								
0028+450	45.02	4	00	28	6.7	45	05	14	+	90	D 34	1.62	-1.14	N	17.4	9.72	14.6	3.0	-.6	2	
			00	28	10.0	45	05	14		290	D										
0029+394		3	00	29	54.0	39	25	42	+	300	R 11	.97	-.92								
0029+398		1	00	29	34.9	39	50	38		125	U	.29	-.66								
0030+390		1	00	30	13.6	39	03	43		32	U	.20	-1.43								
0030+394		0										.11									
0030+396		0										.10									
0031+391	13	4	00	31	32.4	39	07	53		1199	D 26	6.41	-1.05								41
			00	31	33.7	39	07	31		479	D										

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0031+393		1	00	31	30.0	39	19	42		51	U	.26	-1.28								
0031+395		0										.10									
0031+396		1										.34									
0031+398		0										.11									
0032+390		1	00	32	29.2	39	02	35	+	88	R 31	.35	-1.08								
0032+394		1	00	32	43.8	39	24	15		112	U	.35	-.89								
0032+423	42.01	3	00	32	23.2	42	21	47		278	U	.92	-.94	B	* 1.588	23.26	49.6	-.7	-2.6	12	
0033+397		1	00	33	19.0	39	45	12	+	47	D 12	.29	-1.05								
			00	33	19.6	39	45	23	+	29	D										
0033+425	42.02	3	00	33	55.3	42	35	42	+	98	D 25	1.14	-1.10								
			00	33	55.4	42	36	07		183	D										
0034+387		2	00	34	37.1	38	42	46		151	U	.54	-1.00								
0034+393		1	00	34	54.2	39	21	42		222	U	.20	.08	B	* 1.937	54.24	42.4	-.5	-.4	3	
0034+444	44.02	4	00	34	9.6	44	26	51		607	U	2.25	-1.03								
0035+385A	38.03	3	00	35	1.3	38	31	35	+	343	D 82	2.55	-1.04								
0035+385B		3	00	35	8.1	38	31	16	+	334	D										
0036+398		0	00	36	49.8	39	52	07	+	38	R 9	.13	-.96								
0037+394		0	00	37	36.3	39	28	23		55	U	.12	-.61								
0037+396		0	00	37	34.0	39	38	37		45	U	.14	-.89								
0038+399		1	00	38	41.2	39	56	04	+	90	R 65	.22	-.70	G	0.247	40.64	54.4	6.4	9.6		
0039+373		4	00	39	24.3	37	23	10		844	U	2.00	-.68								
0039+391	39.02	3	00	39	11.6	39	08	54		276	U	1.02	-1.02								
0039+398	39.03	4	00	39	33.7	39	53	23		701	U	1.98	-.81								
0039+412		3	00	39	34.5	41	13	01		350	U	.91	-.75								
0040+470	46.01	4	00	40	37.1	47	00	19	+	184	D 14	1.83	-1.05								
			00	40	37.6	47	00	05	+	298	D										
0041+382A		2	00	41	5.9	38	13	45		206	U	.66	-.91								
0041+393		1	00	41	14.5	39	21	31	+	118	R 10	.33	-.81								
0041+405		2	00	41	10.0	40	30	10		148	U	.42	-.82								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0041+425	42.03	3	00	41	53.9	42	31	37		431	U	1.26	-.84								
0042+381A		2	00	42	24.7	38	06	53		271	U	.65	-.69								
0042+381B		2	00	42	32.1	38	08	07	+	72	T 98	.58	-.96								
			00	42	33.5	38	08	52		14	T										
			00	42	35.4	38	09	37	+	84	T										
0042+386		2	00	42	1.4	38	40	36		141	U	.57	-1.10								
0043+398		0	00	43	10.0	39	50	35	+	13	D 40	.10	-1.08	G	0.328	11.50	17.1	-4.6	2.4	4	27
			00	43	12.2	39	50	04	+	12	D										
0045+393		1	00	45	17.1	39	21	05		192	U	.22	-.11								
0045+395		0	00	45	10.4	39	32	37	+	94	R 11	.18	-.51	N	16.0	10.28	36.4	1.4	.6		4
0045+396		1	00	46	.0	39	37	20	+	107	R 28	.31	-.83	G	0.282	.14	22.7	-1.6	-2.7		27
0045+400	40.03	4	00	45	27.6	40	05	32	+	651	R 42	1.73	-.77	G	0.247	27.59	29.7	.1	2.3		4
0045+404		2	00	45	58.6	40	28	43	+	114	R 10	.56	-1.25								
0046+439		3	00	46	3.8	43	57	13	+	85	D 18	.96	-1.19								
			00	46	3.4	43	57	31	+	125	D										
0049+379	37.03	4	00	49	31.5	37	59	11		717	U	2.55	-1.00								
0049+395		0	00	49	19.4	39	34	20		11	D212	.15	-1.66								
			00	49	35.8	39	35	53		7	D										
0050+401		2	00	50	45.3	40	11	01		223	T446	1.22	-.98	G	0.153	45.11	10.0	-2.4	-2.0	3	
0050+403		2	00	50	45.6	40	18	23		111	T										
0050+402B		2	00	50	44.9	40	15	08		111	T										
0051+397		0	00	51	13.6	39	42	44		49	U	.15	-.88								
0051+404	40.04	4	00	51	40.5	40	25	56	+	761	R 12	2.44	-.91								
0052+380	38.04	3	00	52	10.1	38	05	35	+	351	R 28	1.25	-1.00								
0052+392		2	00	52	53.5	39	17	33	+	60	D 22	.41	-.99								
			00	52	54.0	39	17	11	+	56	D										
0052+395		1	00	52	53.8	39	32	02		79	U	.28	-.99								
0053+394		2	00	53	1.7	39	28	36		175	D 30	.60	-.88	G	0.198	2.71	43.0	-3.0	4.0	4	
			00	53	3.2	39	28	58		21	D										
0053+439	43.03	3	00	53	57.2	43	57	11	+	350	R 18	1.35	-1.06								
0054+396		1	00	54	42.1	39	41	05		81	U	.28	-.97								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0056+389		1	00	56	51.2	38	54	12	+	51	R 14	.22	-1.15								
0057+395		0	00	57	17.5	39	33	26	+	58	R 7	.19	-.93								
0057+397		0	00	57	3.2	39	42	37		7	U	.11	-2.15								
0058+403		2	00	58	4.8	40	19	52	+	241	R 9	.76	-.90								
0059+397		0	00	59	38.0	39	48	01	+	23	D 27	.11	-.85								
			00	59	38.6	39	48	28	+	14	D										
0059+461	46.02	4	00	59	53.1	46	08	39		590	D 32	3.92	-.95								
			00	59	55.8	46	08	23		583	D										
0100+388		2	01	00	55.1	38	48	35		169	U	.57	-.95								
0100+397		0	01	00	9.5	39	43	56	+	35	R 17	.11	-.90								
0103+422	42.04	4	01	03	51.5	42	13	53	+	580	R 8	2.37	-1.10								
0105+441		3	01	05	17.8	44	09	01		348	U	1.15	-.94								
0106+380		2	01	06	36.5	38	00	46		174	U	.45	-.74	B	15.9	36.49	47.8	.1	-1.8		
0106+397		2	01	06	35.1	39	44	04		151	U	.46	-.87								
0107+397		2	01	07	.8	39	42	30		165	U	.61	-1.03								
0107+398		0	01	07	53.9	39	51	25	+	35	R 15	.11	-.90	B	18.6	54.05	22.2	-1.7	2.8		
0107+399		1	01	07	13.2	39	55	16		57	U	.26	-1.19								
0108+396		1	01	08	10.5	39	36	09	+	45	D 20	.32	-.87								
			01	08	12.3	39	36	11	+	60	D										
0108+402		2	01	08	14.1	40	17	47	+	169	R 17	.66	-1.07								
0109+390		1	01	09	44.6	39	02	08		29	D 22	.28	-1.32								
			01	09	46.5	39	02	12		23	D										
0109+415		3	01	09	16.5	41	31	05	+	590	R 40	1.14	-.52	G	0.110	16.36	3.8	1.6	1.2		26
0109+416B	41.01	4	01	09	20.3	41	39	29	+	261	D 76	1.96	-.81	G?	0.180	24.70	28.9	-3.7	-13.7	5	
			01	09	26.8	41	39	07	+	437	D										
0110+386		2	01	10	6.3	38	37	28	+	78	D 22	.49	-.94								
			01	10	8.0	38	37	18	+	69	D										
0110+395		1	01	10	46.6	39	34	40		84	U	.30	-1.00								
0110+398		0	01	10	37.6	39	54	35		39	U	.13	-.94								
0110+401		3	01	10	26.5	40	10	19		474	U	1.08	-.65	B	20.0	26.48	19.6	.2	-.6		50

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0112+400		2	01	12	23.5	40	04	20		125	U	.54	-1.15								
0112+432	43.04	3	01	12	3.5	43	16	25	+	282	R 8	1.09	-1.06								
0113+400	40.05	4	01	13	32.3	40	01	27		285	D 20	1.85	-.86								
			01	13	33.0	40	01	45		330	D										
0114+399		2	01	14	34.7	39	57	00		66	D 22	.44	-1.18								
			01	14	36.6	39	57	06		32	D										
0115+394		1	01	15	3.5	39	28	46		88	U	.29	-.93								
0115+453A	36	4	01	15	3.5	45	20	35	+	1473	R 12	4.27	-.83								
0115+469	46.03	4	01	15	22.9	46	55	34	+	110	T116	1.77	-1.44								
			01	15	27.0	46	57	21	+	150	T										
			01	15	25.9	46	56	52		21	T										
0116+397		1	01	16	25.2	39	43	06		47	U	.21	-1.17								
0116+438		3	01	16	26.4	43	52	06	+	236	D 33	1.05	-.74	G	0.220	27.57	57.3	-2.1	-1.0	5	
			01	16	28.7	43	51	43	+	173	D										
0119+395		1	01	19	34.6	39	32	46	+	54	U	.21	-1.06								
0119+397		2	01	19	40.9	39	45	26	+	145	R 8	.48	-.94								
0120+380		2	01	19	59.5	38	02	51	+	44	D 23	.54	-1.23								
			01	20	.5	38	02	31		68	D										
0120+405	40.06	4	01	20	32.2	40	31	20		520	U	1.68	-.92								
0121+389		2	01	21	15.9	38	57	48	+	153	R 15	.45	-.85								
0122+395		0	01	22	33.9	39	30	17	+	43	D 17	.16	-.78								
			01	22	35.4	39	30	17		16	D										
0123+385		2	01	23	17.5	38	35	06		73	D 13	.49	-.99								
			01	23	17.9	38	35	19		66	D										
0123+396		1	01	23	31.9	39	38	39	+	143	R 24	.33	-.66	G	0.132	31.85	38.1	.6	.9		
0123+402		3	01	23	4.9	40	13	01		229	U	.81	-.99								
0124+387		2	01	24	50.8	38	44	39		129	U	.43	-.94								
0126+392A		1	01	26	33.8	39	14	27		71	U	.32	-1.18								
0127+395		0	01	27	47.6	39	34	48		39	U	.10	-.74								
0127+399		1	01	27	54.1	39	54	36		49	D 11	.34	-1.07								
			01	27	54.5	39	54	47		38	D										
0128+389		1										.20									

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TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0128+394		3	01	28	34.7	39	27	32		293	U	.92	-.90								
0130+381	38.05	3	01	30	48.6	38	07	38		384	U	1.07	-.80	G	0.392	48.66	37.4	-.7	.6		5
0130+384		2	01	30	4.3	38	25	54		81	D 37	.53	-1.01	B	18.7	5.06	38.1	2.2	.9	4	
			01	30	6.2	38	25	24		65	D										
0130+398		1	01	30	17.3	39	47	59	+	112	R 11	.28	-.72								
0131+390		2	01	31	48.6	39	03	50	+	91	D 93	.47	-.86	G	0.282	48.91	18.7	8.6	7.9	5	
			01	31	51.1	39	05	17	+	66	D										
0132+376A	46	4	01	32	29.4	37	38	25	+	635	D147	4.32	-1.05	G	* 0.437	34.08	47.0	3.7	3.8	5	37
0132+376B		4	01	32	40.8	37	39	24	+	495	D										
0132+392		1	01	32	35.6	39	12	48		89	D 77	.35	-.81								
			01	32	42.2	39	13	01		36	D										
0133+381		2	01	33	42.8	38	10	40		30	D 18	.47	-1.10								
			01	33	44.1	38	10	51	+	85	D										
0134+386	38.06	3	01	34	52.7	38	36	00	+	114	D 39	1.17	-.94	G	0.198	53.93	10.0	.2	3.0	4	
			01	34	55.2	38	36	26	+	239	D										
0134+389		1	01	34	20.1	38	56	20	+	93	R 8	.31	-.94								
0136+396	39.04	4	01	36	25.1	39	41	05	+	400	T237	3.89	-1.04	G	* 0.211	33.54	51.9	-.5	-.9	2	34
			01	36	33.5	39	41	51		37	T										
			01	36	43.4	39	42	53	+	600	T										
0137+385		2	01	37	50.0	38	33	19	+	72	D 42	.42	-.96								
			01	37	51.0	38	34	00	+	51	D										
0137+401		3	01	37	36.3	40	09	04		277	U	.82	-.85	B	18.5	36.18	3.9	1.4	.1		50
0138+394		2	01	38	31.5	39	25	30	+	108	R 12	.43	-1.08								
0139+389A		2	01	39	13.0	38	57	33		182	U	.43	-.67	B	20.8	13.02	34.4	-.2	-1.4		
0140+387	38.07	4	01	40	36.0	38	47	06		402	U	1.81	-1.18								
0141+398		1	01	41	33.8	39	48	11		78	U	.29	-1.03	N	14.5	33.59	8.2	2.4	2.8		
0143+393		1	01	43	2.4	39	18	15		85	U	.38	-1.17								
0143+446B	44.04	4	01	43	44.7	44	39	41	+	157	D 90	1.71	-.99	N	14.0	45.30	16.9	1.1	8.6	4	
			01	43	46.1	44	41	10	+	326	D										
0144+391		1	01	44	24.5	39	09	52		91	U	.25	-.79								
0144+399		1	01	44	17.9	39	58	52	+	55	R 15	.21	-1.05								
0144+430	43.05	3	01	44	55.8	43	04	47	+	255	R 19	.92	-1.01	B	18.6	55.86	48.1	-.7	-1.1		

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0144+432		3	01	44	53.6	43	17	19		319	U	.82	-.74								
0146+394		1	01	46	46.9	39	25	26		85	D132	.24	-.56								
			01	46	47.2	39	23	14	+	32	D										
0147+397		1	01	47	16.1	39	44	46	+	100	R 35	.26	-.75	G	0.328	16.32	49.9	-2.5	-3.9		
0147+398		0	01	47	36.9	39	52	23	+	45	R 46	.12	-.77	*G	0.124	36.50	27.0	4.6	-4.0		
0147+400		3	01	47	20.1	40	02	38		551	U	1.54	-.81								
0149+398		2	01	49	8.7	39	49	46	+	258	R 39	.59	-.65	G	0.080	8.66	47.3	.5	-1.3		
0150+406	40.07	3	01	50	46.6	40	41	19		159	D 41	.83	-1.15	N	16.0	47.88	6.5	3.6	3.0	4	
			01	50	49.8	40	41	00	+	32	D										
0152+382		2	01	52	17.5	38	16	52	+	97	D 40	.64	-1.10								
			01	52	20.0	38	16	24	+	60	D										
0152+435	54	4	01	52	25.3	43	30	58	+	987	D 49	5.27	-.98								
			01	52	27.2	43	31	43	+	522	D										
0153+417	41.02	4	01	53	19.6	41	47	28	+	328	D 47	2.14	-1.04	N	19.5	18.66	48.6	-.7	.4	4	8
			01	53	17.6	41	48	10	+	243	D										
0157+393A		1	01	57	7.4	39	19	26	+	35	D 29	.32	-.97								
			01	57	9.9	39	19	22		58	D										
0157+393B	39.05	3	01	57	47.2	39	19	00	+	109	T156	1.12	-.85	G	* 0.072	51.82	30.7	-.2	-1.7	3	34
			01	57	50.9	39	21	30	+	192	T										
			01	57	51.8	39	20	29	+	80	T										
0157+405A	40.08	4	01	57	3.9	40	35	45	+	659	D568	4.39	-.92	G	0.066	22.25	34.7	63.0	-11.7	4	
0157+405B		4	01	57	51.7	40	33	01	+	700	D										
0157+442	44.05	4	01	57	32.7	44	12	46		1021	U	3.26	-.91	N	19.0	32.74	44.6	-.4	1.4		8
0158+391		1	01	58	50.0	39	09	57	+	90	R 37	.26	-.83								
0158+394		0	01	59	.2	39	28	54		95	U	.12	-.18	B	18.4	.20	55.0	.0	-1.0		
0159+390		1	01	59	49.9	39	02	26	+	34	D 45	.24	-1.06	*G	0.198	50.87	1.0	8.5	14.0	4	
			01	59	53.3	39	02	04	+	28	D										
0159+397		1	01	59	18.2	39	45	12	+	72	R 20	.23	-.91								
0200+393		2	02	00	48.6	39	18	19	+	236	R 8	.59	-.72								
0201+390		1	02	01	43.0	39	05	04	+	102	R 12	.27	-.76								
0201+396		1	02	01	36.0	39	40	58		56	U	.28	-1.26								
0201+402		2	02	01	36.3	40	15	21	+	165	R 8	.47	-.82								

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TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0202+380		2	02	02	50.3	38	00	31	+	117	T131	.44	-.78	G	0.247	52.00	17.0	2.4	-2.5	2	
			02	02	52.2	38	01	14		6	T										
			02	02	54.5	38	02	32	+	40	T										
0205+395		2	02	05	4.5	39	30	47	+	160	R 8	.47	-.84								
0205+398		0	02	05	52.9	39	52	50	+	23	R 12	.10	-1.15								
0207+389		1	02	07	6.9	38	57	15	+	82	R 11	.22	-.77	G	* 0.018	7.01	19.0	-1.3	-4.0		9
0207+395		1	02	07	7.3	39	35	52	+	106	R 8	.36	-.96	N	18.0	7.47	51.6	-2.0	.4		
0207+397		1	02	07	52.4	39	47	32		138	U	.39	-.81	G	0.328	52.44	33.4	-.5	-1.4		
0207+399		0	02	07	12.6	39	59	26		41	U	.10	-.70								
0209+386		2	02	09	41.0	38	36	50	+	177	R 17	.64	-1.01								
0209+390		1	02	09	47.2	39	03	17		54	U	.23	-1.14	B	20.0	47.64	15.1	-5.1	1.9		
0209+394		0	02	09	33.8	39	28	15		69	U	.13	-.50	B	18.6	33.86	16.3	-.7	-1.3		
0210+396		0	02	10	30.1	39	40	07	+	39	R 8	.16	-1.11								
0211+393		2	02	11	7.1	39	19	13	+	72	D111	.52	-.81	G	0.198	1.88	58.3	-5.8	.1	5	
			02	10	57.7	39	18	49	+	112	D										
0213+392		1	02	13	11.1	39	16	56	+	42	D 49	.31	-.98								
			02	13	6.9	39	17	05	+	47	D										
0213+398		1	02	13	18.2	39	51	39	+	42	R 12	.21	-1.26								
0213+407		2	02	13	46.1	40	47	02		174	U	.51	-.84	G	0.247	46.36	2.6	-2.9	-.6		
0213+412	41.03	3	02	13	23.8	41	17	58		520	U	1.26	-.69								
0214+393		2	02	14	32.2	39	21	56		18	T 74	.77	-.77	G	0.117	34.12	35.2	2.0	5.6	5	
			02	14	33.3	39	22	29	+	138	T										
			02	14	35.6	39	22	59	+	134	T										
0216+388		1	02	16	48.6	38	52	00		57	U	.30	-1.30								
0216+393		1	02	16	30.4	38	18	51		130	U	.20	-.34								
0216+403		2	02	16	57.3	40	21	58	+	193	R 8	.57	-.85								
0216+423	42.06	4	02	15	58.8	42	19	04	+	429	D 53	2.30	-1.03								
			02	16	3.3	42	19	23	+	191	D										
0217+395		0	02	17	3.1	39	30	56		70	U	.16	-.65	N	18.5	3.16	56.7	-.7	-.7		
0217+417		3	02	17	4.5	41	43	34	+	97	T 61	.85	-.94	B	* 1.430	6.31	56.9	3.6	-5.6	4	2
			02	17	6.5	41	43	50		32	T										

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
			02	17	8.9	41	44	10	+	126	T										
0218+396	39.06	4	02	18	44.9	39	41	18	+	404	D 87	2.93	-1.07								
			02	18	40.4	39	42	28	+	342	D										
0218+397		0										.14									
0218+399A		1	02	18	36.5	39	55	01	+	101	D120	.85	-1.13								
0218+399B		2	02	18	44.9	39	55	48	+	100	D										
0218+402A		2	02	18	58.0	40	15	10	+	51	D 88	.69	-1.00	N	13.5	56.27	41.6	-8.8	-5.1	4	
0218+402B			02	18	53.0	40	14	03	+	142	D										
0219+397		2	02	19	33.3	39	47	02		111	U	.44	-1.08								
0219+421		3	02	19	24.1	42	07	17	+	1100	R270	1.18	-.05	G	0.002	24.66	18.6	-6.2	-1.6		14
0219+428A	66A	4	02	19	30.1	42	48	32		1670	U	5.05	-.87	*B	16.0	30.00	30.0	1.1	2.0		
0219+443		3	02	19	5.1	44	18	57	+	112	D 47	.89	-1.17	B	17.3	6.20	17.8	-.5	.2	4	
			02	19	7.2	44	19	39	+	88	D										
0220+388		2	02	20	4.4	38	49	31		145	U	.54	-1.03								
0220+393A		3	02	20	28.2	39	22	24	+	391	R 21	1.04	-.77	G	0.220	28.34	28.1	-1.6	-4.1		
0220+397	65	4	02	20	36.6	39	47	17	+	2853	R 17	9.08	-.91								
0220+427A	66	4	02	20	10.9	42	46	02	+	3000	D 90	14.83	-.63	*G	* 0.021						22
0219+427D		4	02	20	2.8	42	46	02	+	3600	D										
0221+383		2	02	21	8.8	38	18	49		94	D 85	.43	-.93								
			02	21	15.6	38	18	19		38	D										
0221+393		1	02	21	36.3	39	17	54		22	D164	.21	-1.17								
			02	21	38.2	39	20	35	+	25	D										
0221+396		1	02	21	44.4	39	41	44	+	48	R 57	.22	-1.19								
0222+397		0										.14									
0222+403	40.09	4	02	22	36.7	40	18	51	+	530	T 92	2.19	-.61	G	0.117	37.66	2.4	-.7	.6	2	
			02	22	37.6	40	18	03		61	T										
			02	22	38.4	40	17	21	+	410	T										
0222+422A		3	02	22	22.6	42	16	10		250	U	.81	-.92								
0224+393		3	02	24	.9	39	18	15		374	U	.89	-.68	B	19.5	.97	14.2	-.8	.8		
0224+396		0	02	24	52.0	39	36	28	+	61	R 10	.18	-.85								
0225+381		2	02	25	23.7	38	08	02	+	50	D 37	.47	-1.07	G	0.050	25.09	49.0	.8	-.4	5	
			02	25	26.2	38	07	39	+	70	D										

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0225+389		1	02	25	54.1	38	54	36	+	64	D131	.22	-.70	B	15.8	52.93	22.5	11.9	-8.7	5	
			02	25	53.6	38	56	47	+	26	D										
0225+427		3	02	25	44.1	42	47	31		244	U	.86	-.99								
0226+394		2	02	26	38.6	39	28	52	+	29	T 86	.40	-.79	G	0.247	39.16	52.0	11.9	-8.7	5	
			02	26	42.1	39	29	21	+	63	T										
			02	26	41.1	39	30	13	+	54	T										
0226+396		1	02	26	49.6	39	39	47	+	44	R 12	.21	-1.22								
0226+467	46.05	4	02	26	5.0	46	46	58	+	747	R 8	2.45	-.93	N	20.0	5.20	54.0	-2.0	4.0		8
0227+397		1	02	27	22.7	39	45	27	+	76	R 20	.23	-.87								
0227+398		2	02	27	59.7	39	50	15		207	U	.68	-.93								
0228+392		1	02	28	.8	39	14	05		83	U	.38	-1.19	N	13.5	1.16	2.0	-4.2	3.0		
0228+393	39.08	4	02	28	41.3	39	19	43		382	D 13	2.62	-.97								
			02	28	42.2	39	19	52		382	D										
0228+409A		3	02	28	29.6	40	56	39		322	U	.90	-.81								
0231+385		2	02	31	21.6	38	31	47	+	185	R 10	.47	-.73								
0231+391		1	02	31	43.3	39	10	49		33	D 49	.25	-1.00								
			02	31	46.9	39	10	23		37	D										
0231+405A		2	02	31	50.9	40	30	52		219	U	.43	-.53								
0232+411B	41.04	3	02	32	41.6	41	09	55	+	430	D109	2.75	-.99	B?	17.4	45.82	13.5	-1.0	-2.4	5	
0232+411C		3	02	32	50.8	41	10	31	+	350	D										
0233+390		1	02	33	41.4	38	59	56		106	D 37	.34	-.84								
			02	33	44.5	39	00	08		11	D										
0236+399		1	02	36	58.8	39	59	06		41	U	.25	-1.42								
0236+438		3	02	36	52.4	43	48	11		102	D 54	1.05	-1.28								
			02	36	52.9	43	49	05		102	D										
0237+389		1	02	37	45.2	38	58	38	+	30	D 70	.24	-1.00								
			02	37	39.2	38	58	48	+	37	D										
0237+396		0	02	37	46.1	39	37	48		26	U	.12	-1.20								
0237+435		3	02	37	14.2	43	35	15		276	D 18	1.25	-.92								
			02	37	14.2	43	34	57		110	D										
0239+395		2	02	39	47.3	39	32	02	+	123	T 68	.64	-.86	G	0.180	47.35	39.6	2.7	1.0	5	
			02	39	47.8	39	31	20	+	60	T										
			02	39	48.3	39	30	54	+	30	T										

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0239+397		1	02	39	44.6	39	43	02		124	U	.37	-.86	B	20.0	44.63	.7	-.3	1.3		
0240+404		2	02	40	22.0	40	29	12		138	D113	.40	-.68	G	0.220	24.67	2.2	-10.9	5.0	5	
			02	40	31.6	40	28	45	+	30	D										
0241+393B	39.09	4	02	41	27.8	39	22	12	+	200	T150	3.23	-1.16	G	0.247	23.76	35.7	-24.2	-15.4	4	
0241+393A		3	02	41	17.1	39	20	47		166	T										
			02	41	20.1	39	21	02	+	368	T										
0241+395		1	02	41	29.2	39	31	03	+	86	R 10	.28	-.92								
0242+395		1	02	42	18.4	39	34	06		68	U	.26	-1.05								
0243+439		3	02	43	2.1	43	59	02		285	D 17	1.01	-.62								
			02	43	2.7	43	59	18		172	D										
0244+377	37.07	4	02	44	23.8	37	42	21		564	U	1.73	-.88								
0246+392		2	02	46	22.6	39	16	51	+	142	R 33	.46	-.92								
0246+393	39.10	4	02	46	58.8	39	22	21	+	674	T145	6.17	-.97	G	0.198	4.78	8.7	-.9	.3	2	18
			02	47	4.7	39	22	09		21	T										
0247+393A		4	02	47	11.3	39	22	08	+	1100	T										
0246+396		1	02	46	48.0	39	39	18		79	U	.25	-.90								
0246+428A	42.08	4	02	46	10.8	42	52	13	+	900	D140	4.97	-.81	G	0.132	15.18	.0	3.5	5.4	5	49
0246+428B		3	02	46	20.3	42	53	59	+	880	D										
0247+395		0	02	47	10.0	39	30	28	+	48	R 30	.12	-.72								
0247+404		3	02	47	58.2	40	29	53	+	307	R 15	.94	-.88								
0248+392		2	02	48	46.7	39	16	13		114	D 15	.61	-.94								
			02	48	45.6	39	16	04		69	D										
0248+396		0	02	48	44.7	39	39	13		36	U	.13	-1.00								
0248+467		4	02	47	38.4	46	45	03	+	990	R600	2.40	-.69	G	0.029	39.50	55.0	-11.2	8.0		
0249+383		3	02	49	59.0	38	23	11		616	U	.95	-.34	B?	18.3	58.98	10.9	.2	.1		
0250+384	38.08	4	02	50	41.2	38	29	50		159	D 47	1.64	-1.24								
			02	50	42.3	38	29	04		180	D										
0250+396		2	02	50	6.9	39	41	38		49	D 24	.42	-.88								
			02	50	8.5	39	41	22	+	87	D										
0251+393		1	02	51	31.1	39	19	43		217	D 26	.27	-.10								
			02	51	33.2	39	19	20		21	D										
0252+385		2	02	52	59.2	38	30	40	+	38	D 29	.42	-1.16								
			02	53	.5	38	30	15	+	58	D										

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0252+388		1	02	52	50.0	38	52	17	+	61	R 10	.20	-.93								
0252+399		0	02	52	35.0	39	54	13		25	U	.10	-1.08								
0253+396		0	02	53	35.0	39	39	52	+	34	R 8	.14	-1.11								
0254+406	40.10	3	02	54	36.6	40	38	33		438	U	1.43	-.93								
0255+460	46.07	4	02	55	8.0	46	04	08		654	U	1.74	-.77	B	21.0	8.10	6.0	-1.0	2.0		8
0258+435	43.07	3	02	58	53.9	43	31	00	+	470	R 80	1.31	-.80	G	* 0.065	54.40	4.0	-5.4	-4.0		16
0258+443		3	02	58	13.7	44	18	38		440	U	1.09	-.71								
0259+387		2	02	59	47.0	38	44	56	+	65	D 34	.43	-.88	G	0.132	48.30	55.0	1.7	3.0	4	
			02	59	49.9	38	45	00	+	75	D										
0259+391		1	02	59	18.0	39	09	11	+	100	R 21	.28	-.81								
0700+375	37.18	4	07	00	41.1	37	31	27		540	U	2.07	-1.05								
0700+390		1	07	00	44.7	39	04	06	+	70	D194	.23	-.24								
			07	00	32.2	39	01	55	+	100	D										
0700+398		0	07	00	37.8	39	52	55		20	U	.11	-1.33								
0700+399		0	07	00	56.8	39	53	51	+	75	R 30	.13	-.43								
0701+392		3	07	01	5.1	39	15	54		449	U	1.17	-.75	N	18.7	5.07	54.3	.3	-.3		50
0701+397		0	07	01	44.1	39	42	22	+	22	R 40	.11	-1.26								
0701+401		4	07	01	53.3	40	06	53		172	D 18	1.67	-.80	G?	0.328	53.06	7.4	5.2	-1.6	5	
			07	01	53.6	40	07	11		428	D										
0702+396		0	07	02	24.2	39	39	25		41	U	.13	-.90								
0703+390		1	07	03	22.8	39	00	12		112	U	.31	-.80								
0703+426A	42.23	4	07	03	5.8	42	38	42	+	420	T226	4.30	-.57	*G	* 0.060	12.18	17.3	2.4	-3.3	3	16
			07	03	10.6	42	36	38	+	1080	T										
			07	03	12.4	42	35	14	+	580	T										
0703+426B		3	07	03	30.6	42	36	53		212	D100	1.04	-.85								
			07	03	30.6	42	38	18		138	D										
0703+468		4	07	03	5.9	46	52	36		1598	U	1.90	-.14								
0704+384	38.20	4	07	04	7.6	38	26	56		439	D 16	2.87	-1.06	B	* 0.579	8.41	56.8	-1.3	.2	4	36
			07	04	9.0	38	26	58		307	D										
0704+397		0	07	04	5.8	39	46	24		36	U	.14	-1.06								
0704+399		1	07	04	17.8	39	54	19		67	U	.20	-.86								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0704+418		3	07	04	47.3	41	53	19	+	103	T 60	.84	-1.05								
			07	04	49.6	41	53	40		42	T										
			07	04	51.3	41	54	00	+	74	T										
0705+398		0	07	05	52.3	39	53	24		42	U	.19	-1.18	G?	0.392	51.94	22.2	4.1	1.8		
0706+396		1	07	06	34.5	39	41	39	+	103	R 17	.35	-.96								
0707+380A		2	07	07	16.1	38	04	14		157	U	.43	-.79								
0708+384		2	07	08	25.5	38	28	55		135	U	.46	-.96								
0708+388		2	07	08	43.1	38	53	28	+	210	R 8	.66	-.90	G	0.282	42.80	26.8	3.5	1.2		
0709+393		1	07	09	35.8	39	18	56	+	50	D100	.29	-1.38								
			07	09	22.5	39	18	56			D										
0709+398		1	07	09	22.6	39	53	51	+	116	R 8	.26	-.63								
0709+405		2	07	09	32.7	40	33	11	+	163	R 18	.49	-.86								
0709+409	40.18	3	07	09	1.0	40	56	42		160	D 21	1.29	-1.11	N	20.0	1.01	32.9	1.5	-1.2	5	
			07	09	1.3	40	56	21		154	D										
0710+403		2	07	10	56.3	40	21	03	+	48	D 26	.54	-1.01	G	0.247	56.94	15.3	-1.0	-.8	4	
			07	10	57.4	40	21	26		100	D										
0710+457	45.13	4	07	10	49.3	45	45	27		573	D 64	3.00	-.97	G	0.165	52.28	19.3	.2	.2	4	
			07	10	55.3	45	45	12	+	300	D										
0711+397		2	07	11	51.5	39	45	10		200	U	.79	-1.08								
0711+399		0	07	11	5.3	39	58	00	+	30	U	.10	-.94								
0712+382		2	07	12	26.8	38	13	26		231	U	.50	-.61								
0712+387		2	07	12	8.6	38	44	38		121	D 42	.71	-.96								
			07	12	10.9	38	44	05	+	88	D										
0712+388		1	07	12	24.5	38	53	54	+	101	R 12	.25	-.71								
0712+390		1	07	12	12.0	39	05	23		74	U	.25	-.95								
0712+391		2	07	12	3.9	39	06	15		74	T 25	.50	-1.01								
			07	12	3.7	39	06	39		64	T										
			07	12	3.7	39	06	26			T										
0714+405		2	07	14	42.8	40	30	20		94	D 39	.40	-.87								
			07	14	44.0	40	30	56		97	D										
0717+393		1	07	17	21.3	39	21	46	+	21	D 24	.31	-1.04								
			07	17	23.3	39	21	54		61	D										

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0720+381		3	07	20	42.2	38	08	55	+	145	D 50	.90	-.87	G	0.220	44.03	41.4	-1.5	-.4	2	
			07	20	43.9	38	09	41	+	150	D										
0720+412		3	07	20	11.5	41	14	09	+	204	R 20	.96	-1.21								
0721+394		1	07	21	34.3	39	27	18		52	U	.20	-1.06								
0721+398		1	07	21	.4	39	49	48		64	U	.22	-.97								
0722+385		2	07	22	21.0	38	32	39		189	U	.59	-.89								
0722+393A	39.17	4	07	22	25.1	39	23	32		1010	U	2.91	-.83								
0723+397		2	07	23	28.2	39	46	00	+	47	D 17	.41	-1.20								
			07	23	28.7	39	45	44	+	42	D										
0724+396		1	07	24	45.0	39	36	30	+	50	D 7	.24	-1.09	G?	0.282	45.00	35.0	.6	-1.5	4	
			07	24	45.1	39	36	37	+	10	D										
0726+402		2	07	26	37.9	40	17	19		272	U	.79	-.84								
0726+431	43.14	3	07	26	16.1	43	07	28	+	123	T 19	1.29	-1.17	B	* 1.072	16.70	35.5	.0	-.5	2	48
			07	26	16.7	43	07	35		14	T										
			07	26	17.3	43	07	42	+	154	T										
0727+401		2	07	27	10.0	40	07	42		367	U	.40	-.07								
0728+389		1	07	28	34.4	38	57	51		98	U	.20	-.56								
0728+395		1	07	28	30.8	39	29	59	+	79	R 24	.30	-1.05								
0729+391		1	07	29	57.1	39	11	36		151	U	.26	-.43	B	18.4	57.02	34.9	.9	1.1		
0729+395		1	07	29	36.5	39	31	35	+	80	R 8	.28	-.98	G	0.110	36.28	35.0	2.5	.0		
0729+397		0	07	29	33.2	39	45	07		53	U	.15	-.81								
0729+437		3	07	29	11.1	43	42	07		295	U	.90	-.87								
0730+396		1	07	30	25.0	39	41	51		57	U	.21	-1.02								
0731+438	43.15	4	07	31	49.4	43	50	57	+	745	R 10	2.60	-.98								
0733+389		1	07	33	33.3	38	59	24	+	27	D 63	.24	-1.21								
			07	33	38.4	38	59	02	+	24	D										
0735+388		1	07	35	31.9	38	53	36	+	25	T 55	.36	-.99	G	0.180	32.13	19.6	-.3	-1.6	2	
			07	35	32.1	38	53	18	+	22	T										
			07	35	34.2	38	52	47	+	55	T										
0735+390		2	07	35	37.5	39	03	04	+	34	D 21	.44	-.91	G	0.282	38.50	5.5	-1.2	-.5	4	
			07	35	39.3	39	03	06	+	104	D										
0735+395		1	07	35	41.0	39	32	37		108	U	.28	-.75								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0736+386		2	07	36	48.5	38	40	45		253	U	.77	-.87								
0736+398		0	07	36	4.3	39	54	04	+	66	R 60	.14	-.59								
0736+400		2	07	36	41.5	40	01	47		191	U	.53	-.80	G	0.198	41.17	43.3	3.8	3.7		
0739+396		2	07	39	27.7	39	36	43	+	80	D 27	.49	-1.16								
			07	39	29.1	39	36	21	+	32	D										
0739+397a		3	07	39	44.8	39	46	22		79	D 25	1.02	-1.62	G?	0.198	45.83	28.8	-.6	-.9	5	
			07	39	47.3	39	46	37		51	D										
0739+397b		3	07	39	46.0	39	48	41		495	U	1.02	-.57	B	19.2	45.96	39.5	.5	1.5		
0739+398		0	07	39	13.2	39	51	42		129	U	.17	-.22	B	19.9	13.28	42.4	-.9	-.4		
0740+380C	186	4	07	40	56.8	38	00	31		1072	U	5.55	-1.29	B	* 1.063	56.75	31.6	.6	-.6		12
0740+393	39.18	4	07	40	35.7	39	22	14	+	457	R 10	2.08	-1.19								
0740+474B	47.26	4	07	40	48.8	47	25	56	+	521	R 12	1.63	-.89								
0741+396		0	07	41	7.1	39	41	47	+	22	R 60	.14	-1.45								
0741+399		0	07	41	57.0	39	58	21		46	U	.14	-.87								
0741+407		2	07	41	52.3	40	45	45		56	D 18	.49	-1.06								
			07	41	53.6	40	45	34		70	D										
0742+376	37.19	4	07	42	18.8	37	39	13		427	D 60	2.10	-1.13								18
			07	42	22.6	37	38	33		71	D										
0742+394		2	07	42	39.5	39	24	19	+	123	R 21	.45	-1.02								30
0743+392B		3	07	43	53.8	39	17	19	+	270	R 15	.87	-.92								
0743+399		0	07	43	7.0	39	58	43	+	29	R 13	.10	-.97								
0744+399		1	07	44	9.5	39	57	38		74	U	.20	-.78								
0744+464	46.16	4	07	44	7.2	46	26	25		444	U	1.66	-1.03								
0745+397		0	07	45	38.9	39	46	23		40	U	.11	-.79								
0745+398		2	07	45	18.5	39	48	38	+	100	R 11	.41	-1.11								
0746+399		2	07	46	59.6	39	58	51	+	190	R 20	.75	-1.08								
0747+398A	39.19	3	07	47	2.6	39	49	12		283	U	.97	-.97								
0747+398B		3	07	47	57.2	39	51	47	+	120	R 8	.88	-1.56								
0748+413B		3	07	48	54.1	41	23	37		198	U	.83	-1.12								

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TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0749+398	39.20	3	07	49	15.4	39	52	02		265	U	1.02	-1.06								
0749+399		0	07	49	28.5	39	54	35		53	U	.12	-.64								
0750+400		2	07	50	56.9	40	03	21	+	54	D 53	.44	-.96								
			07	50	59.2	40	04	07	+	75	D										
0750+402		2	07	50	15.7	40	12	54	+	61	D 44	.45	-1.15								
			07	50	18.7	40	12	26	+	43	D										
0751+392		3	07	51	5.6	39	16	04		220	U	.80	-1.01								
0752+398		1	07	52	7.2	39	51	11		33	D 28	.21	-.98								
			07	52	9.4	39	51	25	+	27	D										
0753+383	38.22	3	07	53	30.3	38	18	19		321	D 60	1.51	-1.03								
			07	53	32.2	38	17	23		84	D										
0753+391		2	07	53	3.8	39	09	03	+	69	D 65	.48	-.97	G	0.328	6.20	25.0	-4.3	1.3	5	
			07	53	7.8	39	09	49	+	71	D										
0754+394		0	07	54	41.3	39	28	50	+	35	R 40	.13	-1.03	*B	* 0.096	38.20	34.0	35.8	16.0		52
0754+396	39.21	3	07	54	46.9	39	37	40		449	U	1.53	-.96								
0754+397		0	07	54	36.9	39	44	45	+	25	R 30	.17	-1.50	G	0.282	36.80	45.0	1.1	.0		
0755+379B		4	07	55	9.0	37	55	20	+	2285	R140	5.42	-.68	G	* 0.043	9.05	20.9	-.6	-.9		6
0756+377	37.21	4	07	56	28.4	37	47	04	+	1290	R 19	4.54	-.99								
0756+383		3	07	56	32.9	38	22	37	+	354	R 25	.97	-.79	G	0.153	32.73	38.9	2.0	-1.9		
0756+406		2	07	56	52.6	40	38	11		167	U	.50	-.86	B	18.8	52.64	9.8	-.5	1.2		
0757+395		1	07	57	29.6	39	34	12	+	35	T160	.29	-.77	G	0.057	32.35	55.4	-1.7	1.6	2	
			07	57	32.2	39	32	57	+	40	T										
			07	57	32.5	39	31	35	+	33	T										
0757+399		1	07	57	10.6	39	55	24		91	U	.32	-.99								
0759+392		1	07	59	2.5	39	12	38	+	29	D 43	.20	-.87								
			07	59	3.2	39	13	21	+	37	D										
0759+397		1	07	59	40.9	39	45	02		137	U	.39	-.82								
0800+399		1	08	00	39.3	39	58	35		86	U	.22	-.74	B	21.0	39.00	38.0	3.4	-3.0		
0800+472	47.27	4	08	00	38.2	47	13	12		809	U	2.06	-.73	B	21.0	38.30	14.0	-1.0	-2.0		8
0801+394		0	08	01	17.8	39	29	06	+	43	R 12	.18	-1.12								
0801+399		0	08	01	11.1	39	58	28		32	U	.18	-1.35								

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TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0801+401		2	08	01	58.9	40	06	18		261	U	.58	-.63								
0802+398		2	08	02	36.6	39	49	21		113	U	.41	-1.01	B	20.6	36.63	22.7	-.3	-1.7		
0802+406		2	08	02	8.6	40	36	35		175	U	.49	-.81								
0803+426		3	08	03	16.1	42	41	30	+	88	R 20	.42	-1.22	*G	0.153	18.40	33.0	-25.2	-3.0		
0803+427		3	08	03	40.9	42	43	41	+	80	D 56	.58	-1.24	G	0.142	40.90	30.0	.7	-4.0	5	
			08	03	41.1	42	42	56	+	40	D										
0804+399		0	08	04	25.1	39	55	23	+	38	R 13	.15	-1.07	G	0.220	25.52	19.5	-4.8	3.5		
0805+391		1	08	05	2.4	39	09	47	+	75	R 20	.23	-.88	G	0.198	2.45	45.5	-.6	1.5		
0805+392		1	08	05	24.4	39	15	36	+	60	R 17	.23	-1.05								
0805+406	40.19	3	08	05	40.7	40	41	45		438	U	1.33	-.87	B	21.0	40.64	44.6	.7	.4		
0806+399		2	08	06	11.4	39	58	20		246	U	.49	-.54								
0806+426	194	4	08	06	37.9	42	36	58	+	1970	R 18	5.68	-.83	G	* 0.312	37.58	55.1	3.5	2.9		37
0807+399		1	08	07	24.3	39	58	29	+	30	T 87	.23	-1.08								
			08	07	22.2	39	59	09		10	T										
			08	07	19.7	39	59	37	+	18	T										
0809+404	40.20	4	08	09	31.6	40	28	03		955	U	2.31	-.69	G?	0.282	31.61	3.1	-.1	-.1		8
0810+460B	46.17	4	08	10	58.6	46	05	48		1059	U	4.08	-1.06	G	0.328	58.57	47.5	.3	.5		
0811+388	38.23	3	08	11	53.8	38	50	09	+	307	D 29	1.51	-.93	G	* 0.132	53.78	57.0	1.4	2.2	5	46
			08	11	54.1	38	49	40	+	156	D										
0811+391		1	08	11	4.6	39	09	29		21	D135	.27	-1.01								
			08	11	11.2	39	07	37		53	D										
0812+382	38.24	3	08	12	44.2	38	13	29	+	409	R 31	.97	-.68	G	0.132	44.41	30.7	-2.5	-1.7		50
0812+398		0	08	12	38.9	39	52	45	+	20	R 30	.10	-1.26								
0812+399		0	08	12	38.1	39	58	26	+	30	R 23	.12	-1.08								
0812+406		2	08	12	46.9	40	38	33		213	U	.70	-.93	N	18.4	46.79	34.0	1.2	-1.0		
0813+381		2	08	13	34.7	38	08	22	+	125	R 11	.44	-.99								
0813+393		1	08	13	33.3	39	20	25	+	49	D 31	.21	-.85	G	0.328	34.62	34.2	-1.4	-1.7	4	
			08	13	35.7	39	20	40	+	22	D										
0813+398		1	08	13	23.1	39	53	06		70	D 84	.31	-.78	G	0.282	29.63	46.0	-.3	.0	2	
			08	13	29.6	39	53	46	+	45	D										

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0814+383		2	08	14	2.0	38	20	47		138	U	.56	-1.10								
0814+425		3	08	14	51.7	42	32	08		1212	U	1.42	-.12	B	19.1	51.70	6.8	0.0	-1.2		
0814+441		3	08	14	8.6	44	08	55		241	U	.84	-.98								
0815+397		0	08	15	25.0	39	47	40	+	43	R 23	.16	-1.03								
0818+472A	197.1	4	08	18	.9	47	12	10	+	1651	R 18	4.68	-.82	G	* 0.128	.90	11.2	.0	-1.2		42
0819+397		0	08	19	10.5	39	45	26		50	U	.19	-1.05								
0820+392		2	08	20	5.5	39	16	18		86	D 26	.58	-1.04	G	0.282	5.65	20.2	-1.7	-2.2	1	
			08	20	6.7	39	15	56		67	D										
0820+431	43.16	4	08	20	2.5	43	06	36		456	D 9	3.06	-.95	N	19.5	2.76	37.2	1.5	.8	4	8
			08	20	3.3	43	06	40		456	D										
0821+394	39.23	4	08	21	37.4	39	26	29		1237	U	2.59	-.58	B	* 1.216	37.31	27.7	1.0	1.3		36
0821+395		0	08	21	30.0	39	33	33		14	D 55	.13	-1.23								
			08	21	34.1	39	33	04		13	D										
0821+399		0										.11									
0821+447	44.17	4	08	21	49.5	44	46	23		483	D 23	2.26	-.90	B	* 0.893	50.27	14.2	.8	1.3	4	44
			08	21	51.2	44	46	08		235	D										
0822+390		1	08	22	12.6	39	02	02	+	80	R 9	.22	-.79								
0822+394		4	08	22	5.5	39	29	33		1187	U	1.76	-.31								
0822+398		0	08	22	12.5	39	50	54		26	U	.15	-1.37								
0823+384B		2	08	23	56.7	38	29	53	+	174	R 30	.41	-.67	G	0.198	56.74	48.4	-.5	4.6		
0823+399		0	08	23	35.3	39	58	29		26	D 16	.12	-.77	N	16.5	35.30	37.3	1.1	-.3	4	
			08	23	35.5	39	58	45		19	D										
0824+397		1	08	24	47.0	39	45	20	+	60	D 58	.27	-.95	B	17.7	48.59	38.5	3.0	1.5	4	
			08	24	50.7	39	46	00	+	20	D										
0827+378	37.24	4	08	27	55.2	37	52	18		1873	U	5.17	-.80	B	* 0.914	55.10	18.2	1.2	-.2		12
0827+387		2	08	27	8.8	38	47	57		143	U	.60	-1.12								
0827+395		0	08	27	1.6	39	33	52	+	50	R 20	.11	-.62	G	0.165	1.76	51.8	-1.8	.2		
0827+399		0										.11									
0827+458	45.16	4	08	27	7.9	45	53	41		270	D 16	2.36	-1.16								
			08	27	6.6	45	53	31		267	D										
0828+381		2	08	28	35.8	38	06	43	+	170	R 30	.46	-.78	G	0.220	35.90	46.0	-1.2	-3.0		

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0829+395		2	08	29	58.3	39	31	34		190	U	.63	-.94								
0829+425		3	08	29	26.4	42	35	13	+	400	R 4	.82	-.56	B	18.5	26.28	12.5	1.3	.5		
0831+393		1	08	31	38.4	39	21	13		73	U	.21	-.83								
0831+399		1	08	31	12.4	39	55	34	+	60	D 99	.25	-.85								
			08	31	20.2	39	54	52	+	25	D										
0832+395		1	08	32	14.6	39	33	09		81	U	.20	-.71								
0832+399		0	08	32	19.9	39	55	24	+	15	T 56	.11	-.85								
			08	32	22.1	39	55	04		6	T										
			08	32	23.5	39	54	46	+	16	T										
0834+399		0	08	34	8.9	39	57	53	+	39	R 10	.10	-.74								
0834+450A	45.17	4	08	34	26.1	44	59	55	+	650	D119	3.45	-.77	G	0.153	27.30	57.3	-2.1	-3.3	4	
			08	34	28.1	45	01	53	+	650	D										
0836+399		0	08	36	55.2	39	54	29	+	34	R 5	.12	-.99								
0836+402	40.21	4	08	36	54.1	40	15	03	+	140	D 49	1.87	-.98	G?	0.282	54.37	39.9	-1.4	-1.4	4	
			08	36	54.4	40	14	14	+	395	D										
0836+426	42.26	3	08	36	35.6	42	38	33		488	U	1.12	-.65	B	20.0	35.50	34.1	1.1	-1.1		
0837+399		0	08	37	50.3	39	55	31		28	D 26	.16	-1.24								
			08	37	52.6	39	55	30		5	D										
0838+396		0	08	38	54.6	39	37	58		51	U	.15	-.84								
0840+400		2	08	40	50.5	40	03	53		126	U	.58	-1.20								
0840+424A		4	08	40	11.5	42	26	20		1125	U	2.28	-.55								
0841+386		2	08	41	14.8	38	41	49		311	U	.49	-.36								
0841+397		0	08	41	6.0	39	45	03	+	19	R 8	.13	-1.50								
0841+403		3	08	41	37.9	40	19	09	+	226	D 34	1.12	-.91	N	19.5	39.40	10.1	.0	-1.1	4	50
			08	41	40.9	40	19	09	+	123	D										
0841+407		2	08	41	52.4	40	42	08		89	D 24	.40	-.74								
			08	41	50.5	40	42	17		67	D										
0842+401		2	08	42	47.6	40	07	55	+	200	R 35	.57	-.82								
0843+425		3	08	43	57.5	42	33	50	+	130	T 96	.91	-.80	N	18.5	58.65	44.8	3.8	-2.8	4	
			08	43	59.2	42	34	55		32	T										
			08	44	.3	42	35	21	+	164	T										
0844+396		0	08	44	35.5	39	41	06	+	34	R 8	.11	-.92								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0847+406		2	08	47	24.8	40	40	28	+	61	R 14	.48	-1.62								
0849+394		0	08	49	24.9	39	30	10		17	D 32	.11	-.92								
			08	49	24.9	39	30	42		17	D										
0849+424	42.27	3	08	49	15.6	42	26	40		205	D 15	1.31	-.90	B	18.0	15.71	47.5	-.7	.0	4	8
			08	49	15.7	42	26	55		212	D										
0850+383		2	08	50	42.1	38	23	06		32	D 56	.40	-1.09								
			08	50	40.6	38	22	13		67	D										
0852+384		2	08	52	43.2	38	25	03	+	141	R 16	.44	-.89								
0854+399B		3	08	54	34.0	39	56	18	+	234	D159	1.46	-1.16	G?	0.282	29.70	4.5	-8.6	8.0	4	1
0854+399A		2	08	54	23.9	39	58	07	+	100	D										
0855+397		0	08	55	2.1	39	42	33	+	33	R 55	.10	-.87								
0855+419		3	08	55	44.0	41	54	59	+	270	R 12	.96	-.99								
0856+386		2										.40									
0856+397		0	08	56	24.9	39	41	56		19	D 20	.12	-.86								
			08	56	25.9	39	42	12	+	20	D										
0856+406	40.22	3	08	56	45.1	40	36	18		229	U	1.07	-1.21								
0857+391		3	08	57	41.3	39	08	08		283	D 18	1.28	-.87	G	0.220	41.63	2.6	-1.2	-.2	5	50
			08	57	42.0	39	07	51	+	139	D										
0857+402		2	08	57	15.9	40	16	43		194	U	.56	-.83								
0858+386		2	08	58	20.8	38	38	59		284	U	.74	-.75								
0858+388		2	08	58	3.0	38	53	55	+	74	R 8	.42	-1.36								
0858+452	45.18	4	08	58	54.7	45	12	43	+	728	R 10	2.17	-.86								
0859+470	47.29	4	08	59	39.9	47	02	56		2145	U	2.81	-.21	B	* 1.462	39.90	57.0	.0	-1.0		44
0900+380B		2	09	00	49.8	38	04	28	+	133	R 24	.45	-.96								
0900+389		2	09	00	58.0	38	58	57	+	111	D 81	.55	-1.01								
			09	01	2.6	38	57	55	+	40	D										
0900+395		0	09	00	11.7	39	30	59	+	23	D 62	.15	-1.10								
			09	00	15.1	39	30	11	+	14	D										
0900+428	42.28	4	09	00	58.8	42	50	02	+	1237	R 8	2.65	-.60	G?	0.328	58.74	.6	.7	1.5		8
0902+383		2	09	02	17.7	38	19	26	+	300	R 20	.70	-.66								

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TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0902+384		2	09	02	2.5	38	26	34		149	U	.61	-1.11								
0902+414	41.18	4	09	02	48.0	41	28	31	+	655	R 8	2.03	-.89	N	17.5	47.88	30.4	1.3	.6		5
0902+416		3	09	02	7.2	41	40	39		329	U	.93	-.81	N	19.5	7.48	35.8	-3.1	3.3		
0903+428		3	09	03	9.7	42	51	08	+	244	R 8	1.20	-1.25								
0904+386		2	09	04	34.9	38	39	46	+	150	R 14	.43	-.83	B	17.8	35.15	46.5	-2.9	-.5		
0904+396		1	09	04	25.8	39	36	39	+	13	D 23	.26	-1.20								
			09	04	27.8	39	36	34		43	D										
0904+399		0	09	04	15.5	39	56	26		85	U	.16	-.50	N	17.5	15.42	27.0	.9	-1.0		
0904+417B	41.19	4	09	04	16.0	41	46	36	+	407	D 59	4.09	-.85	G	0.282	18.46	50.6	-.7	-1.6	4	
			09	04	20.8	41	47	02		977	D										
0905+380A	217	4	09	05	41.3	38	00	29	+	1865	R 17	7.09	-1.05								
0905+399	39.24	3	09	05	1.3	39	55	37		181	D106	.94	-1.10								
			09	05	10.2	39	55	11	+	50	D										
0906+383		2	09	06	53.3	38	18	31		112	U	.61	-1.33								
0906+421		3	09	06	30.8	42	09	33	+	225	R 17	1.29	-1.37	G	0.083	31.11	34.0	-3.4	-1.0		
0906+430	216	4	09	06	17.3	43	05	59		3930	U	11.90	-.87	B	* 0.670	17.26	58.0	.4	1.0		38
0907+381		2	09	07	45.0	38	11	32		261	U	.43	-.39	B	17.4	44.92	32.5	.9	-.5		
0908+380B		3	09	08	39.6	38	02	37	+	135	R 12	1.00	-1.57								
0908+380C	38.27	4	09	08	53.3	38	04	00		561	D 19	4.58	-1.20								
			09	08	54.7	38	03	49		431	D										
0909+395		0	09	09	30.0	39	35	07		26	U	.10	-1.05								
0909+432		3	09	09	44.9	43	17	44	+	329	R 7	1.21	-1.02								
0910+392		2	09	10	39.5	39	14	36	+	43	D 48	.41	-1.08	B	19.0	42.31	37.7	-4.2	-2.2	5	
			09	10	43.7	39	14	35	+	60	D										
0911+384		2	09	11	27.9	38	29	10	+	203	R 12	.65	-.91								
0911+395		1	09	11	28.1	39	35	08		117	U	.39	-.94								
0911+418		3	09	11	31.7	41	50	09	+	226	D 66	1.00	-.59	G	0.071	32.12	41.5	.9	-5.5	4	
			09	11	32.7	41	49	03	+	245	D										
0912+388		1	09	12	25.0	38	50	26		118	U	.37	-.90								
0912+392		1	09	12	55.5	39	12	51		53	U	.21	-1.08	B	20.0	55.43	50.9	.8	.1		

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0913+385		2	09	13	39.3	38	30	41	+	344	R 30	.70	-.56	G	* 0.071	39.15	39.2	1.8	1.8		15
0913+391	38.28	4	09	13	39.5	39	07	02		1200	U	1.65	-.25	B	18.5	39.50	1.7	.0	.3		19
0914+390		1	09	14	32.7	39	01	54	+	90	R 65	.33	-1.02								
0917+458A	219	4	09	17	46.6	45	50	48	+	3540	D143	24.30	-1.06	G	* 0.174	51.00	42.0	-6.2	5.5	4	35
0917+458B		4	09	17	54.2	45	52	47	+	2725	D										
0918+381	38.29	4	09	18	37.3	38	06	40		55	T 49	2.45	-1.07	N	18.8	39.27	58.1	-5.5	-1.8	4	
			09	18	38.7	38	06	55		201	T										
			09	18	40.4	38	07	14		371	T										
0918+395		0	09	18	23.9	39	30	26	+	17	D 19	.12	-.90	N	18.7	24.27	29.0	3.8	2.0	4	
			09	18	25.3	39	30	36	+	21	D										
0918+444	44.18	3	09	18	22.1	44	26	53	+	224	D146	1.03	-1.07	G	0.198	17.29	31.5	-15.9	1.5	4	
			09	18	9.5	44	26	13	+	40	D										
0919+381		3	09	19	8.0	38	06	52	+	298	R 13	1.05	-.99								
0920+390		2	09	20	6.3	39	02	32		351	U	.43	-.16								
0920+408		3	09	20	50.6	40	47	45	+	114	T 51	.86	-1.01	G	0.090	50.94	51.6	4.0	.6	5	
			09	20	51.5	40	47	59	+	109	T										
			09	20	55.0	40	47	58		15	T										
0921+400		2	09	21	40.8	40	03	27		143	U	.47	-.93								
0922+397		1	09	22	44.1	39	48	06		35	U	.24	-1.51								
0922+407		2	09	22	50.7	40	42	49		314	U	.48	-.33	B	19.5	50.59	48.6	1.2	.4		
0922+422		3	09	22	47.8	42	16	35	+	210	R 6	.97	-1.20	B	17.9	47.56	37.2	2.7	-2.2		
0922+425	42.29	3	09	22	11.9	42	30	26	+	244	R 14	1.15	-1.22								
0923+392	39.25	4	09	23	55.3	39	15	24		2972	U	3.33	-.09	B	* 0.698	55.29	23.4	.1	.6		12
0923+398		1	09	23	38.3	39	50	47		145	U	.35	-.69	G	0.392	38.48	50.0	-2.1	-3.0		
0926+388		2	09	26	34.5	38	49	12		124	U	.50	-1.09	B?	18.5	34.45	13.6	.6	-1.6		
0926+392		1	09	26	8.4	39	12	23		86	U	.27	-.90								
0929+395		1	09	29	18.2	39	31	52		93	U	.24	-.74								
0930+389		3	09	30	.7	38	55	10		260	U	.92	-.99								
0930+395		1	09	30	34.6	39	35	03		41	D 17	.34	-1.02								
			09	30	35.4	39	35	17		51	D										
0931+399	39.26	4	09	31	59.7	39	55	30	+	876	R 14	3.16	-1.01	G	0.328	59.77	31.2	-.8	-1.2		24

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TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0932+397		1	09	32	5.2	39	46	26		65	U	.28	-.93								
0934+387		2	09	34	48.1	38	45	20		129	D108	.67	-.97								
			09	34	48.2	38	47	08		65	D										
0935+397		1	09	35	33.0	39	47	53		143	U	.32	-.63	N	19.5	33.06	52.5	-.7	.5		
0935+428A	42.30	3	09	35	6.6	42	52	07		387	U	1.35	-.98								
0936+398		1	09	36	17.9	39	53	19		83	U	.33	-1.08								
0936+399		0										.14									
0936+405	40.23	4	09	36	13.4	40	30	13	+	279	D 26	2.12	-1.09	N	18.5	13.85	36.7	1.7	2.3	2	
			09	36	14.0	40	30	39	+	251	D										
0937+391	39.27	4	09	37	57.5	39	07	48	+	219	D 40	1.88	-.86	B	* 0.618	59.10	31.0	1.1	-.6	5	7
			09	38	.1	39	07	21	+	408	D										
0937+396		0	09	37	38.7	39	42	14	+	37	R 7	.15	-1.10								
0938+399A		0										.16									
0938+399B	223.1	4	09	38	17.5	39	57	49	+	934	D 71	4.86	-.88	G	* 0.107	18.23	22.2	.2	1.3	4	31
			09	38	19.0	39	58	58	+	650	D										
0942+394		0	09	42	51.9	39	28	49		48	U	.10	-.57								
0942+399		0	09	42	56.6	39	56	43	+	56	R 12	.19	-.96								
0943+399		3	09	43	13.9	39	58	10	+	230	R 8	.95	-1.11								
0944+390A		1	09	44	7.3	39	05	10	+	60	R 13	.24	-1.09								
0944+390B		1	09	44	23.6	39	00	46		96	U	.27	-.81	G?	0.392	23.61	46.6	-.1	-.6		
0944+397		1	09	44	44.8	39	47	07	+	68	R 10	.25	-1.02								
0945+408	40.24	4	09	45	50.1	40	53	43		1280	R 3	2.45	-.51	B	* 1.252	50.16	44.1	-.7	-1.1		44
0945+419	41.20	4	09	45	23.7	41	55	41	+	580	R 23	2.21	-1.05								
0947+405		2	09	47	23.7	40	31	59		110	U	.44	-1.09								
0947+424	42.31	3	09	47	9.3	42	26	54	+	251	D 55	1.48	-.99								
			09	47	14.2	42	27	07	+	166	D										
0948+390		2	09	48	58.6	39	04	38	+	117	R 13	.40	-.96	G	0.328	58.48	34.6	1.4	3.4		
0950+402		2	09	50	58.1	40	14	07		92	D 26	.55	-1.21								
			09	51	.2	40	13	57	+	25	D										
0951+398		1	09	51	27.0	39	52	38	+	40	D 54	.27	-1.02	*G	0.153	27.44	25.8	-2.5	-12.2	5	
			09	51	27.5	39	51	44	+	33	D										

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
0951+408		3	09	51	38.7	40	50	51		102	D 17	.92	-1.03	B	18.5	39.03	56.5	1.9	1.0	4	
			09	51	39.7	40	51	04		146	D										
0951+422		3	09	51	7.0	42	15	19		424	U	1.08	-.73	G?	0.282	7.10	21.0	-1.1	-2.0		
0953+382		2	09	53	43.4	38	17	19	+	181	R 14	.65	-1.00								
0953+398		2	09	53	5.9	39	49	31		295	U	.50	-.41	N	18.5	5.90	32.1	.0	-1.1		
0954+436		3	09	54	39.2	43	41	08	+	380	R 12	1.11	-.84	G	0.247	39.06	8.8	1.5	-.8		
0955+380		2	09	55	30.4	38	01	37		129	U	.46	-1.00								
0955+387	38.30	3	09	55	1.0	38	44	24	+	150	D 18	1.42	-1.10	N	20.0	1.66	19.9	-.1	-.9	4	1
			09	55	2.3	38	44	14	+	197	D										
0955+390		3	09	55	42.9	39	02	44		392	U	1.05	-.77								
0955+396		2	09	55	55.0	39	36	52		148	U	.49	-.94								
0956+391		1	09	56	41.0	39	12	02		29	D 78	.28	-1.23								
			09	56	43.8	39	10	51		29	D										
0956+404		2	09	56	34.8	40	29	52	+	100	R 8	.47	-1.21								
0956+475	47.31	4	09	56	8.0	47	35	37		450	U	2.10	-1.21								
0957+399		1	09	57	44.4	39	55	12		64	U	.24	-1.04								
0958+390A		1	09	58	23.8	39	02	24	+	115	R 10	.34	-.85								
0958+391		1	09	58	17.7	39	06	14	+	41	R 22	.22	-1.32								
1004+446	44.19	4	10	04	13.5	44	39	38		734	D 35	4.22	-1.06								
			10	04	15.1	44	40	09	+	360	D										
1007+417	41.21	4	10	07	26.1	41	47	16	+	1500	R 25	4.10	-.79	*B	* 0.613	26.13	25.3	-.3	-9.3		45
1007+422		3	10	07	22.8	42	14	19		412	U	1.00	-.70								
1008+395		0	10	08	56.2	39	34	54	+	52	R 12	.17	-.93								
1008+423		3	10	08	53.5	42	19	23		570	U	1.12	-.53								
1008+467A	239	4	10	08	39.3	46	43	09	+	1262	R 11	6.55	-1.29								
1009+389A		1	10	09	26.0	38	58	56		100	U	.29	-.83								
1009+434	43.18	3	10	09	7.0	43	27	55	+	163	R 10	.91	-1.35								
1012+389		1	10	12	59.2	38	56	50		85	U	.33	-1.06	N	20.5	59.50	48.2	-3.5	1.8		
1012+395		0	10	12	52.6	39	35	04	+	70	R 40	.18	-.74								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1012+425	42.32	3	10	12	30.8	42	34	38		116	D 22	.90	-1.06								
			10	12	32.3	42	34	53		116	D										
1013+410	41.22	4	10	13	3.5	41	01	55	+	386	D122	2.86	-1.15	G	0.117	58.78	44.8	3.5	2.8	5	
1012+410B		3	10	12	52.8	41	01	37	+	271	D										
1014+392	39.29	4	10	14	16.5	39	16	23		1253	U	3.08	-.71	G	* 0.206	16.56	20.7	-.7	2.3		33
1014+397A	39.30	3	10	14	15.5	39	47	18	+	279	D128	2.15	-.97	G	* 0.060	19.36	9.5	6.7	17.4	5	34
1014+397B		3	10	14	23.5	39	45	46	+	349	D										
1015+383		2	10	15	27.8	38	20	15	+	103	T 43	.64	-.86	B	17.5	28.92	34.5	-1.4	-2.5	2	
			10	15	28.8	38	20	32	+	50	T										
			10	15	30.5	38	20	44	+	60	T										
1016+388B		2	10	16	48.1	38	48	29		17	T 73	.43	-.96								
			10	16	49.6	38	48	35		25	T										
			10	16	53.9	38	48	56		84	T										
1016+396		0	10	16	29.5	39	37	45		11	U	.14	-1.99								
1016+397		0	10	16	55.0	39	46	07	+	28	D101	.11	-.60	G	0.100	56.98	27.8	3.2	1.3	5	
			10	17	.0	39	44	43	+	23	D										
1016+443		3	10	16	46.5	44	23	30		315	U	1.00	-.91	G?	0.328	46.18	30.4	3.4	-.4		
1018+393		1	10	18	41.1	39	19	03		53	U	.24	-1.18								
1018+405		2	10	18	47.1	40	34	48	+	197	R 20	.67	-.96								
1019+382A		2	10	19	27.2	38	18	02		185	U	.56	-.87								
1019+394	39.31	3	10	19	58.7	39	24	00	+	367	R 8	1.40	-1.05								1
1019+395		1	10	19	18.8	39	32	39		74	U	.28	-1.04								
1019+397		0	10	19	40.4	39	47	00		41	D 19	.17	-.86	B	16.0	40.57	.8	-2.0	-.8	1	
			10	19	42.1	39	47	02		16	D										
1020+400		4	10	20	14.6	40	03	28		1103	U	1.77	-.37	B	16.3	14.56	26.8	.5	1.2		50
1021+384		2	10	21	.9	38	24	00		226	U	.57	-.73								
1022+432	43.19	4	10	22	30.6	43	12	58	+	833	R 15	2.68	-.92								
1023+393		1	10	23	8.6	39	20	39	+	83	R 15	.32	-1.06	*G	0.142	8.05	40.5	6.4	-1.5		
1024+463	46.21	4	10	24	13.1	46	17	35	+	520	D 68	4.09	-1.01								
			10	24	13.4	46	18	43	+	612	D										
1025+390B	39.32	3	10	25	49.4	38	59	57		637	U	1.54	-.69	G	0.180	49.33	58.3	.8	-1.3		50
1025+394		1	10	25	55.9	39	26	08		64	U	.29	-1.18								

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TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1027+383		2	10	27	43.7	38	18	58		55	D 64	.72	-.87								
			10	27	48.8	38	18	34		182	D										
1027+390		1	10	27	31.7	39	02	25	+	76	R 35	.23	-.87								
1027+392		3	10	27	21.7	39	13	18		379	U	.83	-.61								
1028+400		2	10	28	21.8	40	02	21		116	U	.42	-1.01								
1028+402		3	10	28	7.2	40	16	36	+	137	D 55	.82	-1.17								
			10	28	7.6	40	15	43	+	48	D										
1030+398		2	10	30	27.5	39	51	20		397	U	.70	-.44								
1030+415		3	10	30	7.8	41	31	34		587	U	1.02	-.43	B	* 1.120	7.78	34.0	.2	.0		45
1033+387A		3	10	33	29.8	38	45	47	+	57	D160	.33	-.88	G	0.247	29.70	5.2	7.2	-2.5	5	
			10	33	30.9	38	48	27	+	51	D										
1033+387B		3	10	33	52.8	38	42	19	+	105	R 20	.36	-.97								
1033+388		1	10	33	38.0	38	52	17	+	33	T166	.29	-.51	G	0.124	43.11	42.5	.2	-.6	5	
			10	33	42.1	38	51	09	+	30	T										
			10	33	45.4	38	49	57	+	88	T										
1033+399		0										.13									
1033+408		3	10	33	30.8	40	51	11		213	D 22	.95	-.97	N	20.0	31.77	2.0	-5.9	-1.0	4	
			10	33	31.7	40	50	51	+	63	D										
1034+397		0	10	34	39.6	39	43	24	+	15	D 33	.10	-1.22								
			10	34	41.4	39	43	49		6	D										
1034+404	40.26	4	10	34	18.3	40	27	13	+	638	R 10	2.15	-.95								
1035+398		1	10	35	23.0	39	48	37		37	D 14	.21	-.82								
			10	35	23.9	39	48	27		37	D										
1037+399		1	10	37	16.1	39	58	33	+	29	T 94	.28	-1.06								
			10	37	18.7	39	58	08	+	16	T										
			10	37	22.8	39	57	39	+	27	T										
1038+398		1	10	38	36.5	39	49	38	+	103	R 8	.36	-.98								
1039+397		2	10	39	8.7	39	42	41	+	246	R 22	.66	-.77	*G	0.124	7.48	34.0	14.0	7.0		
1039+424		3	10	39	11.8	42	26	13		250	U	.87	-.98								
1040+395		2	10	40	10.5	39	29	39	+	94	D 75	.76	-1.00								
			10	40	15.9	39	30	22	+	119	D										
1040+397		0	10	40	25.3	39	45	22	+	50	R 15	.13	-.75								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1040+398		0	10	40	43.8	39	49	57	+	40	R 40	.11	-.79	G	0.198	43.80	57.5	.0	-.5		
1041+392		2	10	41	34.5	39	16	36	+	189	R 14	.60	-.91								
1042+392	39.33	3	10	42	23.6	39	12	25	+	573	R 10	1.50	-.75								
1042+393		1	10	42	29.7	39	20	06	+	68	R 50	.28	-1.11								
1042+397		0	10	42	48.6	39	45	56		107	U	.17	-.36								
1043+394		1	10	43	5.7	39	29	02		42	D 21	.30	-.93								
			10	43	7.1	39	28	48	+	50	D										
1044+454		4	10	44	38.6	45	24	43		412	U	1.67	-1.10								
1047+387		2	10	47	56.0	38	47	43	+	153	U	.42	-.79								
1047+396		0	10	47	20.8	39	41	43	+	26	D 54	.16	-1.09								
			10	47	25.4	39	41	31	+	14	D										
1049+384		3	10	49	22.3	38	27	39		613	U	1.24	-.55								1
1050+391		1	10	50	44.2	39	09	46	+	129	R 12	.37	-.83								
1052+380		3	10	52	56.1	38	02	02		198	U	.82	-1.11								
1052+389		1	10	52	35.1	38	56	16	+	21	D 21	.24	-1.14								
			10	52	36.9	38	56	22	+	35	D										
1053+384		2	10	53	23.7	38	24	25	+	77	D 40	.75	-.87								
			10	53	24.0	38	25	05	+	171	D										
1053+394		1	10	53	59.0	39	27	35		52	D 22	.28	-.95								
			10	54	.6	39	27	22		31	D										
1054+396		0	10	54	37.8	39	41	18		74	U	.18	-.70								
1055+381		2	10	55	19.5	38	10	06	+	68	D 32	.60	-.76								
			10	55	21.1	38	10	32	+	161	D										
1055+396		0	10	55	45.8	39	36	54		42	D188	.16	-.67								
			10	55	57.6	39	34	48	+	26	D										
1055+404A		3	10	55	11.1	40	26	23		311	U	.94	-.87								
1055+404B	40.27	3	10	55	48.3	40	29	56		237	D 38	1.08	-1.08								
			10	55	49.3	40	30	28		35	D										
1056+387		3	10	56	28.3	38	41	59	+	257	R 20	.86	-.95								
1056+396		3	10	56	22.8	39	41	10	+	242	R 8	.82	-.96								
1056+399		0	10	56	15.0	40	00	07		74	U	.14	-.50								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1056+432A	247	4	10	56	08.2	43	17	29	+	2700	R 14	7.01	-.75	B	18.9						
1058+393		1	10	58	42.2	39	20	40		245	U	.34	-.26								
1100+398		0	11	00	36.2	39	49	56		35	U	.14	-1.09								
1101+384		3	11	01	40.5	38	28	43		576	U	1.14	-.54	G	* 0.030	40.49	41.9	.1	1.1		9
1101+395		0	11	01	50.5	39	31	02		53	U	.19	-1.00								
1101+396		0	11	01	15.4	39	37	06	+	20	R 10	.13	-1.46								
1103+393		2	11	03	21.4	39	22	17	+	173	R 15	.52	-.86								
1104+390		1	11	04	54.0	39	05	18	+	70	R 48	.20	-.82								
1104+397		0	11	04	10.9	39	45	26	+	25	D 77	.12	-.96								
			11	04	17.6	39	45	28	+	10	D										
1105+390		1	11	05	36.4	39	02	07	+	119	R 46	.23	-.52	G	0.247	36.81	9.0	-4.8	-2.0		
1105+392	39.34	4	11	05	49.8	39	15	16		669	D 56	2.31	-.87	N	18.5	51.48	57.4	4.3	-1.9	4	23
			11	05	53.9	39	14	35	+	90	D										
1106+380		3	11	06	43.4	38	00	48		1121	U	1.29	-.11								
1107+379	37.29	4	11	07	2.6	37	54	35		975	D 50	5.63	-.79	G	* 0.346	4.53	45.9	-1.0	.7	5	34
			11	07	6.1	37	54	57		1090	D										
1107+389		1										.24									
1108+394		0	11	08	28.7	39	28	51	+	30	R 45	.10	-.94								
1108+399		3	11	08	33.7	39	56	32	+	329	R 12	.93	-.81	N	20.0	33.54	32.3	1.8	-.3		29
1108+411B	41.23	4	11	08	52.8	41	06	35	+	771	R 22	3.20	-1.12	*G	* 0.074	53.45	42.2	-7.3	-7.2		43
1109+437	43.21	4	11	09	51.2	43	42	47		554	D 53	4.77	-1.09	*B	* 1.680	52.28	6.1	5.6	-10.1	2	51
			11	09	52.8	43	41	56		634	D										
1110+391		1	11	10	41.9	39	09	59		55	D 28	.32	-1.12								
			11	10	41.3	39	09	31		22	D										
1111+391		1	11	11	44.6	39	06	56		139	U	.37	-.77								
1111+396A		0	11	11	53.0	39	43	35	+	27	R 13	.19	-1.53								
1111+396B		0										.17									
1111+398		0										.18									
1111+408	254	4	11	11	53.0	40	53	41	+	2487	R 13	10.92	-1.16	B	* 0.734	53.28	41.9	-3.2	-.9		12
1112+435		3	11	12	34.4	43	31	05	+	246	R 15	.83	-.95	G?	0.247	34.61	.5	-2.3	4.5		

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1115+380A		2	11	15	21.0	38	04	31		202	U	.70	-.97								
1115+399		0	11	15	12.8	39	57	08	+	53	R 8	.15	-.81								
1116+388		1	11	16	11.4	38	50	50		85	U	.20	-.67								
1116+392		1	11	16	19.0	39	14	36	+	70	T 83	.36	-.91	B	18.0	19.50	17.5	.0	-.5	2	
			11	16	19.5	39	15	17		12	T										
			11	16	20.5	39	15	58	+	31	T										
1117+441		3	11	17	29.7	44	11	05	+	124	D 32	1.01	-.92								
			11	17	31.7	44	11	29	+	187	D										
1118+390		1	11	18	28.2	39	01	17	+	23	D 89	.20	-1.03								
			11	18	31.0	38	59	54	+	31	D										
1119+398		1										.22									
1121+399		2	11	21	4.3	39	54	18		268	U	.46	-.42								
1121+435		3	11	21	48.2	43	32	09		408	U	1.20	-.85								
1121+444		3	11	21	10.4	44	24	49	+	127	D 29	1.39	-1.11								
			11	21	10.7	44	25	18		209	D										
1122+390		1	11	22	1.2	39	02	16	+	89	R 36	.22	-.71	G	* 0.007	1.09	13.0	1.3	3.0		9
1122+397		1	11	22	58.9	39	42	11		84	D 32	.29	-.80								
			11	23	1.2	39	41	52		21	D										
1123+395		1	11	23	46.2	39	35	13		127	U	.36	-.82	B	16.4	45.85	15.1	4.0	-2.0		
1127+380		2	11	27	14.6	38	04	37		227	U	.78	-.97								
1128+385		2	11	28	12.5	38	31	51		794	U	.52	.33	B	18.6	12.55	52.3	-.6	-1.3		10
1128+392		2	11	28	31.4	39	17	07	+	34	D 8	.56	-1.21								
			11	28	32.1	39	17	07	+	86	D										
1128+396		1	11	28	48.5	39	37	54	+	32	D 25	.25	-.92								
			11	28	48.0	39	38	19	+	45	D										
1128+436		3	11	28	4.9	43	41	10		271	D 56	1.55	-1.20								
			11	28	5.9	43	42	05		64	D										
1128+455		4	11	28	56.4	45	31	23		1886	U	4.37	-.66	N	18.5	56.41	25.3	-.1	-2.3		
1130+387		2	11	30	17.0	38	43	31		171	U	.48	-.81								
1131+388		2	11	31	22.1	38	52	24	+	155	D 36	.75	-.75	G	0.328	23.78	27.0	-1.5	-3.0	4	
			11	31	25.2	38	52	24	+	133	D										
1131+391		1	11	31	33.1	39	07	52		132	R 21	.31	-.67								

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TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1131+437	43.22	4	11	31	57.3	43	44	36	+	1408	R 8	3.90	-.80								
1132+396		0	11	32	22.0	39	39	37	+	33	R 7	.11	-.94								
1132+406		2	11	32	.3	40	38	04	+	95	D 25	.77	-.95	G	0.392	.61	47.7	-.2	1.6	5	
			11	32	.8	40	37	39	+	135	D										
1132+410	40.29	3	11	32	5.3	41	00	28	+	263	R 8	1.10	-1.12								
1133+395		0	11	33	46.3	39	33	51	+	53	R 10	.14	-.76								
1133+432		3	11	33	15.3	43	15	20		1319	U	.88	.32								
1134+406		2	11	34	47.1	40	39	07		176	U	.53	-.86								
1135+390		1	11	35	18.8	39	01	45	+	44	R 8	.22	-1.26								
1135+401		2	11	35	1.1	40	07	25		370	U	.66	-.45								
1136+383		3	11	36	55.2	38	20	19	+	401	U	.84	-.58								
1136+390		2	11	36	30.5	39	03	55	+	187	R 8	.49	-.76								
1136+420	42.33	3	11	36	19.5	42	05	17		360	U	1.09	-.87								
1137+396		0	11	37	45.1	39	38	33	+	22	D 27	.11	-.72	N	17.0	46.22	40.7	-.2	-2.2	4	
			11	37	47.3	39	38	44	+	22	D										
1140+394		1	11	40	35.9	39	26	02	+	40	D 26	.32	-.98								
			11	40	36.9	39	25	38	+	52	D										
1140+399		0	11	40	40.7	39	55	19		19	D 30	.14	-1.02								
			11	40	42.0	39	54	52		19	D										
1141+374	37.32	4	11	41	41.8	37	23	43	+	662	D265	4.75	-.95	G	0.110	49.60	15.6	-1.8	3.9	4	15
			11	41	57.1	37	26	56	+	746	D										
1141+392		2	11	41	9.3	39	15	06	+	89	D 47	.60	-1.04								
			11	41	12.4	39	15	37	+	71	D										
1141+400		2	11	41	55.1	40	00	02		122	U	.41	-.95	N	18.5	55.26	2.2	-1.8	-.2		
1141+466	46.23	4	11	41	.4	46	37	59	+	814	R 12	2.60	-.91	G	0.057	.29	59.4	1.1	-.4		8
1142+392		2	11	42	56.5	39	13	36	+	156	R 8	.57	-1.02	*B	18.0	56.29	26.0	2.4	10.0		
1143+405		2	11	43	50.0	40	31	38		106	U	.49	-1.20								
1143+456	45.22	4	11	43	37.0	45	37	17		633	U	2.26	-1.00								
1144+398		0	11	44	40.0	39	53	32	+	46	R 12	.12	-.75								
1144+402		3	11	44	21.0	40	15	13		1001	U	.93	.06	B	18.0	20.98	13.9	.2	-.9		

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1144+404		2	11	44	5.5	40	24	41		169	U	.51	-.87								
1148+387	38.31	4	11	48	53.3	38	42	33	+	555	R 6	1.83	-.94	B	* 1.303	53.32	34.4	-.2	-1.4		47
1148+477	47.33	4	11	48	32.3	47	45	36	+	503	R 13	2.29	-1.19	B	* 0.867	32.28	36.1	.2	-.1		53
1149+390		1	11	49	1.8	39	02	55	+	54	R 8	.33	-1.42								
1149+398		1	11	49	4.7	39	51	25	+	40	D107	.21	-.86								
			11	49	12.5	39	50	26	+	30	D										
1150+388		2	11	50	9.4	38	48	26		246	U	.66	-.77								
1150+401		2	11	50	24.0	40	10	39	+	68	D 33	.55	-1.13								
			11	50	26.8	40	10	28	+	62	D										
1150+438		3	11	50	6.2	43	53	00		284	U	.96	-.96								
1151+383		3	11	51	25.9	38	21	49	+	493	R 15	1.42	-.83	G?	0.392	25.75	47.3	1.8	1.7		
1151+384B	38.32	4	11	51	14.6	38	28	24	+	494	D 63	2.47	-.70	G	0.142	17.41	30.4	-.5	-2.8	5	18
			11	51	20.0	38	28	31	+	520	D										
1151+395		0										.12									
1151+408		3	11	51	19.0	40	53	34		653	U	.88	-.23								
1151+456	45.23	4	11	51	44.8	45	40	50	+	270	D 62	2.28	-.93	G?	* 0.192	44.84	11.5	.2	.5	5	45
			11	51	44.9	45	39	48	+	427	D										
1153+407B		3	11	53	17.6	40	47	26	+	200	R 11	.90	-1.18								
1153+451	45.24	4	11	53	32.2	45	06	55		243	D 30	1.79	-.80								
			11	53	35.0	45	07	00		405	D										
1154+397		0	11	54	12.7	39	45	00	+	75	R107	.14	-.49	G	0.142	12.98	5.9	-3.2	-5.9		
1154+398		0	11	54	54.8	39	52	48		40	U	.15	-1.03								
1156+389		1	11	56	29.7	38	57	01	+	43	D 22	.31	-.87								
			11	56	30.7	38	56	42	+	59	D										
1157+396		0	11	57	17.3	39	41	00	+	31	U	.16	-1.28								
1157+460	46.25	4	11	57	57.2	46	05	24		1042	U	2.88	-.80								
1158+393		1	11	58	21.8	39	19	28		60	D 21	.30	-1.14								
			11	58	21.8	39	19	50		10	D										
1159+395		3	11	59	16.3	39	35	52		548	U	.83	-.33								
1200+393		1	12	00	26.9	39	22	28	+	88	R 15	.25	-.82								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1201+394		3	12	01	33.7	39	29	00		430	U	1.11	-.74	N	19.0	33.69	.1	.1	-.1		1
1201+396		0	12	01	9.4	39	39	37	+	52	R 9	.16	-.88								
1202+388		1	12	02	58.4	38	50	45	+	97	R 14	.30	-.88								
1202+397		0	12	02	15.9	39	46	29		54	U	.11	-.56								
1203+384		2	12	03	43.7	38	29	02	+	97	D 32	.67	-.98	B	18.1	44.74	17.3	-2.8	-2.3	4	
			12	03	45.3	38	29	28	+	95	D										
1204+399		1	12	04	4.6	39	57	45		206	U	.34	-.39	B	18.2	4.55	44.8	.6	.2		
1204+401	40.30	3	12	04	33.9	40	11	21		230	U	1.00	-1.15								
1205+390		1	12	05	25.2	39	05	31	+	40	U	.21	-1.30								
1205+392	39.35	4	12	05	19.1	39	12	40	+	210	D 24	1.83	-1.33	G	* 0.243	19.53	41.5	4.1	-1.5	5	34
			12	05	21.2	39	12	40	+	126	D										
1206+399a		0	12	06	31.6	39	56	57	+	44	R 30	.10	-.64	G	0.142	31.36	55.0	2.7	2.0		
1206+399b		0	12	06	36.7	39	55	24	+	42	R 30	.10	-.68	G	0.142	36.50	24.0	2.3	.0		
1206+439B	268.4	4	12	06	42.0	43	56	00	+	1926	R 10	5.69	-.85	B	* 1.400	42.06	1.9	-.6	-1.9		12
1207+386		3	12	07	36.7	38	37	54	+	145	D 21	.94	-.90								1
			12	07	38.5	38	37	56	+	153	D										
1208+396		2	12	08	1.0	39	41	02		304	U	.76	-.72	G	* 0.003	1.03	1.6	-.3	.4		11
1209+396		2	12	09	56.7	39	39	41	+	260	R 22	.72	-.80								
1209+399		0	12	09	24.3	39	59	30		59	U	.19	-.92								
1211+388		1	12	11	59.5	38	51	49	+	95	R 10	.34	-1.00								
1212+380		3	12	12	26.1	38	05	31		278	U	1.08	-1.06								
1212+389		1	12	12	43.2	38	56	03		83	U	.30	-1.01	B	21.0	43.32	3.5	-1.4	-.5		
1213+389		1	12	13	53.2	38	54	57	+	79	R 18	.27	-.96								
1216+400		2	12	16	34.8	40	04	30		119	U	.49	-1.11								
1216+402		3	12	16	7.5	40	17	25		312	U	1.00	-.91								
1217+427		3	12	17	25.1	42	46	29		240	U	.87	-1.01								
1218+395		1	12	18	41.0	39	35	35	+	39	D 40	.33	-1.07								
			12	18	44.0	39	35	14	+	45	D										
1218+398		0	12	18	24.2	39	53	44	+	13	D 18	.11	-1.13								
			12	18	25.4	39	53	56	+	13	D										

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1218+421		3	12	18	44.1	42	08	56	+	350	R 43	1.08	-.88	*G	0.071	44.58	52.5	-5.3	3.5		
1219+382		2	12	19	42.4	38	15	52	+	119	R 11	.58	-1.24								
1220+393		1	12	20	52.6	39	21	02		51	D 33	.29	-1.04								
			12	20	55.5	39	20	59	+	26	D										
1220+408	40.31	3	12	20	6.8	40	52	59		448	U	1.50	-.95								
1221+394		1	12	21	41.7	39	25	27		81	U	.28	-.97								
1221+397		0	12	21	39.6	39	46	33		52	U	.12	-.65								
1221+398		0	12	21	18.8	39	49	06	+	47	D 33	.13	-.31	G	0.247	19.41	23.5	-.1	-2.0	4	
			12	21	20.0	39	49	37	+	40	D										
1222+390		1	12	22	53.2	39	05	06	+	62	R 8	.21	-.96								
1222+398	39.36	0	12	22	4.0	39	49	21	+	30	D 54	.16	-.99								
			12	22	2.7	39	48	28	+	15	D										
1222+423	272	4	12	21	59.6	42	22	45	+	506	D 52	4.02	-.99								
			12	22	1.6	42	23	33		636	D										
1223+395		2	12	23	22.6	39	30	59		476	U	.51	-.05								
1225+403		2	12	25	45.9	40	20	48		105	U	.47	-1.17								
1225+442		3	12	25	15.7	44	17	16		339	U	.93	-.79	G	0.220	15.72	16.8	-.2	-.8		8
1226+395		0	12	26	44.9	39	32	26		31	U	.10	-.92								
1228+397		2	12	28	24.7	39	46	12	+	35	D 52	.60	-.89	B	17.7	25.90	33.8	2.9	-1.8	4	
			12	28	27.6	39	46	52		157	D										
1228+419A		3	12	28	8.5	41	55	28	+	801	R120	1.77	-.62	*G	* 0.0016	9.00	13.0	-5.6	15.0		
1229+397		2	12	29	32.5	39	47	13	+	124	R 8	.40	-.92								
1229+405		3	12	29	14.2	40	34	05	+	261	R 10	.91	-.98	B	19.0	13.84	2.4	4.1	2.6		
1230+398		0	12	30	16.0	39	53	06		26	D 51	.14	-1.00								
			12	30	18.7	39	53	45		13	D										
1231+394		2	12	31	57.3	39	25	21	+	121	R 8	.47	-1.06								
1231+432	43.24	3	12	31	53.7	43	13	44	+	273	R 22	.82	-.86								
1232+394		1	12	32	45.7	39	27	30		79	U	.30	-1.05								
1232+397A		2	12	32	3.9	39	47	06	+	259	R 20	.46	-.45								
1232+397B	39.37	3	12	32	39.0	39	42	10		218	U	1.13	-1.29								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1232+399		0	12	32	24.8	39	55	15		60	U	.13	-.61								
1232+414A	41.24	4	12	32	2.2	41	26	32	+	278	D 82	1.81	-.95	G	0.165	5.00	6.0	.0	-1.0	4	29
			12	32	7.8	41	25	38	+	260	D										
1233+399		0	12	33	50.2	39	56	33		25	U	.11	-1.16								
1233+418		3	12	33	10.8	41	53	38		641	U	1.47	-.65	G	0.247	10.85	37.3	-.6	.7		
1234+396		2	12	34	26.2	39	36	58		284	U	.54	-.50								
1236+444A		3	12	36	10.5	44	30	12	+	193	R 20	.45	-.66								
1236+444B		3	12	36	15.3	44	27	27	+	48	D237	.41	-.84	G	0.180	23.24	24.9	9.7	-9.4	4	
			12	36	33.0	44	25	04	+	93	D										
1239+382		2	12	39	46.4	38	15	21	+	50	D 34	.44	-1.34	G	0.153	47.74	22.0	1.3	.5	4	
			12	39	49.3	38	15	24	+	30	D										
1239+390		2	12	39	8.8	39	04	57		107	D 18	.75	-1.03								
			12	39	9.3	39	05	15		95	D										
1239+396		1	12	39	23.8	39	37	14	+	118	R 14	.31	-.76	G	0.198	23.56	13.3	2.8	.7		
1239+442B	44.20	3	12	39	56.1	44	12	19	+	307	D 44	1.21	-.92	B	18.0	57.00	34.3	2.7	3.2	4	
			12	39	58.4	44	12	56	+	68	D										
1240+381		2	12	40	27.0	38	07	25		420	U	.41	.02	B	* 1.316	27.03	25.1	-.4	-.1		20
1240+395		1	12	40	29.0	39	32	11	+	60	R 8	.36	-1.40								
1241+411		3	12	41	57.1	41	08	00		933	U	.80	.12	G?	0.165	57.09	1.4	.1	-1.4		
1242+391		1	12	42	6.3	39	10	43		34	D 43	.25	-1.29								
			12	42	7.1	39	10	00	+	14	D										
1242+410		4	12	42	26.4	41	04	29		1457	U	2.01	-.25	B?	19.5	26.37	30.2	.3	-1.2		8
1244+389	38.33	4	12	44	22.0	38	58	06		271	D 22	2.17	-1.09								
			12	44	23.8	38	57	58		271	D										
1244+397		0	12	44	57.0	39	44	10		41	U	.16	-1.07								
1245+389		1	12	45	57.6	38	58	36	+	130	R 19	.38	-.84								
1245+396		1	12	45	40.9	39	38	08		38	D 36	.33	-.98	B?	19.2	41.02	30.2	.0	-.8	5	
			12	45	41.1	39	38	44		56	D										
1245+399		1	12	45	26.7	39	56	36	+	72	D 19	.30	-.85								
			12	45	27.1	39	56	55	+	30	D										
1246+385C		2	12	46	50.8	38	33	45		85	D 57	.70	-.91								
1246+385B			12	46	50.8	38	32	48	+	134	D										

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1247+450A	45.26	4	12	47	2.7	45	01	02	+	464	D 18	1.68	-.73	B	* 0.799	3.61	10.9	-3.3	-1.9	4	53
			12	47	3.9	45	01	16	+	197	D										
1249+393		1	12	49	23.7	39	19	28		108	U	.26	-.69								
1249+432	43.25	3	12	49	41.0	43	14	00		437	U	1.40	-.91								
1249+475	47.35	4	12	49	56.8	47	31	48		342	D 42	3.29	-1.14								
			12	50	.8	47	32	02		428	D										
1250+384		2	12	50	9.6	38	27	00		183	U	.61	-.94	G	0.392	9.52	59.7	.9	.3		
1250+390		2	12	50	27.8	39	05	53		119	U	.46	-1.06								
1251+398		0	12	51	49.1	39	49	24		35	D 29	.15	-.93	B	19.2	49.36	38.5	.5	-.1	4	
			12	51	49.7	39	49	53		11	D										
1253+374	37.35	4	12	53	55.3	37	29	37		306	D 38	2.27	-1.03								
			12	53	56.6	37	30	12		306	D										
1253+432	43.26	3	12	53	24.5	43	14	39		410	U	1.42	-.97								
1254+476	280	4	12	54	40.9	47	36	32	+	4051	R 12	13.71	-.96								39
1255+448	44.22	4	12	55	45.1	44	51	27	+	540	D 28	2.90	-.98								
			12	55	42.7	44	51	39	+	290	D										
1256+392		2	12	56	42.2	39	16	30	+	215	R 15	.80	-1.03	*B	17.6	41.87	22.9	3.8	7.1		
1257+383	38.34	4	12	57	52.3	38	20	29		337	D 32	2.07	-.98								
			12	57	54.4	38	20	50		253	D										
1257+399		0	12	57	25.5	39	57	01	+	16	D 30	.11	-.97								
			12	57	25.9	39	56	31	+	16	D										
1258+395		0	12	58	22.5	39	35	00	+	31	D 44	.18	-.87	N	18.5	23.56	10.1	4.7	2.0	5	
			12	58	25.6	39	35	25	+	28	D										
1258+404	280.1	4	12	58	13.1	40	25	19		772	D 17	4.51	-1.04	B	* 1.656	13.99	15.9	-2.2	.1	4	12
			12	58	14.5	40	25	13		424	D										
1259+395		0	12	59	41.2	39	31	29	+	26	R 40	.10	-1.05	*B	19.5	41.76	32.0	-6.5	-3.0		
1300+397		1	13	00	29.1	39	46	00	+	87	R 10	.27	-.89	*B	17.8	29.06	7.6	.5	-7.6		
1301+382	38.35	4	13	01	23.7	38	12	06	+	224	D 22	1.65	-.99	G	* 0.126	24.51	17.9	-2.2	-3.0	5	34
			13	01	24.9	38	12	23	+	245	D										
1301+393		1	13	01	.9	39	19	34	+	28	D 26	.32	-1.00	N	19.5	1.09	21.6	4.7	1.4	4	
			13	01	2.1	39	19	12	+	61	D										
1302+388	38.36	4	13	02	38.8	38	48	32	+	437	R 8	1.97	-1.18								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1305+393		1	13	05	19.0	39	20	43	+	139	R 12	.35	-.72								
1306+396		1	13	06	9.6	39	42	00		64	D 20	.35	-1.18								
			13	06	11.4	39	42	03		14	D										
1308+392		1	13	08	20.6	39	12	57		159	R 15	.36	-.64								
1309+412A		3	13	09	25.8	41	11	58		163	D351	.83	-.69	G	0.087	26.92	54.7	2.0	-1.7	4	
			13	09	28.4	41	17	48	+	180	D										
1311+419		3	13	11	48.5	41	56	13	+	84	D 35	1.37	-1.40								
			13	11	50.5	41	56	41	+	147	D										
1312+393		2	13	12	49.5	39	19	34	+	212	R 8	.71	-.95	B	20.0	49.40	33.2	1.2	.8		
1313+387		3	13	13	5.9	38	45	20	+	90	D125	1.05	-1.22								
			13	13	8.9	38	47	20	+	131	D										
1313+392		1	13	13	25.2	39	14	22		139	U	.28	-.55	G	0.392	25.18	24.3	.2	-2.3		
1314+453A	45.27	4	13	13	59.8	45	20	18		582	U	1.75	-.86								
1315+395		2	13	15	34.5	39	31	05	+	160	R 16	.69	-1.15								
1315+396	39.38	3	13	15	2.8	39	41	14		618	U	1.16	-.49	B	17.3	2.86	15.4	-.7	-1.4		48
1317+380		2	13	17	36.2	38	03	07	+	249	R 11	.75	-.86	B	18.6	36.12	9.4	.9	-2.4		
1317+389		1	13	17	45.0	38	56	05		253	U	.31	-.16	B	19.7	44.67	1.7	3.8	3.3		
1317+393		1	13	17	28.0	39	18	35	+	99	R 22	.28	-.81								
1318+398		1	13	18	12.6	39	48	36		61	U	.20	-.93								
1319+388		1	13	19	4.7	38	51	13	+	84	R 5	.29	-.97								
1319+397		0	13	19	57.6	39	46	29		47	U	.16	-.96								
1319+398		1	13	19	33.9	39	53	55		111	U	.31	-.80								
1318+428C	285	3	13	19	.0	42	50	35	+	821	D121	4.87	-.76	G	* 0.079	5.16	55.5	2.6	-8.0	4	32
1319+428A		4	13	19	10.8	42	51	00	+	1025	D										
1320+389		1	13	20	5.8	38	59	51		29	D 19	.20	-1.03								
			13	20	7.0	38	59	37	+	25	D										
1321+415	41.25	4	13	21	10.4	41	30	55	+	290	D 14	1.72	-.87	G	0.392	11.10	52.1	-1.1	.4	4	8
			13	21	11.6	41	30	50	+	280	D										
1322+398		0	13	22	30.9	39	49	19	+	52	R 14	.19	-1.01								
1324+390		1	13	24	47.4	39	04	50		45	D 66	.24	-.97								
			13	24	48.3	39	05	56	+	25	D										

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1324+431	43.28	3	13	24	50.2	43	10	07	+	115	D 19	1.08	-1.00								
			13	24	51.9	43	10	00		187	D										
1327+390		1	13	27	51.5	39	00	55		109	U	.38	-.98								
1327+391		1	13	27	15.2	39	09	43		53	U	.23	-1.15								
1327+398		0	13	27	51.6	39	52	41	+	14	D 58	.12	-1.20								
			13	27	56.6	39	52	52	+	12	D										
1327+474C		4	13	27	46.3	47	27	09	+	1300	R400	3.07	-.67	G	* 0.0017	45.98	21.5	3.2	-12.5		
1328+388		2	13	28	58.2	38	50	26		212	U	.42	-.54								
1328+396		1	13	28	45.7	39	36	43		161	U	.34	-.59								
1330+380		2	13	30	32.1	38	01	12	+	71	D124	.47	-.86								
			13	30	40.8	38	00	01	+	86	D										
1330+389		1	13	30	14.1	38	55	33	+	62	R 18	.20	-.92								
1330+406		2	13	30	43.4	40	39	51		61	D 17	.69	-.91								
			13	30	44.7	40	39	42		154	D										
1331+381		2	13	31	15.7	38	11	21	+	147	R 26	.49	-.94								
1332+385		2	13	32	12.2	38	33	02	+	46	D 31	.51	-.79	G	0.110	12.42	19.3	.9	-1.8	4	
			13	32	12.8	38	33	33	+	139	D										
1333+392		1	13	33	38.7	39	15	13		110	U	.37	-.95								
1333+412	41.26	4	13	33	9.2	41	15	27		312	D 12	2.98	-1.13	G	* 0.187	9.61	22.4	-.9	-1.0	5	8
			13	33	9.8	41	15	17		390	D										
1334+417		3	13	34	16.6	41	46	29		347	U	.89	-.74								
1335+391		1										.24									
1336+391A	288	4	13	36	38.3	39	06	24	+	2820	R 13	10.00	-.99	G	* 0.246	38.47	21.5	-2.0	2.5		37
1336+393		2	13	36	4.0	39	21	10		95	U	.40	-1.13								
1336+396C		3	13	36	57.6	39	41	50	+	218	D 31	1.18	-.95								
			13	36	58.0	39	41	19		134	D										
1336+397A		0	13	36	5.3	39	43	45		64	U	.16	-.72								
1336+397B		1	13	36	44.2	39	45	07		22	U	.22	-1.80								
1336+395		1										.26									
1336+396A		0										.14									
1336+396B		0										.18									

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1336+401		2										.41									
1337+385		2	13	37	13.3	38	30	20		94	U	.44	-1.21								
1337+396		0										.19									
1338+391		1										.20									
1338+394		0	13	38	57.8	39	29	59	+	76	R 20	.18	-.68	B	19.0	57.56	58.0	2.8	1.0		
1339+438	43.29	3	13	39	45.9	43	50	30	+	137	D 20	1.10	-.97								
			13	39	47.7	43	50	23		181	D										
1339+472	47.38	4	13	39	41.3	47	12	29		301	D 14	2.21	-1.02	B	19.7	41.73	25.0	.7	-1.5	4	8
			13	39	42.3	47	12	18		301	D										
1340+439		3	13	40	59.9	43	58	27		462	U	1.20	-.75								
1341+392	39.40	3	13	41	10.1	39	13	39	+	335	R 23	1.22	-1.01	N	20.5	10.10	36.0	.0	3.0		
1342+389A		3	13	42	14.8	38	56	31	+	213	R 15	.86	-1.09	B	17.5	15.11	31.3	-3.6	-.3		
1342+389B		2	13	42	37.9	38	59	37	+	188	R 19	.46	-.70	*G	0.220	37.29	35.0	7.1	2.0		
1343+386	38.37	3	13	43	26.6	38	38	12		829	U	1.52	-.48	B	* 1.844	26.76	12.1	-1.9	-.1		36
1343+430	43.30	4	13	43	26.7	43	05	16	+	1042	R 11	2.72	-.75								
1344+397		0	13	44	55.2	39	43	55		86	U	.16	-.49								
1345+398		0	13	45	4.5	39	50	54	+	55	R 8	.18	-.93								
1346+392		1	13	46	47.6	39	15	00		32	D 72	.22	-.91								
			13	46	47.9	39	16	12		37	D										
1347+391		2	13	47	16.6	39	06	34		154	U	.44	-.82								
1347+396		1	13	47	.5	39	37	59		163	U	.38	-.66								
1347+398		1	13	47	6.7	39	51	45		54	D 21	.39	-.98								
			13	47	8.2	39	51	33	+	58	D										
1347+403		2	13	47	44.2	40	21	12		109	U	.43	-1.08								
1348+392		2	13	48	23.6	39	14	12		155	U	.50	-.92	B	19.0	23.49	13.3	1.3	-1.3		
1348+396		1	13	48	30.4	39	36	35	+	40	D 45	.22	-.65								
			13	48	32.4	39	35	56		56	D										
1349+394		0	13	49	39.6	39	26	43		21	D222	.10	-.78								
			13	49	42.9	39	30	20		16	D										
1349+388		1	13	49	10.9	38	53	02		233	U	.22	.05	B	20.5	10.63	3.4	3.1	-1.4		10

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1350+395		1	13	50	21.3	39	33	22	+	75	R 13	.29	-1.06								
1350+432	43.31	3	13	50	24.0	43	14	09		153	U	.83	-1.33								
1352+383		2	13	52	53.8	38	19	23		45	D 29	.44	-1.09								
			13	52	55.2	38	19	47		64	D										
1352+397		1	13	52	.2	39	42	02		66	D108	.32	-1.01								
			13	52	7.2	39	43	14	+	22	D										
1352+403		2	13	52	.1	40	20	51	+	43	D 24	.47	-1.10								
			13	52	.5	40	20	27	+	72	D										
1353+380		2	13	53	43.8	38	00	06		68	D 27	.49	-.90								
			13	53	46.1	38	00	00		87	D										
1353+397		2	13	53	52.7	39	47	50		127	D 87	.53	-1.00								
			13	54	.0	39	48	12		21	D										
1354+397	39.41	3	13	54	8.0	39	44	22		138	D 42	.96	-1.40								
			13	54	9.5	39	43	43		23	D										
1355+380		2	13	55	30.6	38	04	03	+	72	D 34	.60	-.89	B	20.0	30.54	4.8	.7	-1.9	1	
			13	55	32.3	38	04	31	+	121	D										
1356+393		2	13	56	40.8	39	18	35		248	U	.74	-.86								
1356+397		0	13	56	7.1	39	47	39	+	19	T 53	.12	-.73	*B	20.2	8.84	16.8	-3.9	1.2	2	
			13	56	8.5	39	47	18		6	T										
			13	56	10.8	39	47	07	+	22	T										
1357+392		1	13	57	5.3	39	15	05	+	53	R 37	.20	-1.04	B?	21.2	4.97	5.6	3.8	-.6		
1357+394A		1	13	57	48.8	39	27	34		185	U	.21	-.10								
1357+394B		2	13	57	56.0	39	25	14		47	D 20	.41	-1.08	B	20.5	56.50	25.7	-.1	-1.9	5	
			13	57	56.9	39	25	32		56	D										
1358+433	43.32	3	13	58	28.2	43	18	32		207	D 30	1.37	-.94								
			13	58	31.0	43	18	32		207	D										
1359+419		3	13	59	30.0	41	56	42	+	315	R 11	1.16	-1.02								
1401+387	38.38	4	14	01	4.6	38	42	50	+	290	D 72	1.64	-.85	B	15.8	4.55	20.8	5.6	-5.0	5	28
			14	01	5.5	38	41	38	+	262	D										
1401+395		0	14	01	59.3	39	35	30	+	52	R 7	.17	-.93								
1402+382		2	14	02	20.4	38	15	01	+	58	D 29	.64	-1.04	B?	16.5	21.90	34.3	-1.2	2.7	2	
			14	02	21.8	38	14	37		112	D										
1403+395		1	14	03	6.6	39	31	41	+	61	R 18	.24	-1.07								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1406+397		1	14	06	20.4	39	45	42		102	U	.32	-.90								
1407+388		1	14	07	21.1	38	52	12		51	D 14	.26	-.73	G?	0.392	21.30	14.5	-2.3	-2.5	1	
			14	07	22.3	38	52	08		51	D										
1408+398		2	14	08	37.3	39	48	42	+	137	R 26	.42	-.88								
1408+399		0	14	08	13.7	39	59	32	+	44	R 32	.17	-1.06								
1409+390		2	14	09	37.7	39	00	23		179	U	.46	-.74								
1409+394		1	14	09	22.7	39	28	21	+	118	R 29	.29	-.70								
1410+438	43.33	3	14	10	16.4	43	52	25	+	208	D 73	1.34	-1.08								
			14	10	20.2	43	51	24		130	D										
1411+391A		1	14	11	48.8	39	07	26		81	U	.22	-.78								
1411+391B		1	14	11	55.0	39	09	39		59	D 38	.39	-.84								
			14	11	57.2	39	09	10		74	D										
1411+397		0	14	11	9.3	39	44	56	+	61	R 19	.13	-.59	G?	0.328	9.43	57.3	-1.5	-1.4		
1411+427		3	14	11	39.4	42	43	14	+	127	D 32	.96	-1.08								
			14	11	42.0	42	43	30	+	114	D										
1412+392		2	14	12	59.5	39	13	11		199	U	.47	-.67								
1412+397		0	14	12	45.0	39	43	56		20	D 39	.17	-1.22								
			14	12	48.3	39	44	06		16	D										
1413+398		0	14	13	12.8	39	51	37	+	30	R 10	.13	-1.15								
1414+398		1	14	14	30.8	39	51	35	+	46	D139	.24	-1.07	*G	0.031	29.20	12.0	.0	5.0	2	
			14	14	29.2	39	49	17	+	15	D										
1415+391		1	14	15	49.9	39	12	11	+	11	D 47	.21	-.97								
			14	15	52.4	39	11	33		50	D										
1416+400		2	14	16	56.6	40	00	37		98	D 24	.67	-.96	B	20.0	56.66	24.8	-.1	.2	4	
			14	16	56.7	40	00	13		98	D										
1417+383		2	14	17	51.6	38	20	09	+	199	R 15	.54	-.78								
1417+385		2	14	17	43.0	38	35	32		750	U	.47	.37	B	19.3	43.05	32.6	-.6	-.6		10
1417+397		0	14	17	44.1	39	46	05	+	15	D 57	.12	-1.01								
			14	17	49.1	39	46	04	+	18	D										
1418+388		2	14	18	58.6	38	49	25		132	U	.55	-1.12								
1419+397		2	14	19	21.6	39	47	11		315	U	.57	-.46	G	0.328	21.56	13.0	.5	-2.0		
1419+399		0	14	19	23.9	39	57	08	+	80	R 40	.14	-.44	B	17.4	23.92	8.0	-.2	.0		

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1419+419	299	4	14	19	6.4	41	58	30		2640	U	8.17	-.89	G	* 0.367	6.17	29.0	2.6	1.0		40
1420+386		2	14	20	9.8	38	40	20	+	74	D 29	.41	-.92								
			14	20	11.3	38	40	44	+	52	D										
1422+395		1	14	22	16.1	39	35	07	+	52	T136	.26	-.44	*G	0.247	17.20	24.8	12.7	12.2	2	
			14	22	18.3	39	35	37	+	50	T										
			14	22	27.9	39	35	04	+	47	T										
1422+401B		2	14	22	27.9	40	06	51		235	U	.74	-.90								
1424+380	38.39	2	14	24	02.9	38	02	26	+	40	R 45	.78	-2.32								
1426+394		2	14	26	12.9	39	25	41	+	196	R 8	.44	-.63								
1426+398		1	14	26	41.9	39	50	34	+	53	T183	.27	-.79								
			14	26	45.5	39	52	02		5	T										
			14	26	48.6	39	53	21	+	40	T										
1427+404		2	14	27	56.6	40	24	16	+	99	D 42	.41	-.96	*G	0.132	56.20	44.7	13.7	-9.6	4	
			14	27	58.2	40	24	54		21	D										
1428+380		2	14	28	31.5	38	03	07	+	70	D 41	.52	-1.09	G	0.220	32.18	24.4	.1	-.0	5	
			14	28	33.0	38	03	45	+	59	D										
1428+385		3	14	28	52.1	38	32	54		87	T 41	.82	-1.01								
			14	28	53.4	38	32	35		70	T										
			14	28	55.5	38	33	05		70	T										
1429+392		1	14	29	36.1	39	12	48		91	U	.31	-.96								
1429+395		1	14	29	42.3	39	32	18		131	U	.38	-.83								
1430+399		0	14	30	25.1	39	57	40		19	D 66	.14	-1.23								
			14	30	29.9	39	58	22		10	D										
1432+382		3	14	32	56.4	38	17	54	+	458	R 26	1.09	-.68	G	0.282	56.24	55.2	1.9	-1.2		50
1432+389		1	14	32	14.0	38	56	03		57	U	.22	-1.06								
1432+397A		0	14	32	19.8	39	42	39		42	U	.10	-.68								
1432+397B		0	14	32	57.1	39	47	13	+	54	R 8	.15	-.80	N	18.5	56.88	14.0	2.5	-1.0		
1432+428B		3	14	32	32.4	42	49	25		740	U	.93	-.18								
1435+383		2	14	35	32.9	38	20	42		182	U	.54	-.85	B	17.0	33.32	43.7	-4.9	-1.6		
1435+429	42.38	3	14	34	59.0	42	57	15	+	263	R 20	.92	-.98								
1436+399		0	14	36	36.9	39	55	56		11	D 99	.13	-1.05								
			14	36	45.5	39	55	50		23	D										

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1437+397		1	14	37	11.9	39	46	33		36	D 10	.24	-.85								
			14	37	12.4	39	46	24		45	D										
1437+427	42.39	4	14	37	50.8	42	46	53	+	330	D 54	1.91	-1.18								
			14	37	52.8	42	47	43	+	94	D										
1438+382		2	14	38	11.5	38	14	46		185	U	.67	-1.01								
1438+385		3	14	38	22.4	38	33	05	+	947	R 8	.96	-.01								
1438+406		2	14	38	30.9	40	41	17	+	89	R 8	.40	-1.18	G	0.282	31.24	14.3	-3.9	2.7		
1441+409		4	14	41	3.1	40	57	10		869	U	1.68	-.52								
1442+383		2	14	42	17.8	38	21	12	+	151	R 8	.42	-.80								
1442+384		2	14	42	37.3	38	29	57	+	85	D 72	.46	-.94								
			14	42	43.5	38	29	52	+	53	D										
1442+441	44.23	3	14	42	57.4	44	09	55		110	D 49	1.24	-1.13								
			14	42	58.8	44	10	42		185	D										
1444+395		0	14	44	30.1	39	33	32		58	U	.15	-.74								
1444+417A	41.28	3	14	44	28.3	41	45	46	+	200	D 97	1.58	-1.03	B	18.2	32.36	50.1	3.2	1.4	4	
1444+417C			14	44	37.0	41	45	57	+	225	D										
1445+410		3	14	45	17.3	41	00	15	+	309	R 8	.98	-.91	G	0.180	17.32	15.2	-.2	-.2		
1446+399		0	14	46	5.4	39	55	26	+	20	D205	.18	-.74								
			14	46	16.1	39	58	08	+	50	D										
1446+440	44.24	3	14	46	40.1	44	04	39	+	260	D 51	1.42	-.90								
			14	46	42.4	44	05	24	+	193	D										
1447+380		2	14	47	18.1	38	01	06	+	88	R 8	.42	-1.22								
1447+400		3	14	47	32.5	40	00	28		124	D 40	.97	-1.08								
			14	47	34.2	40	01	04		122	D										
1447+402	40.33	3	14	47	4.6	40	13	36	+	228	T 99	1.12	-.77	G	0.132	4.76	9.6	10.3	-8.0	5	50
			14	47	4.8	40	12	29	+	102	T										
			14	47	9.3	40	12	12		91	T										
1449+380		2	14	49	1.2	38	03	50	+	159	R 20	.52	-.93								
1449+421	41.29	4	14	49	14.2	42	07	00		682	U	2.37	-.98								
1450+390		1										.39									
1450+391B	39.42	3	14	50	7.8	39	08	37	+	500	D235	1.68	-.61	*G	0.100	12.20	30.3	-42.1	28.7	5	
1450+392		1	14	50	10.0	39	12	31	+	270	D										
1450+396		0	14	50	7.4	39	38	07	+	35	R 12	.11	-.90								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
1451+396		1	14	51	53.3	39	37	59		87	U	.32	-1.02								
1452+394		1	14	52	59.4	39	23	58	+	97	R 11	.24	-.71	G	0.198	59.42	56.1	-.2	1.9		
1453+397		1	14	53	34.5	39	46	03		73	U	.29	-1.08								
1454+394		1	14	54	24.2	39	26	48		74	U	.26	-.98								
1455+399		0	14	55	52.1	39	57	58	+	20	R 18	.10	-1.26								
1455+421	42.40	3	14	55	49.0	42	10	48		372	U	1.39	-1.03								
1457+388A		2	14	57	14.0	38	48	39	+	77	D 23	.51	-1.11	*B	18.9	14.61	25.0	1.6	6.5	4	
			14	57	15.5	38	48	24	+	47	D										
1458+433		3	14	58	40.5	43	21	41		370	U	1.08	-.84								
1459+399		1	14	59	8.1	39	53	49		38	D177	.22	-.65								
			14	59	21.9	39	55	09		58	D										
2300+382		2	23	00	54.4	38	12	28	+	168	R 7	.61	-1.01	G	0.220	54.52	22.2	-1.4	5.8		
2301+394A		1	23	01	28.4	39	25	13		138	U	.35	-.73								
2301+394B		1	23	01	36.1	39	27	01		55	D 15	.33	-1.00	N	20.0	37.00	55.2	2.3	-3.2	2	
			23	01	37.2	39	26	52		37	D										
2301+398		1	23	01	5.7	39	50	02	+	28	D 29	.37	-1.35	N	19.5	7.40	3.0	-3.0	1.9	5	
			23	01	8.2	39	50	07	+	38	D										
2301+430	42.54	3	23	01	30.9	43	01	34	+	141	T 98	1.01	-1.18	N	16.0	30.14	11.2	-6.4	-9.9	5	
			23	01	28.8	43	00	51		3	T										
			23	01	27.2	43	00	04	+	80	T										
2301+443	44.42	4	23	01	27.8	44	22	55		966	U	4.04	-1.12								
2302+396		0	23	02	27.2	39	40	32	+	9	R 30	.15	-2.20								
2302+402	40.48	4	23	02	34.8	40	12	41		1163	U	2.53	-.61								
2303+391A	39.72	4	23	03	45.3	39	09	33	+	360	D167	1.85	-.89	G	* 0.206	43.84	6.4	1.9	-10.9	4	34
			23	03	42.7	39	12	18	+	236	D										
2304+377		4	23	04	39.4	37	46	28		1354	U	2.60	-.51	N	20.0	39.45	27.5	-.6	.5		18
2304+398		1	23	04	44.3	39	53	55	+	63	R 25	.20	-.90								
2304+429		3	23	04	12.1	42	54	15	+	138	D 41	.84	-.94								
			23	04	13.8	42	54	52	+	116	D										
2305+404	40.49	3	23	05	33.1	40	25	34	+	189	R 7	.99	-1.30								
2306+392		1	23	06	34.6	39	17	19	+	71	R 12	.22	-.89								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
2308+393		1	23	08	34.5	39	22	41	+	36	D 45	.24	-.89								
			23	08	36.3	39	22	01	+	41	D										
2308+395		0	23	08	19.4	39	35	24		33	U	.11	-.94								
2308+400		2	23	08	32.7	40	03	51		146	U	.63	-1.15								
2311+387		2	23	11	37.5	38	45	29		238	U	.50	-.58								
2311+396A		0	23	11	27.9	39	36	42	+	35	R 76	.12	-.96	N	18.5	27.90	42.0	.0	.0		
2311+396B		0	23	11	53.8	39	38	51		32	U	.13	-1.10								
2311+469	46.47	4	23	11	29.2	46	55	55		1802	U	4.34	-.69	B	17.5	29.14	54.5	.6	.5		51
2313+406		2	23	13	46.9	40	39	24	+	127	D 59	.68	-.99	*G?	0.026	47.42	47.3	-2.5	7.2	4	
			23	13	47.5	40	38	25	+	66	D										
2315+396	39.73	4	23	15	25.0	39	35	49	+	586	D 49	2.50	-.81	N	20.0	25.02	11.9	.9	1.6	4	24
			23	15	25.2	39	36	38	+	308	D										
2316+398		1	23	16	56.4	39	53	51	+	42	D 26	.29	-.93	B	20.0	57.66	38.5	-3.5	4.0	5	
			23	16	58.2	39	53	35	+	47	D										
2318+389		1	23	18	54.8	38	55	48		59	U	.22	-1.03								
2320+396		1	23	20	19.6	39	41	23	+	33	D 22	.26	-.82								
			23	20	20.2	39	41	44	+	58	D										
2320+416B		3	23	20	18.7	41	41	22	+	145	D116	1.07	-1.49	*G	0.165	19.46	45.7	-2.7	-14.5	5	
2320+417			23	20	24.2	41	43	00	+	15	D										
2321+423		4	23	21	30.3	42	19	19		688	U	2.07	-.86								
2322+396		0	23	22	52.8	39	41	06		126	U	.12	.04	B	17.8	52.78	7.2	.2	-1.2		
2322+403		3	23	22	23.8	40	23	52		317	U	.85	-.77								
2323+388		2	23	23	14.8	38	49	25		154	U	.48	-.89								
2323+398		0	23	23	6.0	39	51	57		19	D 16	.17	-1.28								
			23	23	6.4	39	51	41		14	D										
2323+435A		4	23	23	18.4	43	30	28		1626	U	2.90	-.45	G	0.220	18.34	28.2	.6	-.2		13
2324+394B		1	23	24	42.6	39	26	48	+	47	R 18	.26	-1.34								
2324+394A		0										.17									
2324+395		0										.19									
2324+399		1										.24									

TABLE II. (continued)

1	2	3	4 5 6	7 8 9	10	11	12	13	14	15	16	17	18	19	20	21	22
2324+405	40.50	4	23 24 30.0	40 31 30	+	990	D 16	5.52	-.58								
			23 24 30.9	40 31 43	+	1652	D										
2325+396		1	23 25 33.2	39 39 10		58	U	.28	-1.23								
2326+395		2	23 26 24.9	39 31 25		198	U	.51	-.74								
2326+422		3	23 26 33.0	42 15 23		201	D 16	1.02	-.72	G	0.247	33.10	31.5	-1.1	-.4	5	
			23 26 33.0	42 15 39		207	D										
2327+391		1	23 27 58.5	39 10 39		86	U	.36	-1.12								
2327+407		2	23 27 42.7	40 47 52	+	234	R 6	.48	-.56	B	17.5	42.88	51.8	-2.0	.2		
2327+422		3	23 27 39.5	42 17 13	+	312	R 17	1.12	-1.00								
2328+397		0	23 29 0.0	39 46 58	+	20	R 80	.11	-1.33								
2329+398		0	23 29 45.5	39 50 49	+	22	D 34	.18	-1.12	B	21.2	46.68	43.8	2.7	1.3	5	
			23 29 48.4	39 50 41		21	D										
2330+387		3	23 30 35.4	38 44 38		664	U	1.12	-.41								
2330+389		1	23 30 1.3	38 58 28		76	U	.30	-1.08								
2330+397		0	23 30 27.2	39 45 21		68	U	.18	-.76								
2330+402		4	23 30 26.2	40 14 02		803	U	1.78	-.62								
2330+435	43.58	3	23 30 54.5	43 30 08		146	D 41	1.15	-1.25	N	15.4	56.10	15.0	2.7	-2.5	4	
			23 30 58.2	43 30 17		88	D										
2331+399		3	23 31 17.2	39 55 53	+	255	R 20	.83	-.93								
2332+386		2	23 32 34.8	38 49 34		105	U	.42	-1.09								
2332+399		1	23 32 7.1	39 56 47		147	U	.39	-.76								
2333+397		0	23 32 59.1	39 43 57		15	T 66	.16	-.96								
			23 33 1.4	39 42 56	+	15	T										
			23 33 3.6	39 43 48		17	T										
2333+399		0	23 33 6.2	39 55 07		46	U	.12	-.75								
2334+398		3	23 34 26.6	39 49 28	+	280	T 49	1.01	-.63	G	0.142	28.00	8.8	-1.1	.9	4	
			23 34 28.1	39 49 14		21	T										
			23 34 29.0	39 48 47	+	151	T										
2335+392		2	23 35 43.0	39 16 56	+	170	R 14	.56	-.93	B?	19.0	43.32	55.5	-3.7	.5		
2336+381		3	23 36 24.1	38 06 42	+	183	D 21	.90	-.96	N	17.5	25.00	41.8	-1.2	5.7	4	
			23 36 25.7	38 06 53	+	82	D										
2337+398		1	23 37 15.0	39 52 07		87	U	.24	-.80	G	0.282	15.02	7.6	-.2	-.6		

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
2338+390		2	23	38	44.9	39	01	51	+	177	R 10	.62	-.98	B	19.7	44.84	49.3	.7	1.7		
2338+393		1	23	38	47.6	39	20	14		74	D139	.23	-.43								
			23	38	56.4	39	18	39		59	D										
2340+386	38.54	3	23	40	.7	38	37	55	+	139	D 27	.90	-1.10								
			23	39	59.8	38	38	20	+	83	D										
2341+396a		1	23	41	37.5	39	37	15	+	76	R 20	.37	-1.24	G	0.198	37.40	14.0	1.2	1.0		
2341+396b		1	23	41	42.9	39	36	00		19	D 73	.37	-1.97	G	0.220	44.10	31.2	-3.3	3.9	5	
			23	41	45.4	39	34	52		11	D										
2341+399		1	23	41	54.5	39	56	12	+	98	R 17	.35	-1.00								
2342+394		0	23	42	29.2	39	29	46	+	23	R 8	.11	-1.22								
2344+429		3	23	44	52.9	42	54	12		441	U	.92	-.58	B	17.5	52.88	12.7	.2	-.7		
2347+397		0	23	47	54.1	39	45	21	+	31	R 15	.12	-1.06								
2348+387		2	23	48	1.4	38	46	39		48	T 80	.43	-.89								
			23	48	3.9	38	46	47		32	T										
			23	48	4.8	38	46	39		58	T										
2348+450	45.50	4	23	48	57.0	45	01	49		594	U	1.80	-.87								
2349+396		2	23	49	31.8	39	38	21		162	U	.51	-.90								
2349+410	41.46	3	23	49	21.5	41	04	33		394	U	1.38	-.98	B	19.2	21.46	34.0	.5	-1.0		
2350+395		0	23	50	17.2	39	31	13		195	U	.15	.21								
2351+394		2	23	51	8.3	39	26	55	+	68	D 35	.59	-1.08								
			23	51	10.2	39	26	27	+	81	D										
2351+398		0	23	51	9.1	39	49	00		21	U	.11	-1.30								
2351+400B	40.51	4	23	51	21.7	40	01	26		265	D 83	1.98	-1.05								
			23	51	28.6	40	01	01		254	D										
2351+456	45.51	4	23	51	50.0	45	36	22		1569	U	2.21	-.27	N	20.0	50.04	23.0	-.4	-1.0		25
2352+385		2	23	52	5.8	38	30	53	+	148	R 8	.47	-.91								
2354+397		0	23	54	32.7	39	46	53		35	D 79	.20	-1.10								
			23	54	36.2	39	45	47		14	D										
2354+471	47.63	4	23	54	52.6	47	08	19	+	860	D258	4.40	-.72	G	* 0.046	57.10	41.2	10.9	-15.7	5	16
2355+471A		4	23	55	3.5	47	10	29	+	900	D										
2355+397		1	23	55	38.7	39	45	11		54	U	.33	-1.42								

TABLE II. (continued)

1	2	3	4	5	6	7	8	9	10	11	12	13	14	15	16	17	18	19	20	21	22
2355+398		2	23	55	28.0	39	49	47		159	U	.56	-.99	B	18.0	27.68	45.0	3.7	2.0		
2356+390		3	23	56	26.5	39	05	47		373	U	.93	-.72								
2356+437	470	4	23	56	2.3	43	47	57		1536	U	5.36	-.98								
2357+398		1	23	57	5.3	39	49	54	+	67	R 11	.32	-1.23	G	0.198	5.24	49.6	.7	4.4		
2358+390		1	23	58	7.9	39	01	22		204	U	.37	-.47								
2358+406	40.52	4	23	58	19.4	40	37	19		1181	U	2.02	-.42								
2358+416	41.47	4	23	58	45.7	41	36	14		212	D 23	2.21	-1.15								
			23	58	46.0	41	36	37		297	D										
2359+394		3	23	59	7.4	39	29	44	+	201	D 45	.96	-.99								
			23	59	11.2	39	29	34	+	70	D										

Notes to TABLE II.

0001+398 Confused by an 8 mag. star, high reddening.

0003+397 Not detected at VLA (<1mJy).

0005+383B Peculiar triple radiosource with strong central comp.; gal. in cluster, see map.

0010+402 Could be cluster member.

0010+405 Compact galaxy.

0013+393 Object plate limit only in red 9" NE.

0015+399 The two radio components could be physically unrelated; the field is empty anyway.

0020+437 A 19.5 magnitude bluish starlike object is 8" N of radio midpoint.

0026+397 Probably a "tail" source; the galaxy coincides with the N. comp.

0028+394 A 20 magnitude galaxy is 7" S from radio centroid.

0029+394 A 15 magnitude neutral color starlike object 8" NE.

0030+394 Not detected at VLA (< 3 mJy).

0030+396 Not detected at VLA (< 3 mJy).

0031+395 Not detected at VLA (< 3 mJy); source marked as "bad" in B3 catalog.

0031+396 Not detected at VLA (< 3 mJy); source marked as "bad" in B3 catalog.

0031+398 Not detected at VLA (< 3 mJy); source marked as "bad" in B3 catalog.

0035+385A Object 18.5 mag 4" preceding the North component. Physical double with 0035+385B

0038+399 Brightest member of a faint cluster.

0039+412 The source is within a spiral arm of M31. Nothing evident as possible id.

0041+382A Stellar object 14 mag. 8" SE.

0045+393 Stellar object 9" SW, 16 mag.

0045+396 A second gal. is 6" NE, probably interacting.

0045+404 The radio map shows a faint halo (< 4 mJy peak) diam. 30".

0049+395 The two components do not appear to be physically connected.

Notes to TABLE II. (continued)

0050+401	Triple source, with B30050+403 and 0050+402B; this latter was not originally included in the VLA sample. Total 408Mhz flux is given.
0058+403	A 15 mag. neutral color starlike object 7" S.
0059+397	A 17.5 mag. object very red and compact, 3" W of the S component.
0059+461	A 19 mag object 3" S of the N component.
0107+399	Stellar object 8" N 16.5 mag., possibly blue.
0108+396	Object coincident with the midpoint, diffuse, elongated, visible only in blue, 20 mag.
0108+402	Spiral galaxy 15 mag. 15"NW.
0109+415	Galaxy in compact group.
0110+401	Also a 19 mag. galaxy 7"W.
0115+453A	The identification listed by Spinrad et al. 1985 is not confirmed: our position differs from the position listed there by 8 arc sec.
0116+438	Also a blue object, 18.5 mag. 4"S of component S.
0119+397	Object at plate limit, only in red, 7" SE.
0120+405	Object very red 9" NE.
0123+396	Galaxy in cluster.
0128+389	Not detected at VLA (< 1 mJy).
0131+390	Galaxy bluish in color.
0132+392	The two radio comp. could be physically unrelated.
0134+389	A 15 magnitude neutral color starlike object 6" S.
0143+446B	The object could be a red, compact galaxy or the chance superposition of two stars.
0146+394	The two radio comp. are probably physically unrelated. The B3 flux is probably uderestimated.
0147+398	Very red galaxy, perhaps in group.
0152+382	The radiosource could be a "tail".
0152+435	There is a faint galaxy 7" SE from the N comp.
0157+393A	The radiosource could be a "tail".
0157+442	Cluster or group of galaxies?
0158+394	Also a 18 mag. galaxy 4" NE.
0207+389	NGC 828
0207+395	The object is confused by a 13 mag. star 7" SE; it could be blue.
0211+393	Galaxy in group.
0213+407	Galaxy in cluster?
0213+412	A faint object, possibly a galaxy, 8"W.
0214+393	Double galaxy? Interacting?
0218+396	Blue starlike object, 17 mag.,6"NW from comp. S.
0218+397	Undetected at VLA (< 3 mJy); source marked as "bad" in B3 catalog.
0218+399A	A 18.5 magnitude galaxy, 30" SW of radio centroid, coincides with an extension of S component; may be the identification. Cluster.
0218+402B	Includes the source B30218+402A; the B3 positions are in error; the 408 MHz flux given is the total.
0219+397	Stellar red object 17 mag. 8" E.
0219+421	Nucleus of NGC 891; the fluxes are unreliable at both frequencies.
0219+428A	Part of 3C66. BLO (see Spinrad et al. 1985)

Notes to TABLE II. (continued)

0220+393A A 18.5 magnitude galaxy 14" NE.

0159+390 Galaxy in cluster, alternate possible identification. Radiosource structure is distorted (see map).

0201+396 A possible second radio component at 020135.4, +394055, 15mJy. The field remains empty.

0202+380 Wide double source, uncertain ident.,four other candidate id. present.

0221+393 Id. uncertain, many candidates in between the two comp. There is also a very diffuse object (galaxy?) 30 arcsec in size, 1 arcmin S of the S comp.

0222+397 Undetected at VLA (<3 mJy).

0225+389 There might be a third radio component at 022552.83, +385519.

0233+390 A 19 magnitude galaxy 7" NE of the radio centroid. A second very faint object (G?) sits on the strong radio component. The radio comp. could be unrelated.

0237+396 A 17.5 mag. galaxy 7" NW.

0240+404 The physical connection between the radio components is doubtful.

0241+393B The source 0241+393A is also in the field, but probably without physical relation (see map).The R-O shift referred to the midpoint of the two components is da=+1.6, dd=+1.3 arcsec.

0246+393 Physically connected with B30247+393A; is 3C73.

0248+467 Associated to B30247+468B; ident. is IC260. Positions and fluxes from D-config. map. The published map (erroneously named 0247+467) is C-config.

0250+384 A 19.5 mag. blue starlike object 7" E of the radio centroid.

0258+435 Includes also 0258+434B listed separately in B3. The source is a "tail". The published map is C-config. but the tabular data are from a D-config. map.

0259+391 A 20.5 mag. blue starlike object 9" S.

0220+427A Part of 3C66, physical double with B3 0219+427D.Classification, positions and fluxes are from a D-config. map; the published map is C-config.

0221+383 A 19.5 mag. red object on the line joining the components, but far from midpoint or baricenter.

0700+390 The radio components could be physically unrelated; there is a 18.5 mag. blue object some 20" from the midpoint (position 070040.21 +390321.5). The B3 flux is underestimated.

0703+426A Galaxy in cluster, coincides with the S comp.; "tail" radiosource. B3 flux is probably underestimated.

0703+426B There is a compact object, 17 mag. possibly in cluster, at 07 03 31.0 +42 37 59.

0708+388 Bluish color.

0709+393 Two radio comp. probably unrelated.

0710+403 Possible cluster or small group.

0712+391 The central radio comp. is visible on a high-resolution, A-array VLA map.

0720+381 The galaxy coincides with the N comp.; the radiosource is not adequately described by a double source model (see map).

0724+396 Bluish color, coincident with the N comp.(the data are from A-conf.) C-map shows further extension toward North ; also a very red object, 18 mag., 7" NE.

Notes to TABLE II. (continued)

0728+395	Stellar object, neutral color, 17.5 mag., 10" E.
0729+395	Galaxy in cluster, brightest member.
0733+389	Red double stellar object.
0735+388	Possibly in cluster.
0736+400	Bluish color.
0739+397	Two physically unrelated sources: the S comp. is identified with a very red, compact object, the N comp. with a 19 mag. blue object. B3 flux is for the whole complex, the spectral index for the unresolved component (N) is a lower limit.
0742+376	The two radio comp. could be physically unrelated. There is a faint, blue stellar object, visible only on the blue printat about 1 arcsec from the S component.
0753+391	A 18.5 magnitude galaxy 12" NE of radio centroid. Also a second, very faint bluish galaxy, closer to midpoint.
0754+394	D map only; the bright blue object is a confirmed quasars with z=0.09. The source is elongated in the direction of the radio-optical displacement.
0754+397	D map only.
0755+379A	See 0755+379B.
0755+379B	Connected to 0755+379A, a double-source model is not a good fit, see map. It is part of 4C37.21, NGC 2484.
0759+392	A 18.5 magnitude galaxy 9" N of the S component. A second very faint object close to the S comp.
0803+426	Very close to 0803+427; both are identified with galaxies in cluster.
0804+399	Multiple galaxy
0805+406	A second 20 mag. object 7" NW
0807+399	Could be a "tail" source.
0810+460B	A second plate limit object 10" N; both objects are visible also in blue
0811+388	Galaxy in group.
0818+472A	Galaxy in group.
0821+399	Undetected at VLA; marked as "bad" in B3.
0821+447	Cluster or group of faint galaxies in the field.
0823+384B	Other galaxies in the field; the source could be a "tail".
0827+399	Undetected at VLA; marked as "bad" in B3.
0836+402	Compact object; a second fainter object 8" S; both are visible also in blue.
0837+399	Could be a "Tail".
0847+406	A second fainter radio component at 084718.8, +404200, 30 mJy, probably unrelated; the spectral index is for the two sources together.
0855+397	Only a D-config. map (45" resolution) available.
0902+414	Galaxy 16 mag. 8" North-East
0902+416	A second 20 mag red object 3" SE.
0906+421	Group, cluster?
0908+380B	Confused at 408 MHz by 0908+380C
0912+392	Only low resolution D map. A 17 mag blue galaxy 10" following.
0914+390	D-map only (45" resolution)
0917+458A	Cluster, 3C219.
0920+408	The physical connection of the fainter component is doubtful.

Notes to TABLE II. (continued)

0936+399	**The source is confused in the B3 Catalog; the VLA map shows two unrelated fainter sources.(delete from the sample)**
0938+399A	Not detected at VLA; (< 3mJy) source marked as 'bad' in B3.
0938+399B	A very interesting source, X-shaped; see map; see also Gregorini et al., 1988.
0944+390B	Blue very faint and slightly fuzzy object.
0951+398	The radio structure is slightly distorted.
0955+387	Also a 19.5 mag galaxy 10" N of the midpoint position.
1007+417	The source is elongated in the direction of the quasar.
1012+395	D-map only.
1014+392	Galaxy in cluster.
1014+397A	A 20 mag. galaxy,closer to midpoint, has been proposed as identification by ref. 28
1016+388B	The last component seems to be physically unconnected. Empty fields anyway.
1019+397	Could be a "Tail".
1022+432	A 18.5 magnitude neutral color starlike object 11" E.
1023+393	Galaxy in cluster. Tail ?. The galaxy lies on the head of the tail.
1033+388	Confused at 408 MHz. B3 flux unreliable.
1037+399	There is a 16.5 mag red starlike object at 103721.9 395753.6
1039+397	Interacting galaxy in cluster,the source is probably a tail. The galaxy is on the head of the tail.
1040+398	D-map only.
1042+392	D-map only.
1042+393	A 19.5 magnitude galaxy 10" S.
1055+396	Two unrelated sources. The S-component coincides witha 16.0 mag galaxy, the N-component with a faint stellar object.(delete from the sample)
1055+404B	Possibly a tail source.
1104+390	D-map only; could be a physical double with 1105+390. Both sources are extended along the line joining them.
1105+390	D-map only; see note for 1104+390.
1107+389	Undetected at VLA (< 6 mJy); marked as 'bad' in B3.
1108+394	D-map only.
1108+411B	Galaxy in group. Probably a tail source.
1109+437	Some faint red objects in the field, probably a distant cluster.
1111+396B	Not detected at VLA (< 1 mJy); marked as 'bad' in B3.
1111+398	Not detected at VLA (< 1 mJy); marked as 'bad' in B3.
1119+398	Not detected at VLA (< 1 mJy); marked as 'bad' in B3.
1121+444	A very red stellar object 18 mag. is 7" from comp. S
1122+390	NGC 3665. The source is extended in a direction roughly perpendicolar to optical major axis.
1123+395	There is also a plate limit object or whisp closer to radio-position, visible only in red.
1133+432	Faint cluster in the field.
1137+396	Faint cluster in the field.

Notes to TABLE II. (continued)

1141+374	The source is resolved as a double even in B3 Catalog, but the southern component was lost because it is out of the Catalog limit in Dec. The 408 MHz flux given is the total. Identification is an interacting galaxy in a compact group.
1142+392	The radiosource is elongated in the direction of the blue object.
1151+395	Undetected at VLA (< 3 mJy); marked as 'bad' in B3.
1151+456	Compact galaxy, z=0.192.
1154+397	Complex source, could be a wide tail. VLA flux uncertain, the source is diffuse (see map).
1206+399	The B3 source becomes two sources (see map); both are identified with galaxies in the same cluster; the B3 flux has been split evenly between the two.
1208+396	Spiral galaxy NGC4151 z=0.003 (ref.)
1209+396	Possibly a "tail", see map.
1211+388	A 19.5 magnitude galaxy 7" SE. The source is extended but not in direction of the galaxy.
1217+427	There is a 19 mag galaxy 6" NE of the radio-position.
1218+421	Bright spiral galaxy (UGC7416)
1220+408	There is a 18 mag stellar object visible only on the blue print, 7" NW.
1221+398	The galaxy could be in a cluster. The B3 flux could be underestimated.
1222+423	3C272. The (uncertain) identification listed by Spinrad et al. 1985 is not confirmed by our data.
1223+395	There is a 19 mag. galaxy 9" SW.
1228+397	There are also two red 18 mag galaxies in the field, one very close to the N component, at 122827.42 394645.
1228+419A	Includes 1228+419B and 1227+419. The ID is NGC 4485. The optical position is not well defined on PSS print.
1232+394	A 18 mag. blue starlike object 7" SE.
1232+397A	There is a 20 mag galaxy 8" W; however the source is not extended in this direction.
1233+418	Brightest cluster member.
1236+444B	The galaxy is very close to a probably real central radio component, not listed (S408 < 100 mJy). There is also a 17.5 mag blue object on the line joining the component, at 123627.57 442553.9 .
1241+411	Very red compact object.
1249+475	There is a blue 19 mag galaxy 9" out of the line joining the components, at 124956.62 473207.0
1255+448	There is a 18 mag. galaxy at 125543.41, +445140 . The source is a tail, see map. The double structure is not a good fit.
1256+392	The blue object lies on the axis of the sources which could be a "tail" or an unequal double.
1259+395	D-map only; the blue object lies on the axis of the sources which is asymmetric and could be an unequal double.
1300+397	The blue object lies on the axis of the sources which is asymmetric and could be an unequal double.

Notes to TABLE II. (continued)

1309+412A	The source includes B31309+413 and B31309+412. The 408 MHz flux is the total.
1311+419	The source is slightly distorted. A 19 magnitude galaxy is out of the line joining the components, 10" from the midpoint. There is a third source in the field, at 131148.7 415711 ,with 1460 MHz flux of 13 mJy, coinciding with a bluish 19 mag galaxy possibly in cluster.
1317+389	There is a 20 mag blue object 7" SE.
1318+428C	Double with 1319+428A. Total 408 MHz flux is given.
1324+390	There is a 19.5 neutral color object on the line joining the components, about 13" N from the radio centroid.
1327+398	D-map only.
1327+474C	M51. The 408 MHz flux is inclusive of B31327+427A and B.
1328+396	There is a 18.5 neutral color object about 8" NE from the radio source.
1332+385	Compact galaxy in cluster.
1333+412	Brightest cluster member.
1335+391	Not detected at VLA (< 3 mJy); source marked as "bad" in B3 catalog.
1336+393	D-map only.
1336+397A	D-map only.
1336+397B	D-map only; source confused at 408 MHz by sidelobes of 3C288=1336+391A.
1336+395	Not detected at VLA (< 6 mJy); source marked as "bad" in B3 catalog.
1336+396A	Not detected at VLA (< 3 mJy); CHI SQUARE very high in B3 catalog.
1336+396B	Not detected at VLA (< 3 mJy); source marked as "bad" in B3 catalog.
1336+401	Not detected at VLA (<1.5 mJy); source marked as "bad" in B3 catalog.
1337+396	Not detected at VLA (< 3 mJy).
1338+391	Not detected at VLA (< 3 mJy); source marked as "bad" in B3 catalog.
1342+389A	Also a 16 mag. red object at 134214.8 385630 .
1342+389B	The proposed identification is uncertain, the source is elongated in direction of R-O discrepancy.
1344+397	There is a 19.5 mag galaxy 8" W.
1349+394	Probably two unrelated sources, if so they should fall below the flux limit of the sample.
1352+397	A 20 mag blue object at 135206.3 394258 is 6" out of the line and closer to the faint component.
1353+380	A 19 mag galaxy at 135346.36 380004.2, very close to the stronger component but outside the line joining the components.
1357+392	D-map only.
1407+388	Galaxy possibly in distant cluster.
1414+398	The galaxy coincides with the S component; the two radio component could be unrelated. They are not resolved in B3.
1419+399	A second radio component at 141916.2 395449 (1460 Mhz flux is 20 mJy) probably unrelated. D-map only.

Notes to TABLE II. (continued)

1420+386	A 20 mag object 12" SE of the second component.
1422+395	The galaxy is between the first two component. In the field of third component there are : a 16 mag very red compact object, 6" SW; a 19 mag blue object (visible only in blue) 5" N (see map). The B3 flux is probably underestimated.
1424+380	D-map only. Very steep spectrum, there is a rich cluster in the field. No obvious identif.
1427+404	The source is distorted and asymmetric, the identification is very probable.
1428+380	Double galaxy.
1430+399	The two component could be physically unrelated.
1436+399	There is a 18 mag blue object 9" W of the first component.
1437+397	There is a third component at 143719.0, +394535 (1460 MHz flux is 35 mJy) coinciding with a 15 mag compact object, perhaps a galaxy.
1438+385	The source is a flat spectrum, but looks resolved. A very red stellar object, m≈18, 8"E and 3"N of the radio position, has been proposed as identification by ref. 50.
1442+441	There is a 19 mag galaxy, perhaps in cluster, some 10" E of the midpoint.
1444+417A	The source is physical double with 1444+417C, listed in B3 but outside the sample area. The 408 MHz flux given is total.
1446+399	The physical connection of the two component is not certain.
1446+440	There is a 19 mag. galaxy at 144641.35, +440449, about 10" S from the radio centroid, not on the line joining the two components. The source is simmetrical and undistorted.
1447+402	Double galaxy in cluster, radio structure distorted.
1450+390	Part of 1450+391B.
1450+391B	Distorted and diffuse, listed in B3 as three sources: 1450+390, 1450+391B, 1450+392. The 408 MHz flux given is the sum and is underestimated.(see map). Identification is very probable, galaxy in cluster.
1457+388A	Very blue object, not on the line joining the component, possible identification.
1458+433	D-map only.
1459+399	B3 position (and flux) correspond to the second component. Physical connection is probable. Empty fields anyway.
2302+396	D map only, probably resolved. The B3 flux is not very good, but the spectrum is certainly very steep.
2311+396A	Tabular data are from a D-config. map; published map is C-config.
2313+406	There is a 18 mag. compact object, perhaps a galaxy, on the line joining the components 7" S from midpoint.
2323+435A	Double Galaxy.
2324+394A	Not detected at VLA (<1.5 mJy); very high CHI SQUARE in B3 catalog.
2324+395	Not detected at VLA (< 4 mJy); source marked as "bad" in B3 catalog.

Notes to TABLE II. (continued)

2324+399 Not detected at VLA (< 1 mJy); source marked as "bad" in B3 catalog.

2328+397 D-map only.

2329+398 Also a 20 mag neutral object 6" N.

2332+388 There is a 19 mag galaxy 7" S.

2332+399 There is a 15 mag neutral object 5" NE.

2333+397 Many objects in the field, no obvious identification.

2338+393 A 17.5 mag. red compact object close to S comp. (pos. 233855.2 391855.0).

2341+396 Also a faint double source on the field (see map), unrelated to the main source, with total flux about 30 mJy. The B3 flux includes this double, which is identified with a 18 mag. galaxy.

2351+394 There is a 16.5 mag. neutral color object about 8" NE from the radio centroid.

2351+398 D-map only. Angular diameter < 20".

2354+397 Physical connection of components doubtful. Two plate limit objects on the line joining the components.

2354+471 Source is DA613. Tabular data are from a D-config. map; the published map is C-config.; the optical position is not very accurate.

2356+437 3C470. See map. The component NE is real, and approximately of 70 mJy, probably unrelated. Empty field.

2357+398 Galaxy in cluster.

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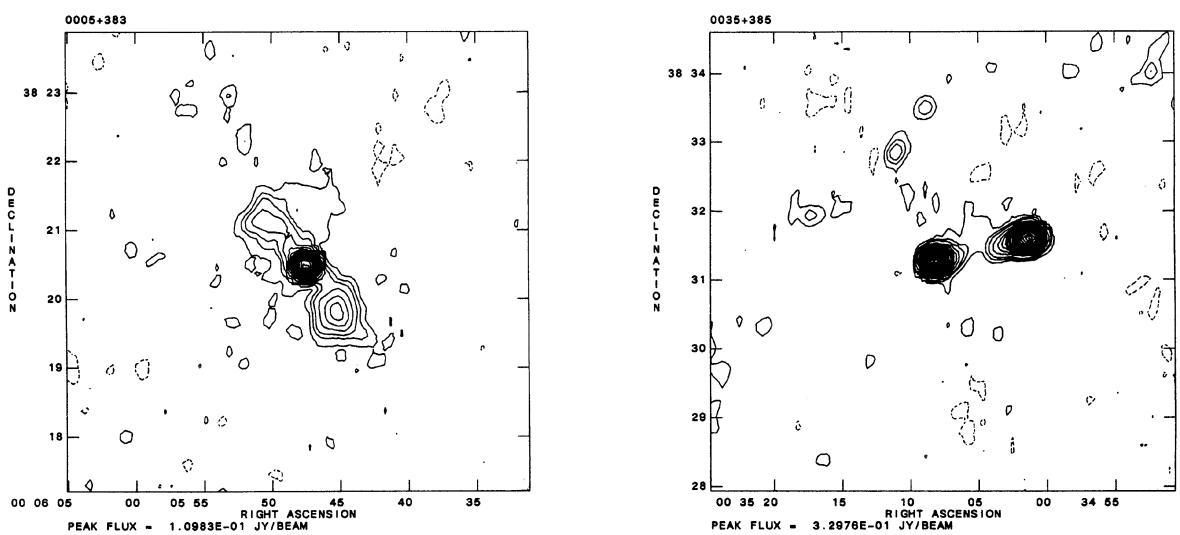
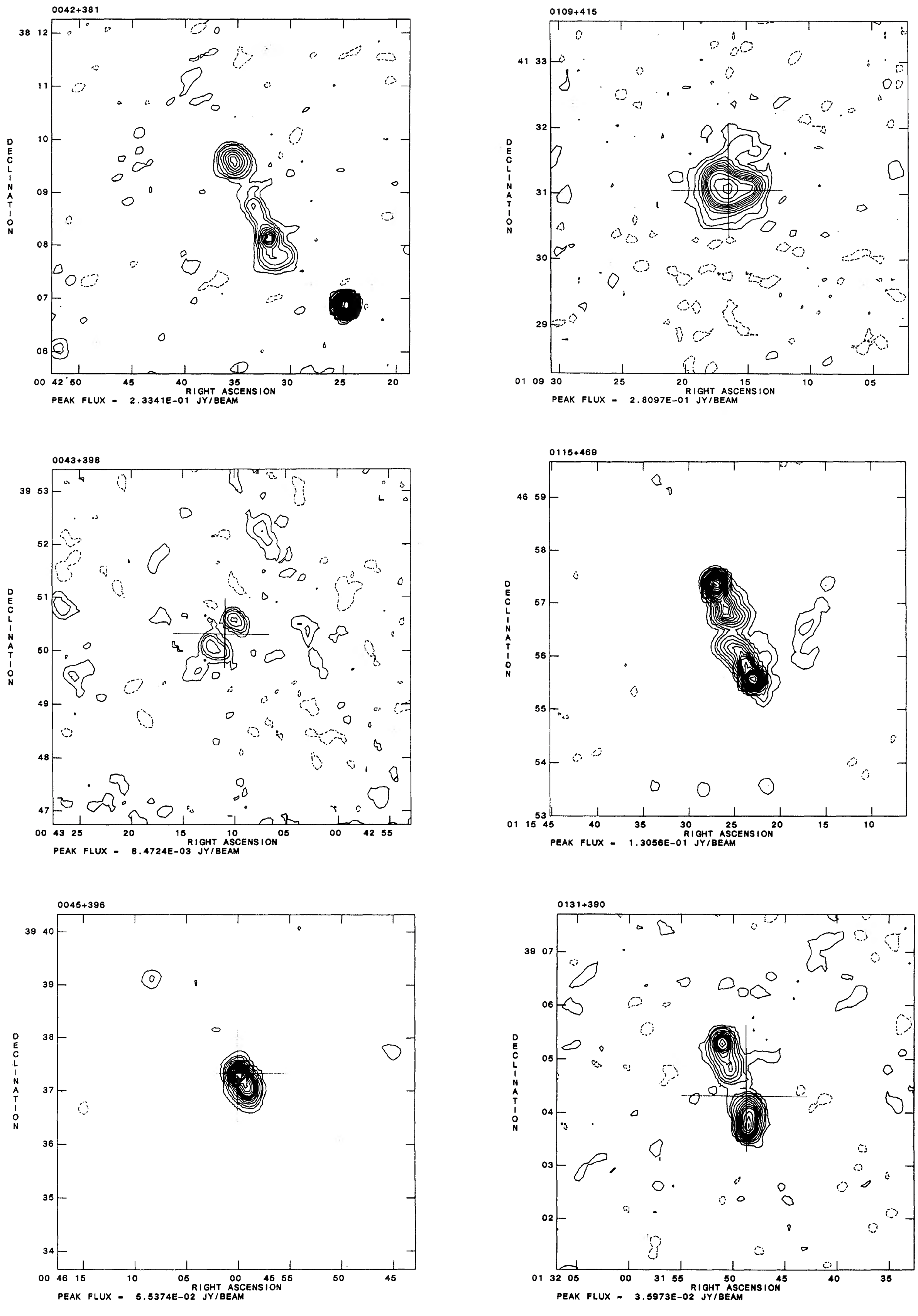


FIG. 3. Maps of radio sources not adequately described by the numerical data of Table II.



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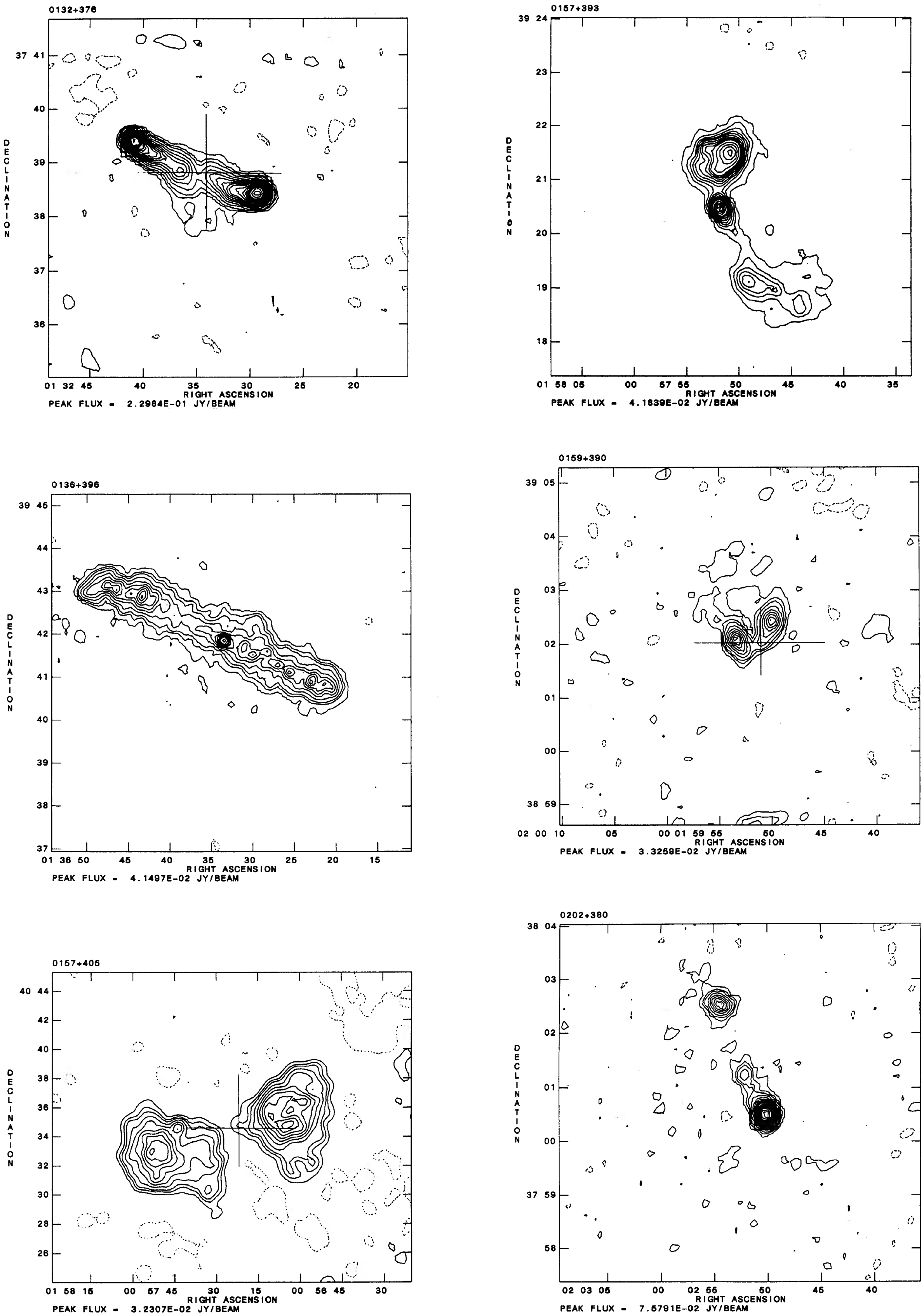

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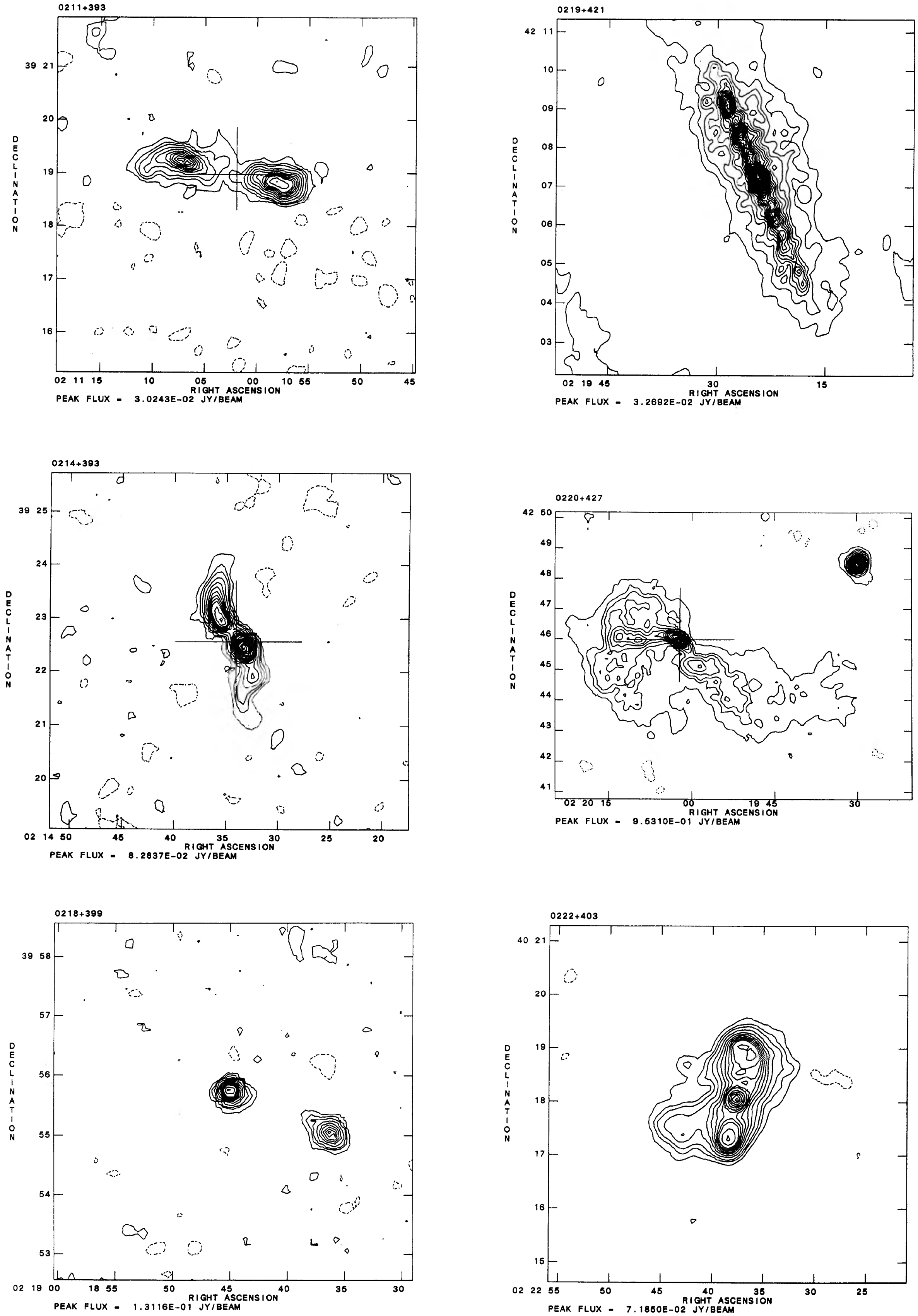

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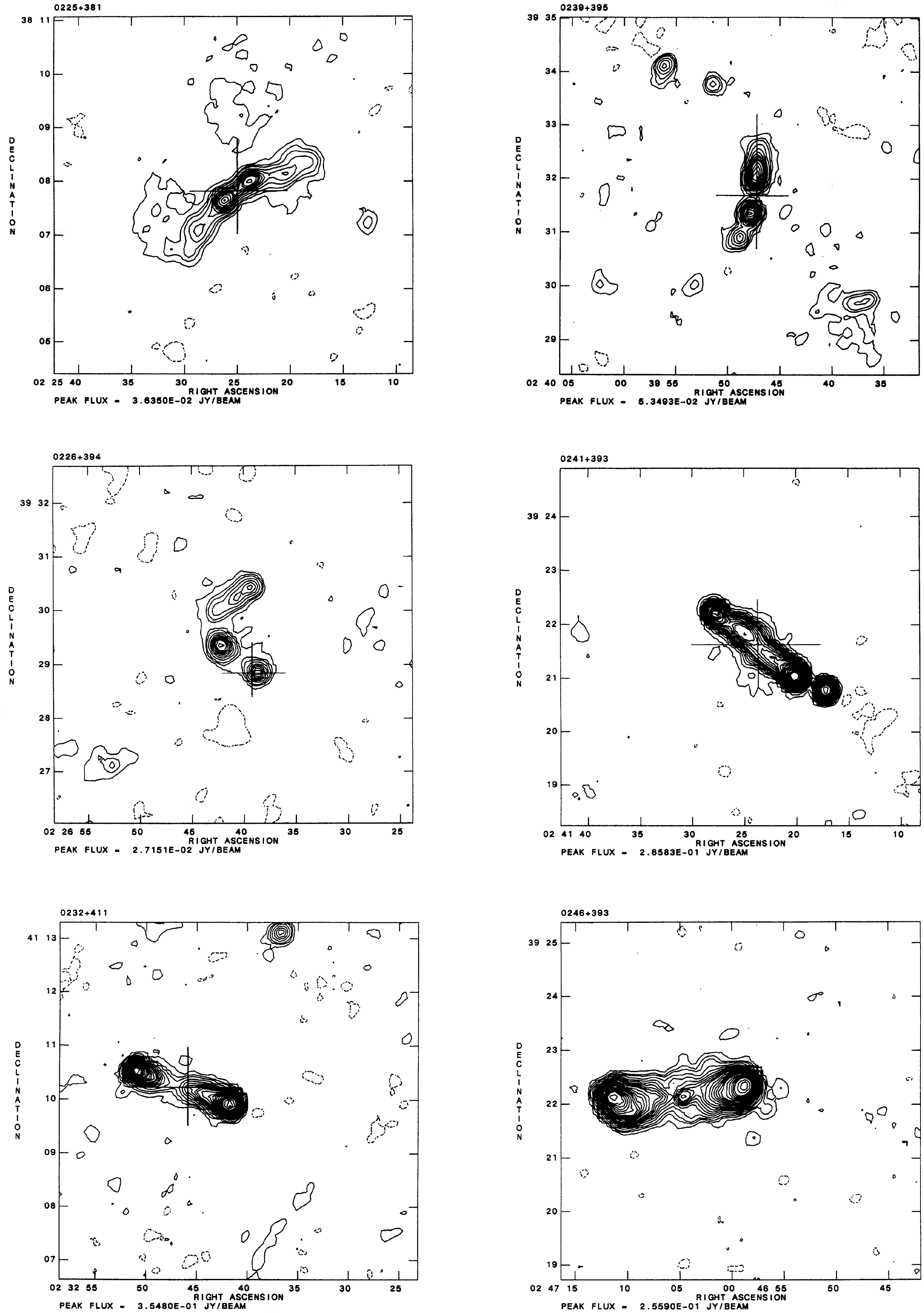

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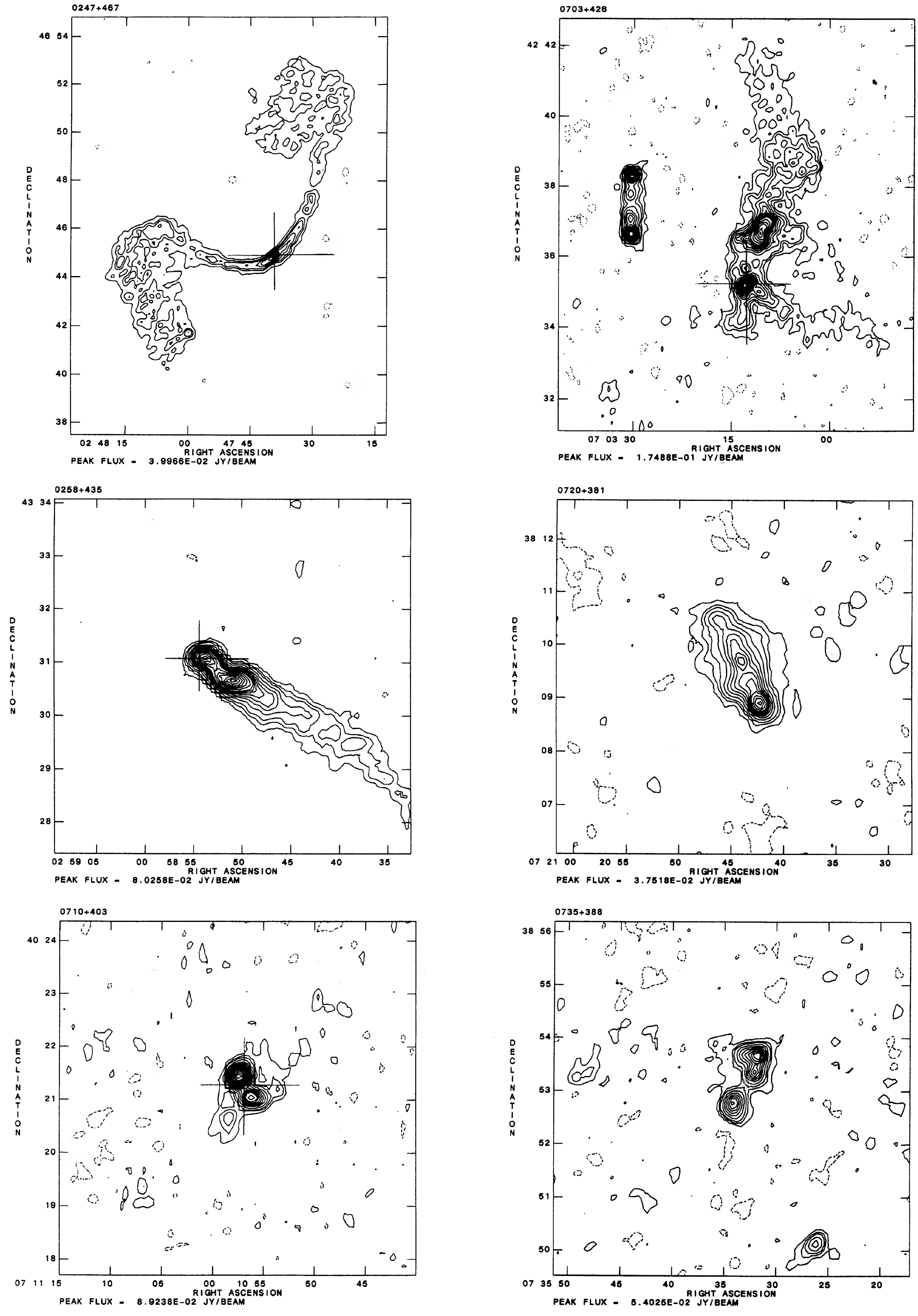

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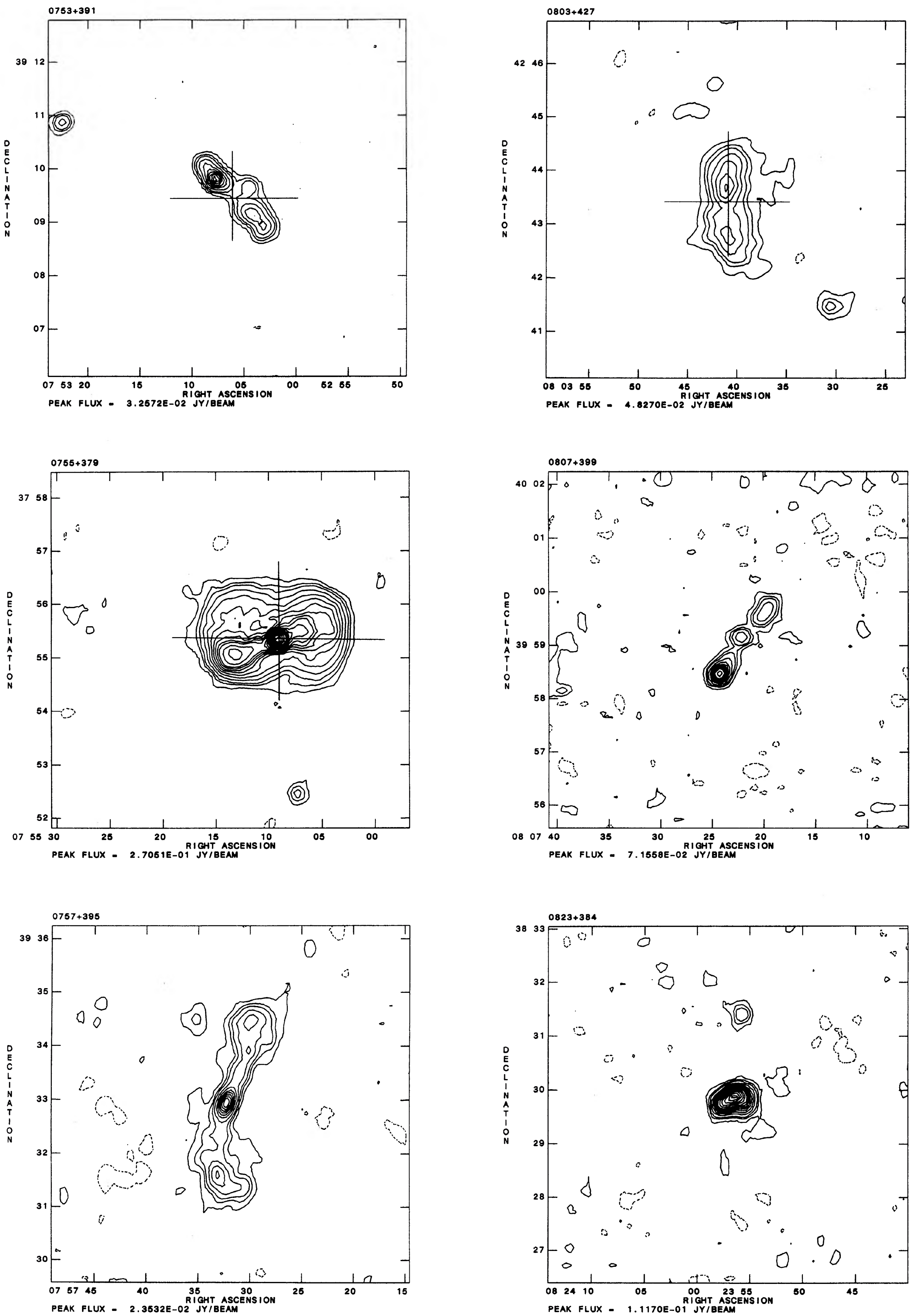

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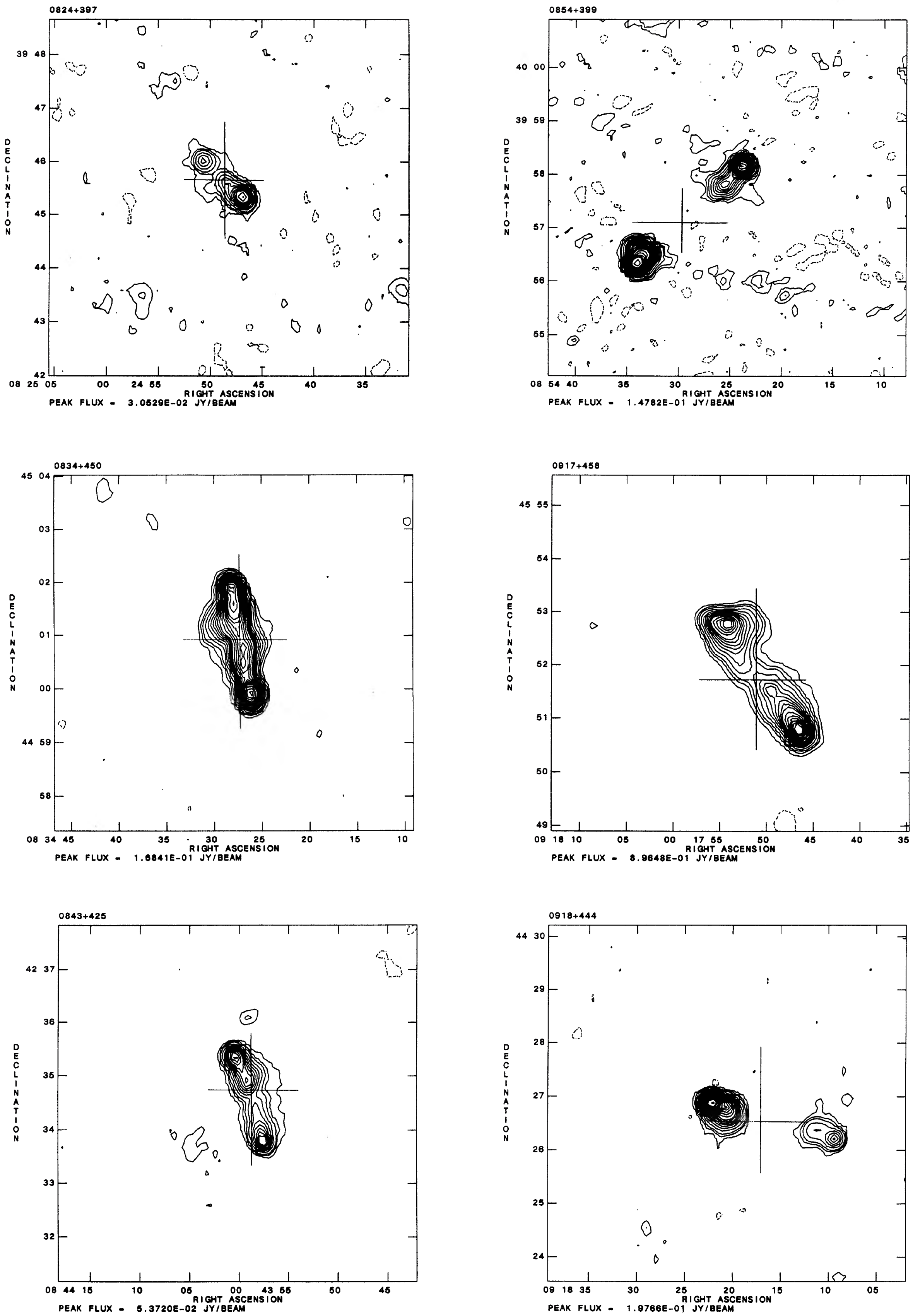

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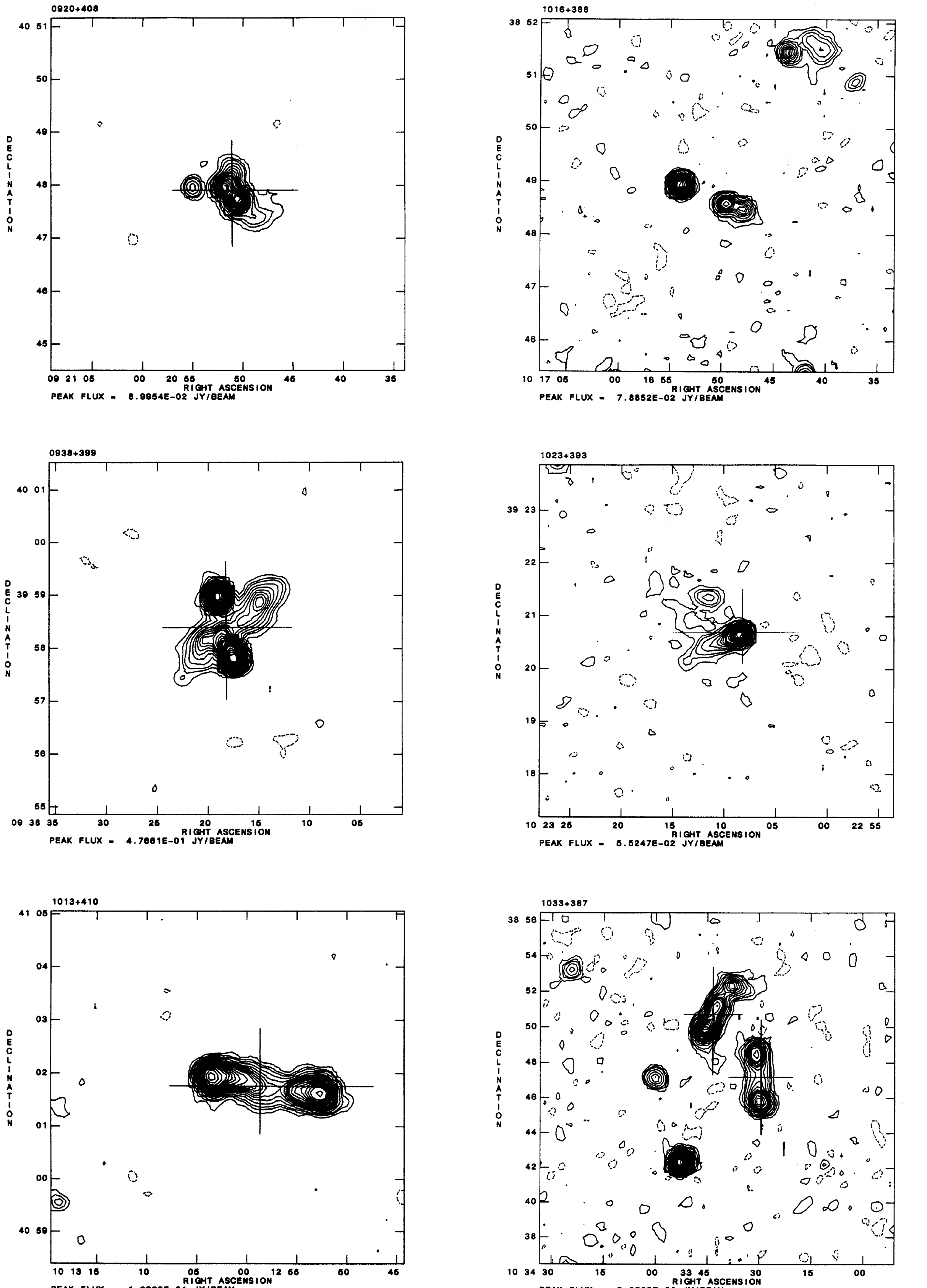

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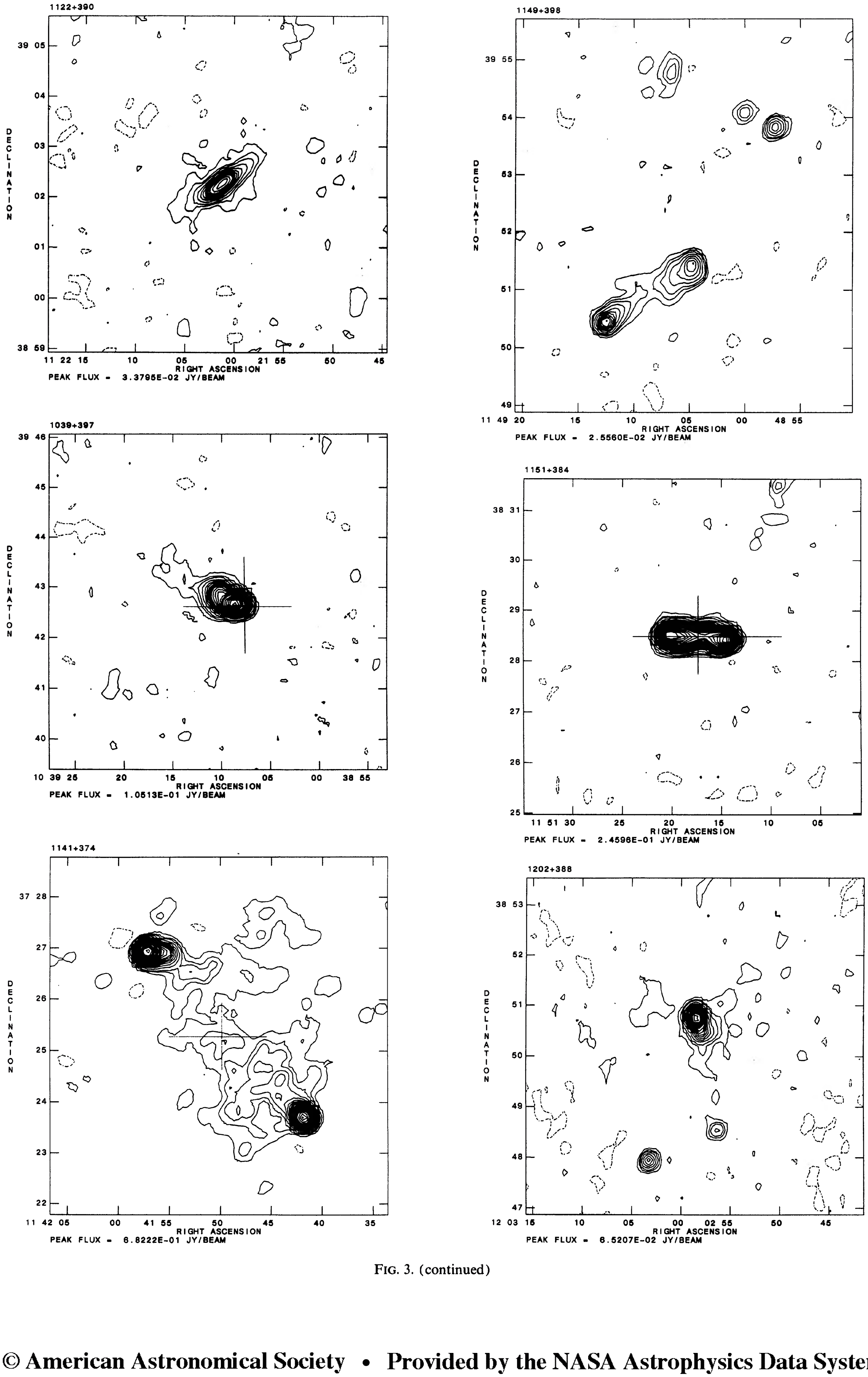

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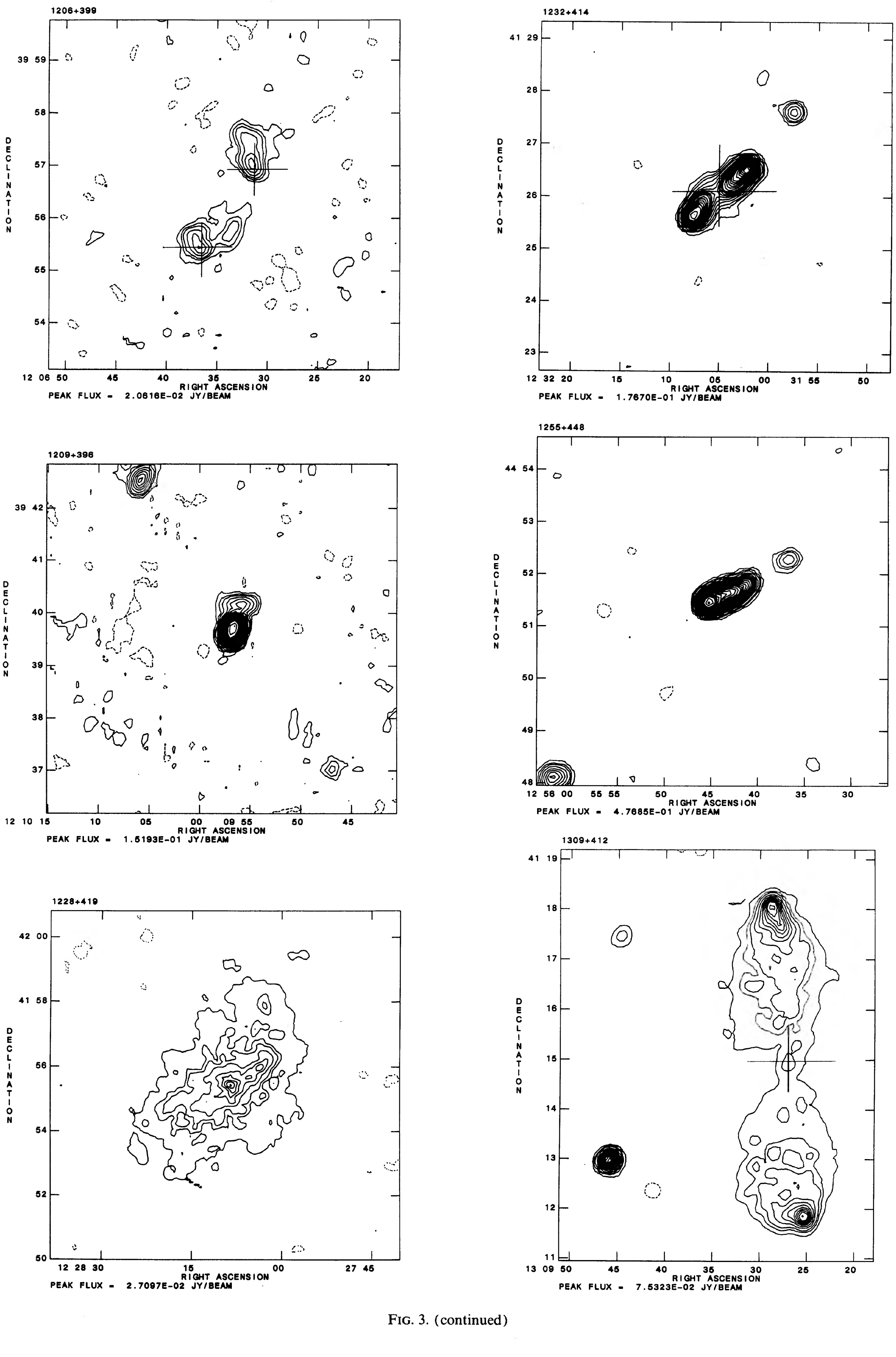

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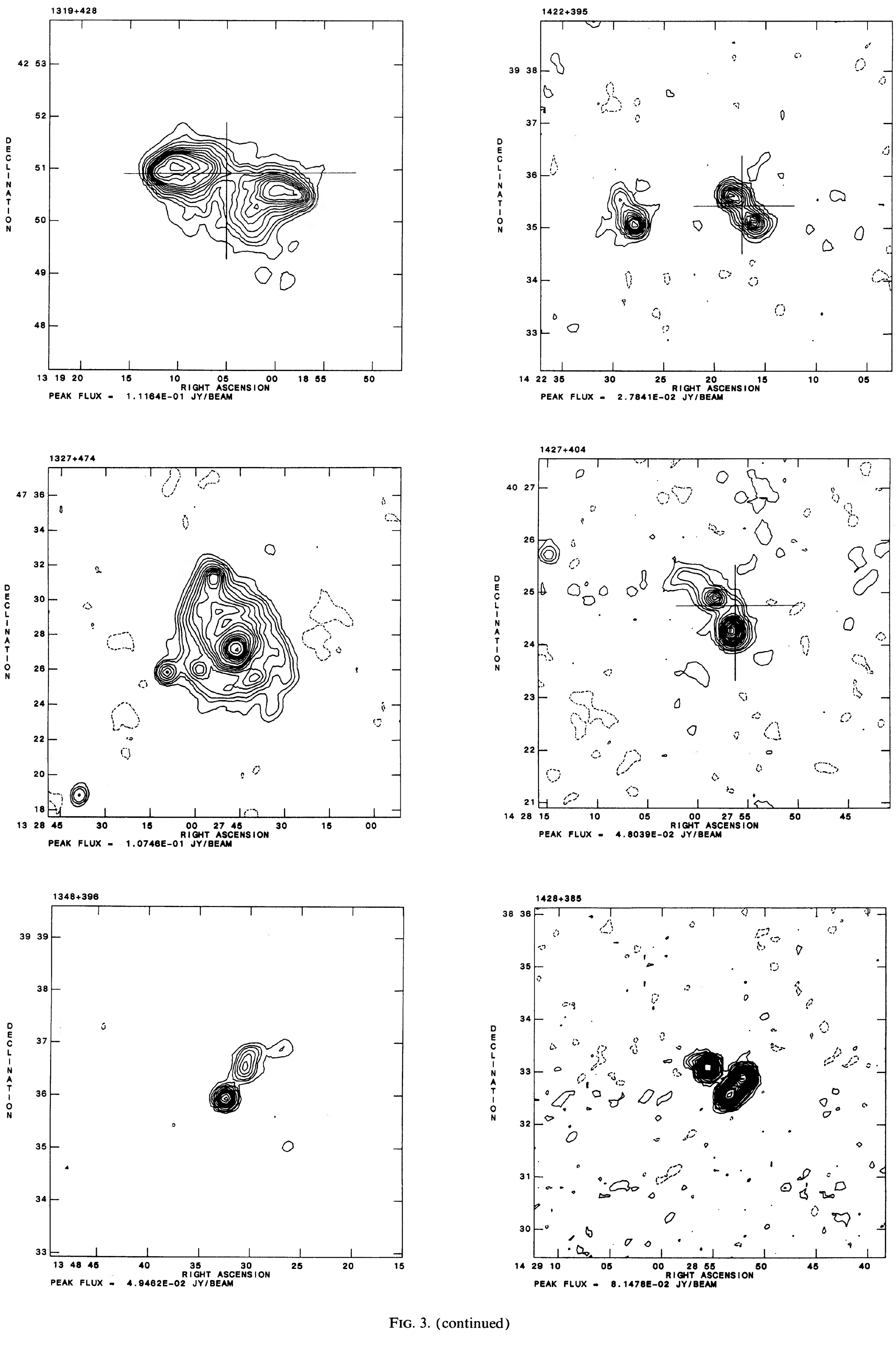

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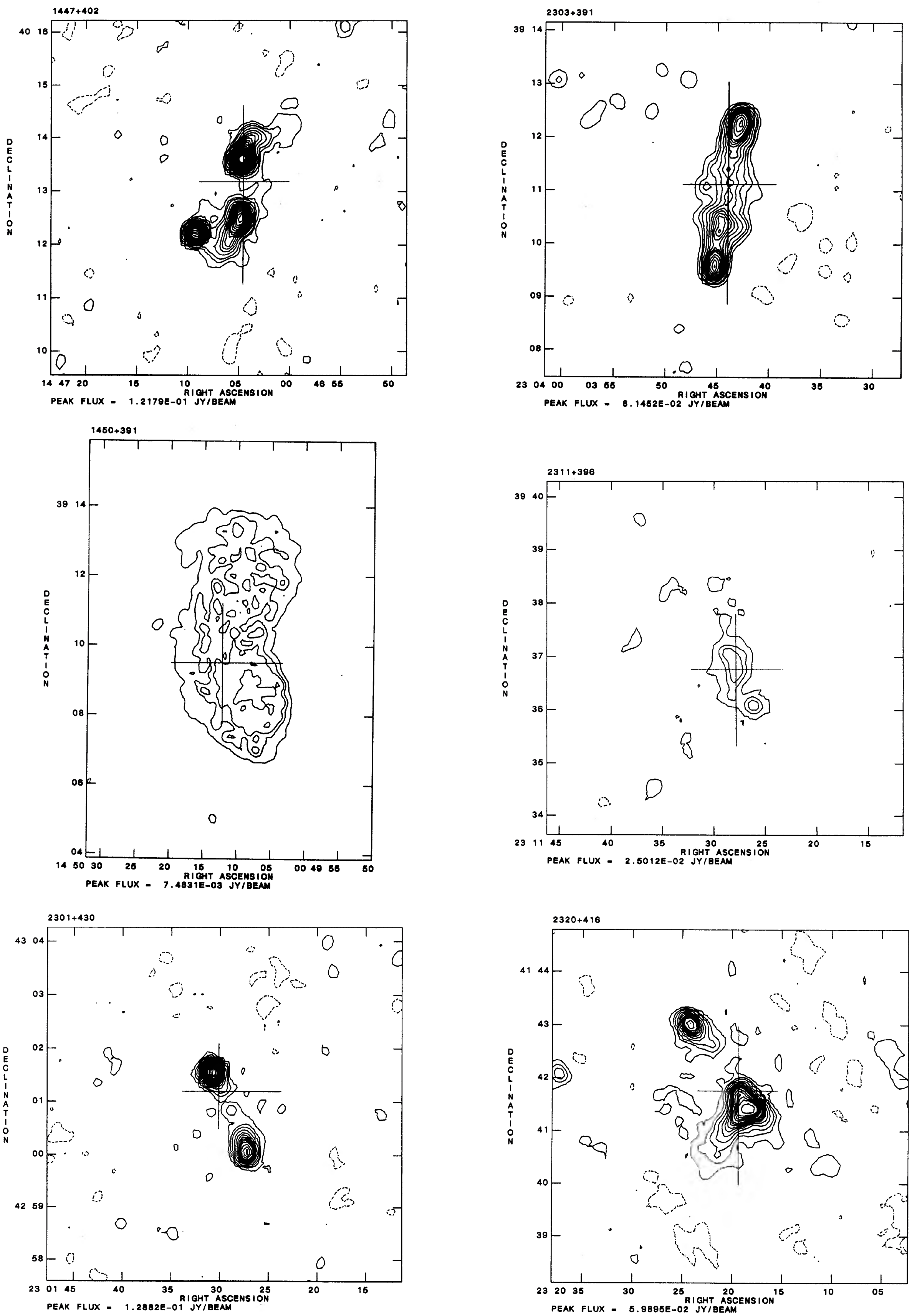

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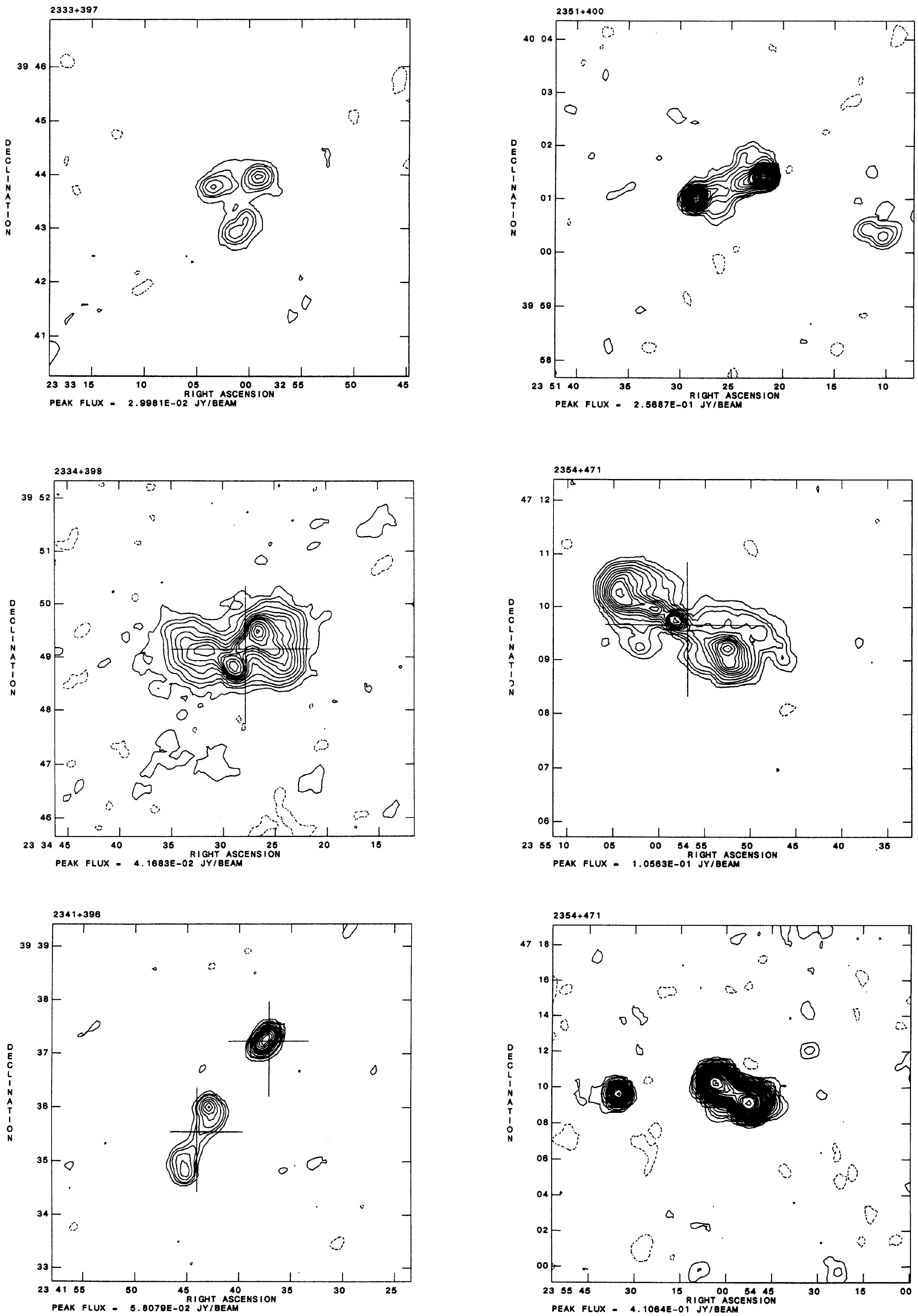

FIG. 3. (continued)

TABLE III. Radio-optical positional comparison for safe identifications.

Spectral index	Number	R.A. mean	R.A. dispersion	Dec. mean	Dec. dispersion
$\geqslant -0.5$	24	-0.03 ± 0.14	0.68	-0.27 ± 0.17	0.82
< -0.5	47	0.29 ± 0.27	1.83	-0.30 ± 0.23	1.55
< -0.5 (less 4 obj.)	43	0.33 ± 0.21	1.32	-0.22 ± 0.22	1.38

II, with an asterisk preceding the optical classification; a good example of this is 0754 + 394.

Sources with LAS$\geqslant$120 arcsec (47 objects in all) have been individually assessed. Most of them are safely identified with bright galaxies; the identifications in this range included in Table II, but considered uncertain, are marked with an asterisk preceding the optical classification, and are discussed in a note.

Table II(b) gives references for sources previously identified. These references are not complete; for sources for which a redshift has been published, only the reference to the redshift measurement is given. For 3CR sources, reference is made also to Spinrad *et al.* (1985), where further references to original papers can be found. Note that some 3CR sources classified as empty fields in Table II are in fact identified with objects beyond the limit of the POSS.

Table V shows the number of identifications proposed in Table II, according to LAS and radio-structure classification. Also shown in Table V are the expected numbers of random coincidences, computed on the basis of area searched, and a plate density of 2×10^{-4} objects/sq. arcsec. The random coincidence numbers shown refer to the total area searched (including thus objects in the notes) and are actually to be considered upper limits to the contamination of Table II.

VI. DISCUSSION

It has been well known since the early days of radio-source optical identification that the fraction of radio sources with optical identifications on the POSS plates sharply decreases with decreasing radio flux density. Figure 4 shows this trend in our data, separately for galaxies (dark symbols) and for blue starlike objects (light symbols). We also plotted in Fig. 4 some data points derived from the 5C 12 identification program (Benn *et al.* 1988) and data concerning the B2 and All-Sky catalogs at 408 MHz (Grueff and Vigotti 1977, and references therein). The present data are shown divided in three groups, including the N optical category, i.e., starlike objects of neutral color; these were classified as quasars in the 5C 12 and All-Sky catalogs, and were forced in either the quasar or the galaxy classification in the B2 sample. The horizontal dotted line is a conservative upper limit to the amount of random contamination in the present data.

The fraction of identification fainter than $m_r \approx 20$ (empty fields, invisible on the PSS) steadily increases from a mere 15% for $S_{408} \geqslant 10$ Jy, to 86% for sources in the 20–60 mJy interval. This trend obviously indicates some sort of positive correlation between the radio and optical flux densities; apart from that, a more detailed interpretation must await a complete identification program, with full redshift information, and the data themselves show the magnitude of the observational effort needed. We note that the identification program for the 5C 12 has been pushed to a limit of $m_v \approx 23$, and at that optical limit fully two-thirds of the sources remain unidentified.

Table VI lists the *average* source spectral index, and the *median* source LAS, computed in each B3 subsample. For the spectral index, standard errors of the mean are quoted; for the largest angular size, the 15% and 85% percentiles of the median are given.

In most cases, the median LAS is smaller than the HPBW of the present observations, so that it is evaluated mostly in connection with slightly resolved sources (R in Table II). LAS for R sources were computed by deconvolving quadratically the sources with the map beam (see Sec. V), which amounts to assuming a Gaussian shape for the beam and the source, an assumption certainly valid for the beam, but unlikely for the source. Figure 5 shows the theoretical correction to apply to LAS, evaluated in this way, if the source is an equal point double instead of a Gaussian. Also shown in Fig. 5 is the correction we actually adopted; it has been estimated statistically, by measuring a number of sources called R in Table II, but for which we already had the information coming from the A array maps available, at a resolution of 1.5 arcsec. We estimate that biases introduced in Table VI by the uncertainty in this correction are negligible compared to statistical errors. Uncorrected median LAS values are also given in Table VI for reference; in Table II, only uncorrected LAS values are given.

TABLE IV. Comparison of 3CR 1.5 GHz flux densities in Table II with KPW (1969) measurements.

3CR	*R*	σ	Number
strong	1.17	± 0.14	11
faint	1.17	± 0.20	11
compact	1.13	± 0.15	11
extended	1.21	± 0.19	11
all	1.17	± 0.17	22

TABLE V. Statistics of the identification proposed in Table II, according to source classification.

	U + R	Double	Triple	Total	
LAS < 60″	681	224	11	916	sources
	188	87	7	282	identif.
	15	14	1	30	random
60″$\leqslant$LAS < 120	8	54	17	79	sources
	6	24	10	40	identif.
	1	7	2	10	random

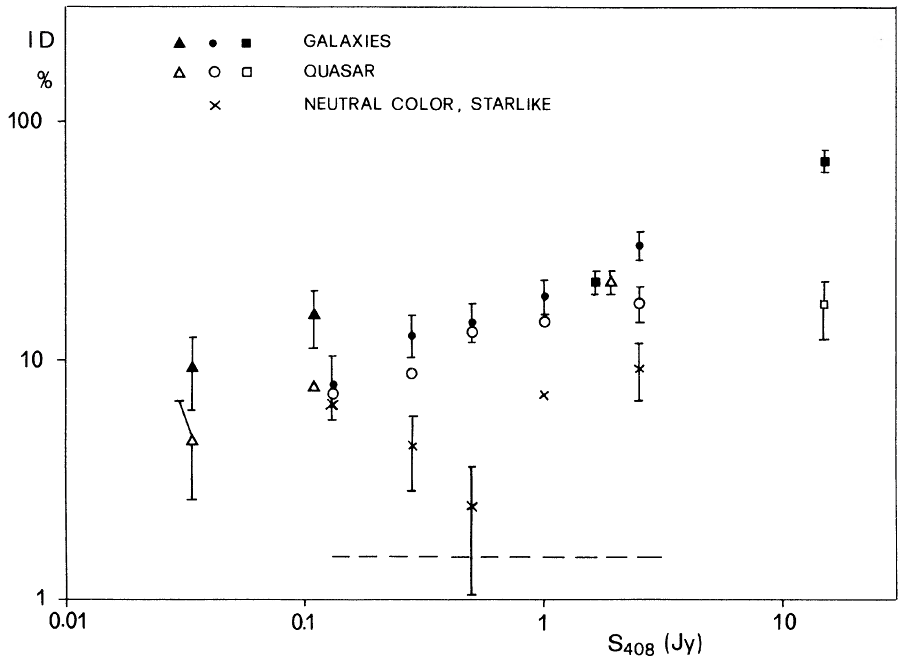

FIG. 4. Fraction of radio sources identifiable on the Palomar Sky Survey versus flux density. The horizontal dashed line indicates the level of expected random contamination in the present data (light and dark circles, crosses). Triangles indicate data from Benn *et al.* (1988), squares are from Grueff and Vigotti (1977).

Figure 6 shows the run of the average spectral index versus flux density at 408 MHz, computed in the B3 sample (light circles) and in the 5C 12 sample (dark circles). Neglecting for the moment the data point at 20 mJy, there appears to be a small but significant trend toward steeper spectra at low flux density. The straight line drawn is the least-square fit (data points weighted according to error), and its slope differs from zero at the 2.6σ level (probability less than 0.5%). The fit with a straight line is very good ($\chi^2 = 1.7$); on the other hand, the inclusion of the last point at 20 mJy will give an overall slope not significantly different from zero, but a rather bad fit (probability $\approx 5\%$).

Thus, there seems to be some evidence of the spectra becoming steeper with flux density decreasing to about 50 mJy, and becoming sharply flatter for flux densities fainter than that. There is good agreement between spectral indexes in the present sample and in the 5C 12 sample, in the overlapping flux-density range. Possible experimental biases producing the flat spectra in the faint 5C 12 interval have been searched, without finding any.

The suspicion that some real change is happening in the source population below ≈ 50 mJy is reinforced by inspecting Fig. 7, which shows the run of the median angular size (LAS) against flux density. The data from the present work and the 5C 12 are plotted with data taken from Oort (1988, and references therein) (error bars for 5C 12 data also mark the 15% and 85% percentiles; note that error bars from Oort (1988) are at the 35%–65% percentiles).

We can see that:

(i) Present data are entirely consistent with data at higher flux densities; (ii) 5C 12 data confirm the trend established by Oort (1988) and co-workers, of very compact angular sizes for faint sources; (iii) Present data define the LAS–

TABLE VI. Average spectral index and median LAS in the five B3 samples.

Sample	Spectral index	Median LAS corrected (70% confid. int.)	Median LAS uncorr.	Median flux (Jy)
0	-0.930 ± 0.025	8–11	10.5	0.13
1	-0.920 ± 0.016	9–12	12	0.28
2	-0.893 ± 0.016	7–10	8	0.5
3	-0.883 ± 0.021	7–14	11	1.0
4	-0.890 ± 0.019	11–18.5	16	2.5

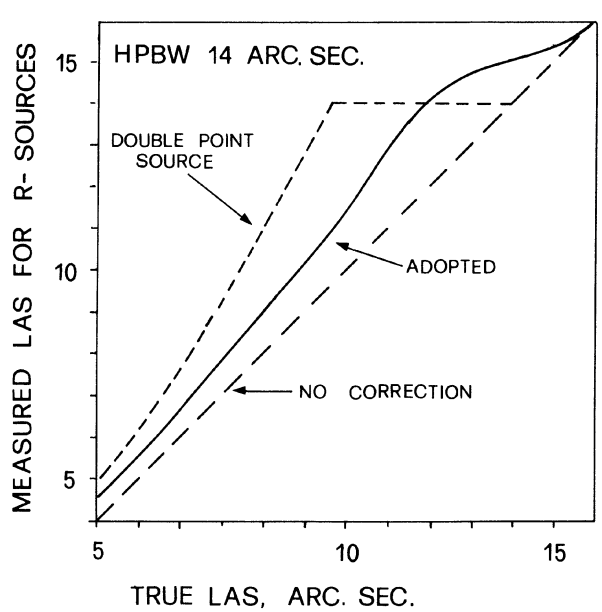

FIG. 5. Correction to apply to the LAS measured for barely resolved sources, if they are in fact equal doubles (small dash), and the empirical correction adopted (full line).

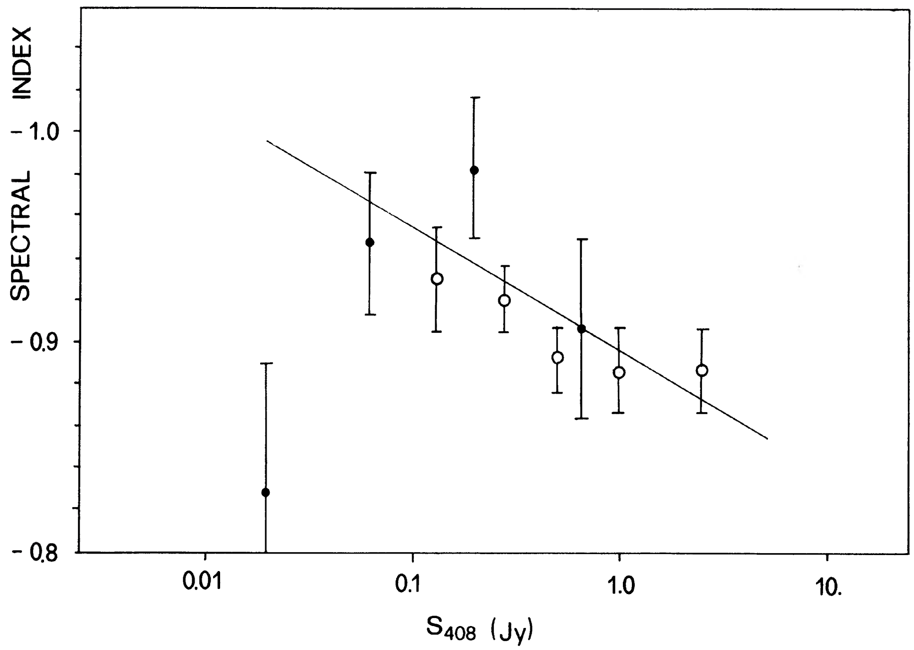

FIG. 6. Average spectral index as a function of flux density. Light circles represent the present data; dark circles are from Benn *et al.* (1988). The full line is the best fit to the data, excluding the point at 20 mJy.

flux-density relation in the interval 100–600 mJy, previously poorly defined, and show that a plateau is probably reached for LAS in this range.

Note that there is good agreement between the B3 and 5C 12 points at $S = 0.10$ Jy, the latter being obtained with data at much higher resolution (5 arcsec, VLA B array at 1.46 GHz).

For identifications with galaxies, Table II gives a rough estimate of redshift; although inaccurate (see Fig. 2), these values can be used to derive metric sizes and absolute powers. A Friedmann model with $q_0 = 0.5$ and $H_0 = 75$ km/s/Mpc has been adopted, and metric sizes were derived from uncorrected LAS values in Table II.

Figure 8 shows the histogram of source metric sizes, separately for low ($z < 0.2$) and high ($z > 0.2$) redshifts, the value 0.2 being the approximate median of the sample redshift distribution. Because only reasonably strong sources can be observed for $z > 0.2$, a lower limit in absolute power of $\log P > 25$ (P in W/Hz) has been adopted, corresponding roughly to the lowest power still represented in the "distant" sample. The 70% confidence levels for the median in the two observed distributions are 60–83 kpc for the distant galaxies and 87–126 kpc for the nearby ones, the medians being 70 and 112 kpc, respectively; these values differ at the 1.5% significance level.

Possible systematic errors in our estimated redshifts are very unlikely to have caused this result; in particular, there could be some hint in Fig. 2 for a systematic underestimation of the redshifts smaller than 0.05, which (if confirmed and corrected) will increase the difference in metric sizes. This

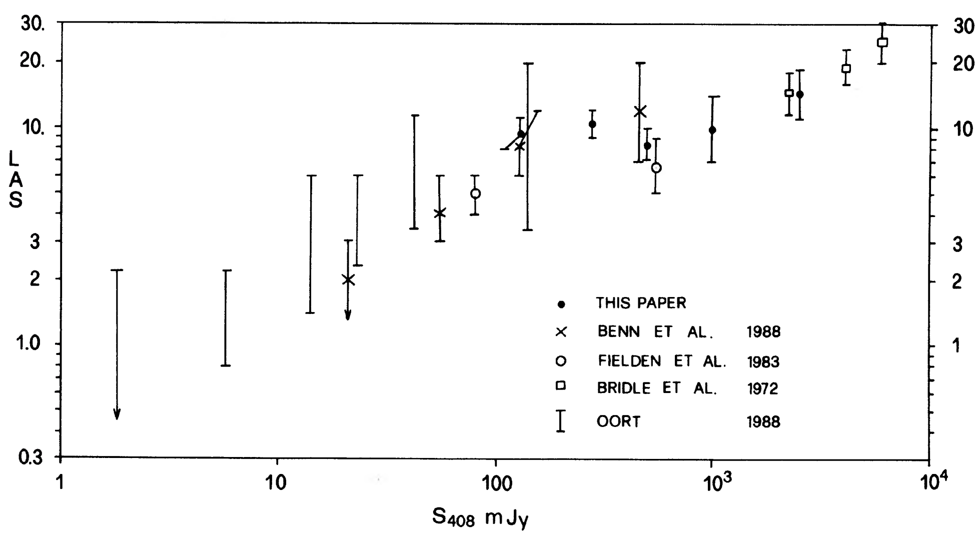

FIG. 7. Median angular size (LAS) as a function of flux density. Error bars in the data from Oort (1988) mark the 35%–65% percentiles, while other error bars mark the 15%–85% percentiles (i.e, they represent roughly ± 1 rms).

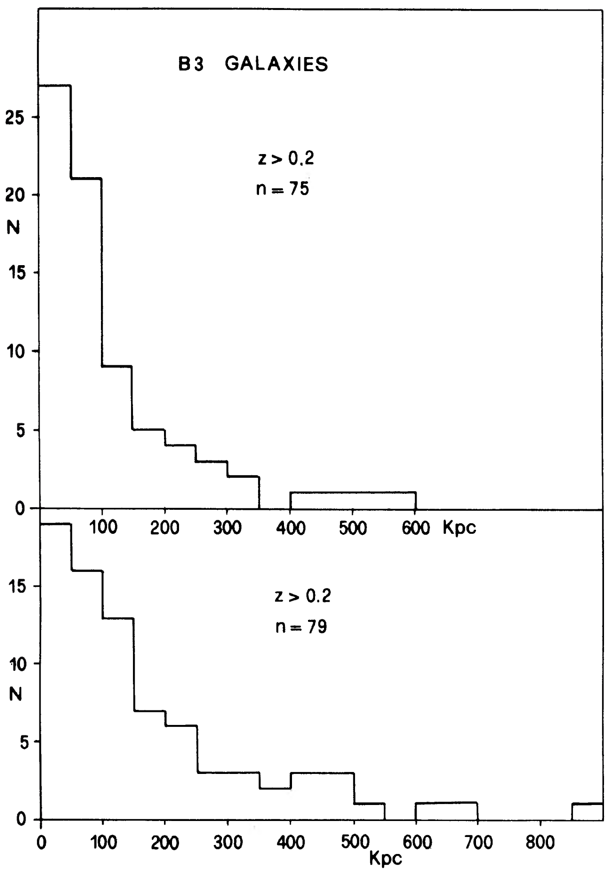

FIG. 8. Histograms of metric sizes for "near" radio galaxies ($z<0.2$) and "far" radio galaxies ($z>0.2$). Nearby radio galaxies are significantly more extended.

difference is also insensitive to a choice of a lower value for q_0, while, of course, a higher q_0 would make things worse. Only with the adoption of a "Euclidean" model, i.e., LAS $\propto 1/z$, the probability for the two medians being equal would be substantially increased.

On the other hand, it is well known that in any flux-density-selected sample there is a correlation between absolute power and distance (Malmquist bias), and Fig. 9 shows this effect in our data (including $\log P \leqslant 25$).

If our sample is split according to absolute power instead of redshift, we find a median metric size of 70 kpc for $\log P \geqslant 26.0$ and 106 kpc for $\log P < 26.0$. Consequently, a smaller metric size could be associated either with higher absolute power or higher redshift.

We can test which correlation is the most likely to be true, using the rank-correlation test (see, for example, Dixon and Massey 1969). In this test, a correlation coefficient R is defined as follows:

$$R = 1 - (6\Sigma_1 d_1^2)/[N(N^2-1)] ,$$

where d_1 is the ith rank difference between the two variables considered, and N is the sample size.

Note that a nonparametric test is needed, due to the strongly non-Gaussian shape of the diameter distribution (Fig. 8). Applying the test to the (absolute power) versus

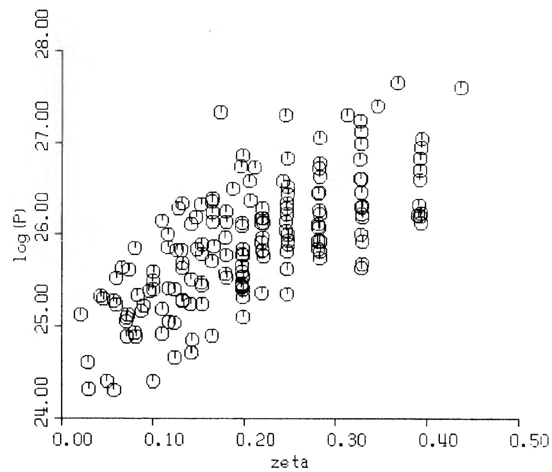

FIG. 9. Correlation between redshift and absolute power for radio galaxies in the present sample, due to Malmquist bias.

(metric diameter) correlation, we find $R = -0.061$, a value not significantly different from zero (the minus sign would imply smaller sizes for higher power). The (redshift)–(metric diameter) correlation gives $R = -0.205$, a highly significant value (bilateral random probability less than 1%). Both values of R are calculated for strong sources only ($\log P > 25.0$).

We thus conclude that the present data indicate a negative correlation (evolution) of metric sizes of strong radio galaxies with redshift, confirming a result obtained by Grueff *et al.* (1977), and already noted for quasars by Wardle and Miley (1974). A similar conclusion has been reached recently by Kapahi (1986).

Figure 10 shows the absolute power distribution of radio galaxies in Table II. The peak at $\log P = 26$ is, of course, expected, given the median sample flux density (≈ 0.5 Jy) and redshift (≈ 0.2); a detailed study of the distribution, and the determination of the luminosity function, are deferred to future papers.

For quasars in Table II, little can be learned without spec-

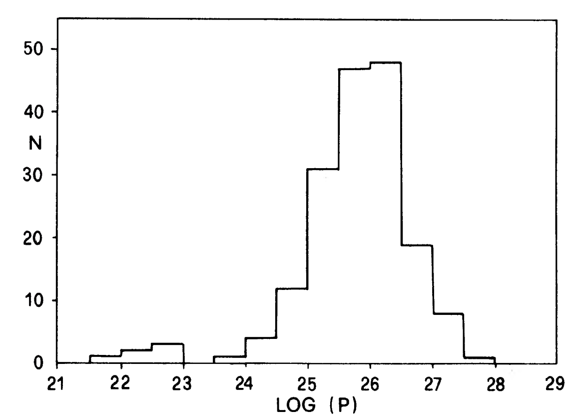

FIG. 10. Distribution of absolute power at 408 MHz for radio galaxies in the present sample. A Friedmann model with $H_0 = 75$ and $q_0 = 0.5$ is adopted.

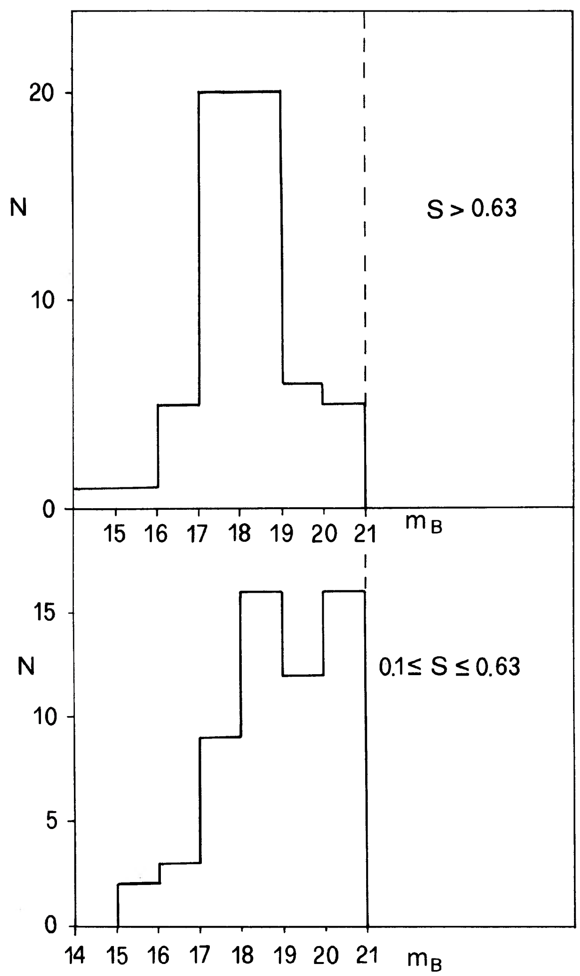

FIG. 11. Distribution of apparent blue magnitudes for B identifications (quasars), radio bright $S_{408} > 0.63$ Jy), and radio faint ($S_{408} \leqslant 0.63$ Jy). The dashed line is the Palomar Sky Survey cutoff.

troscopic redshifts, since the very large dispersion generally observed for quasars' optical luminosity prevents any meaningful redshift estimate from apparent magnitude.

In Fig. 11 we show the magnitude distribution of blue object identifications, divided in two equally populated groups according to 0.4 GHz flux density; one can see that the POSS optical cutoff (dotted line at $m_B \approx 21$) is almost ineffective for "high flux density" blue object identifications (an effect first noted in the 3CR by Bolton (1969)), while it is definitely at work in fainter source samples (see also the discussion in Grueff and Vigotti (1973)).

VII. CONCLUSION

We have mapped a large sample of radio sources selected at low frequency, free as far as possible from unwanted biases, and suitable for study of the cosmological evolution of radio sources. The sample covers the flux-density decade 1–0.1 Jy (at 408 MHz) and should provide a suitable database to extend a program of complete optical identifications beyond the 3CR Catalogue. As a first step in this direction, we give intermediate-resolution, accurate radio positions and structures, as well as optical identifications insofar as permitted by the Palomar Sky Survey red and blue prints. The ongoing program will eventually include high-resolution (1.5 arcsec) radio mapping of all the sources, morphological classification of radio structures, complete identification of empty fields, and spectroscopic redshift determination for all the sources. These, and other data, should lead to a better understanding of the evolutionary properties of the extragalactic radio-source population.

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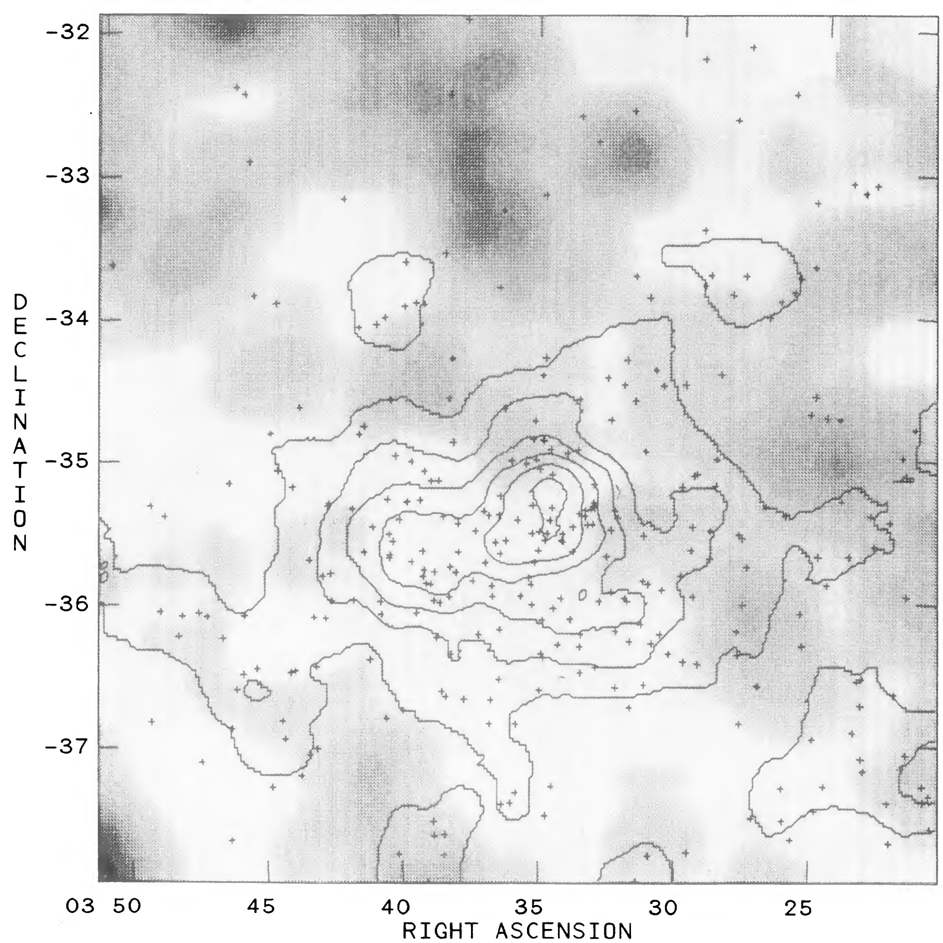

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